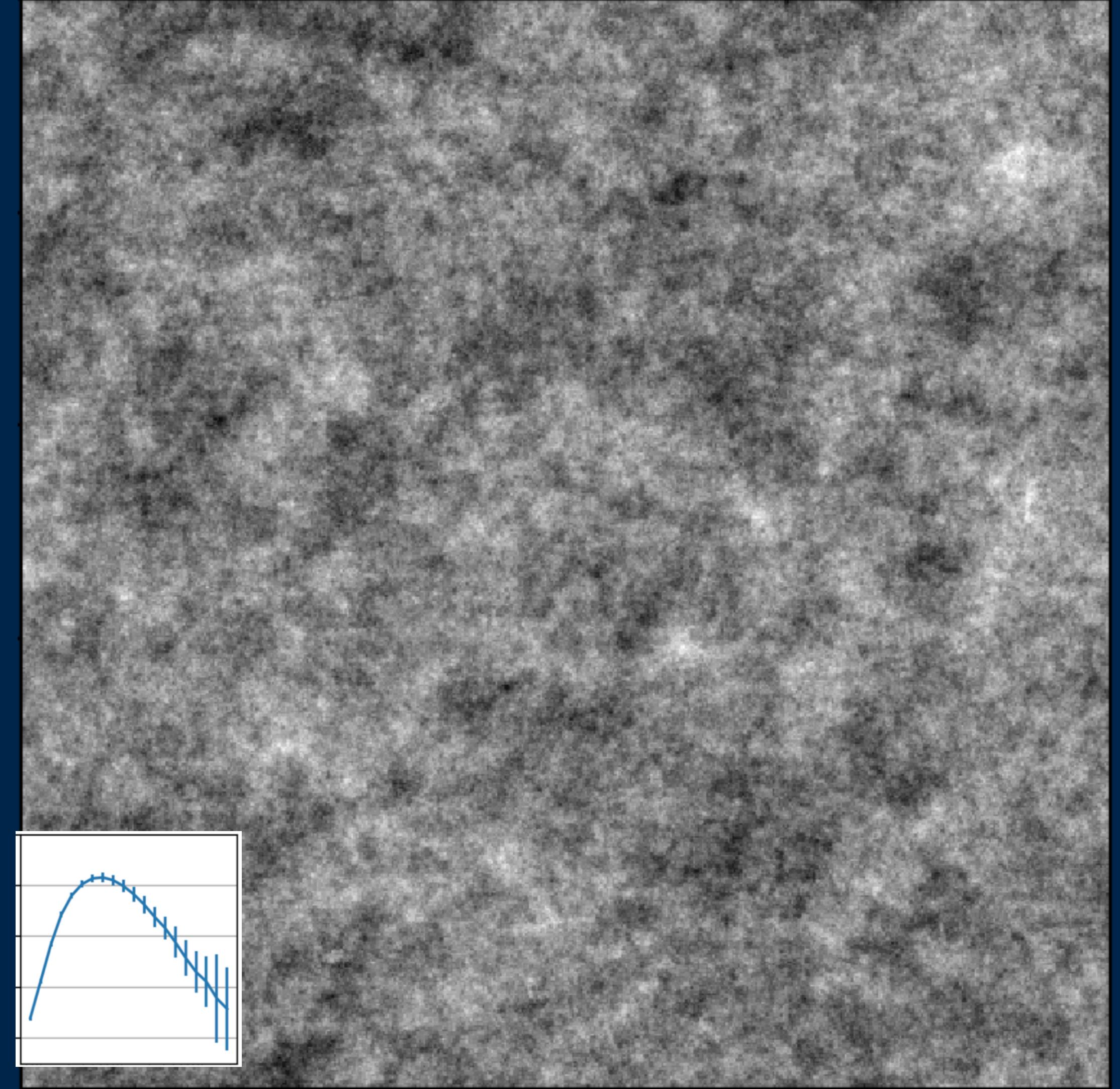
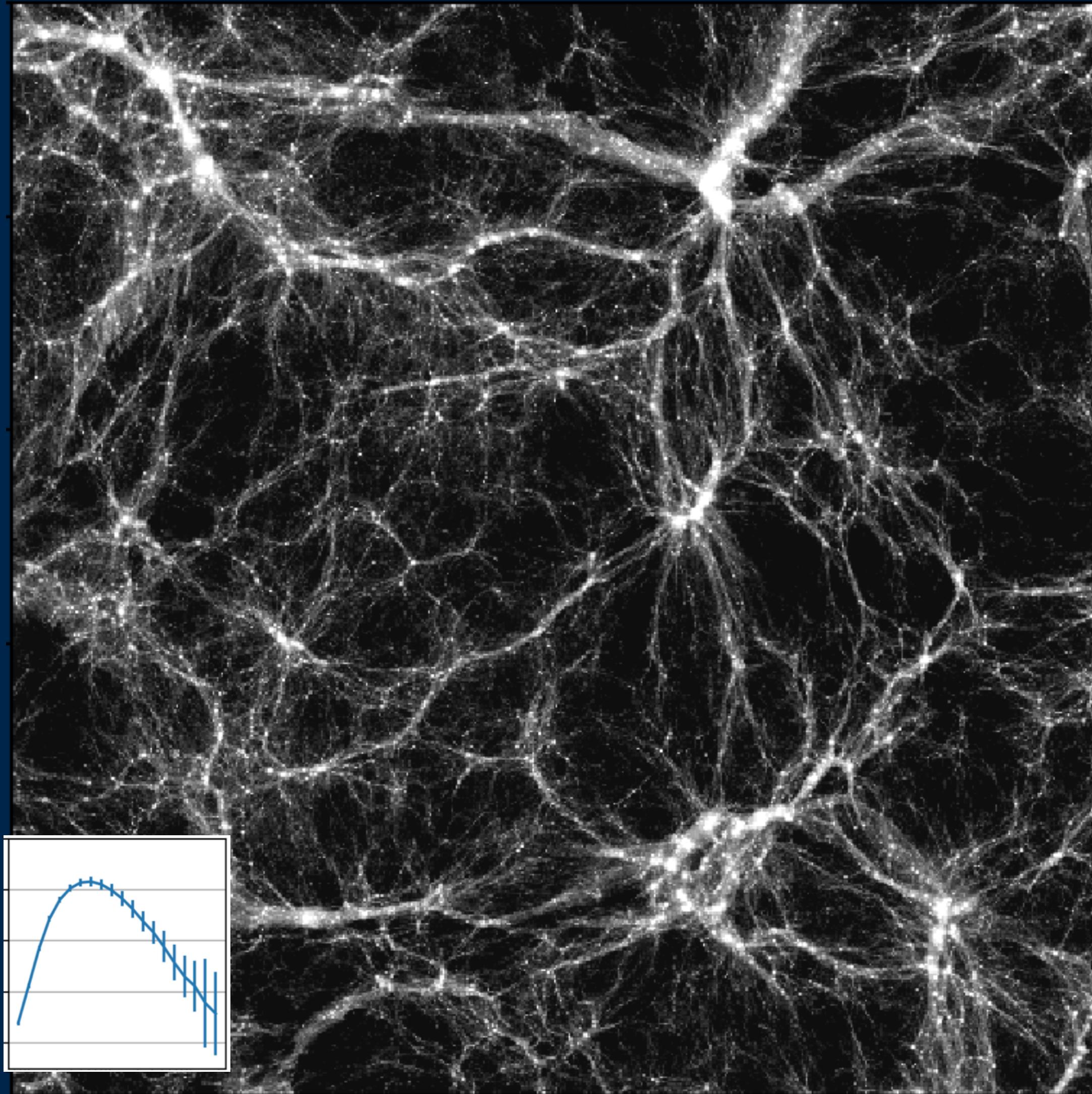


A new vocabulary for textures and its astronomical applications

Sihao Cheng (程思浩)
Johns Hopkins University

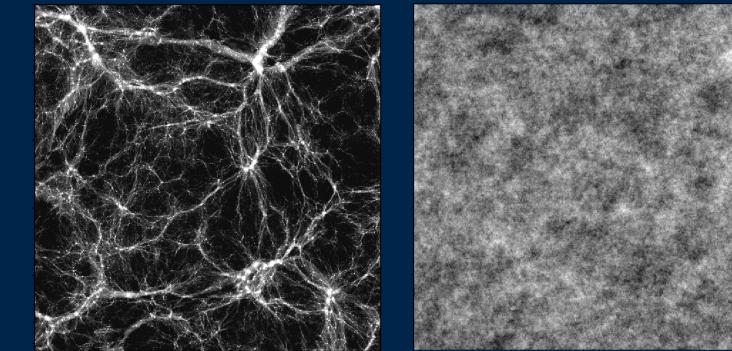
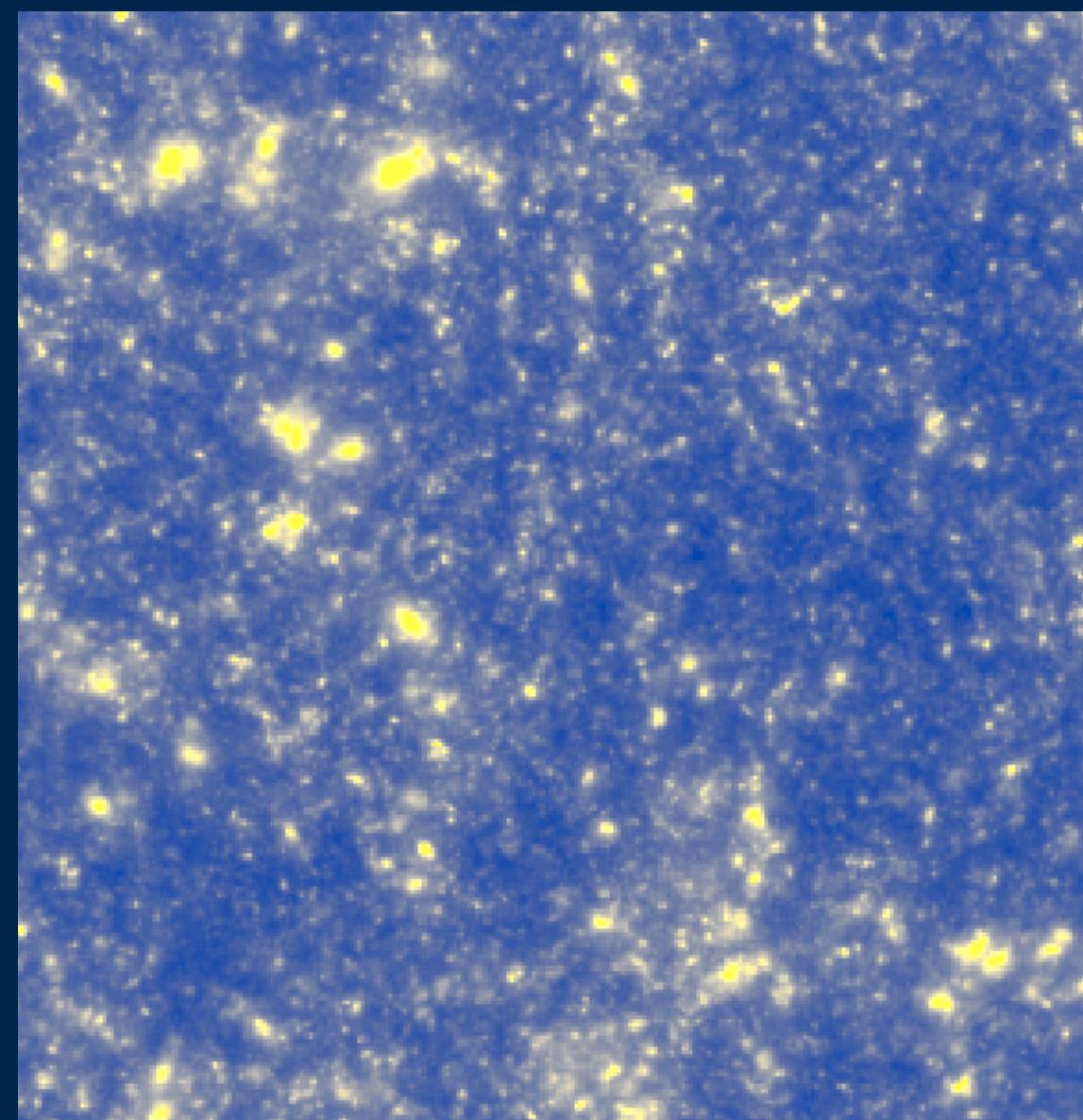
Where the Earth Meets the Sky
May 28

with Brice Ménard, Yuan-Sen Ting, & Joan Bruna



We need to go **beyond** power spectrum to capture non-Gaussianity.

How do we characterize a field?



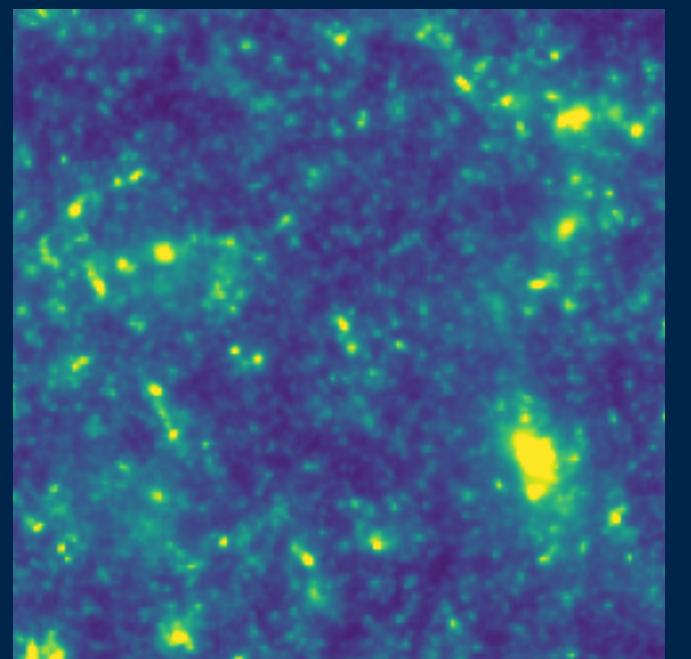
power spectrum
and others

scattering transform
(Mallat 2012)

CNN

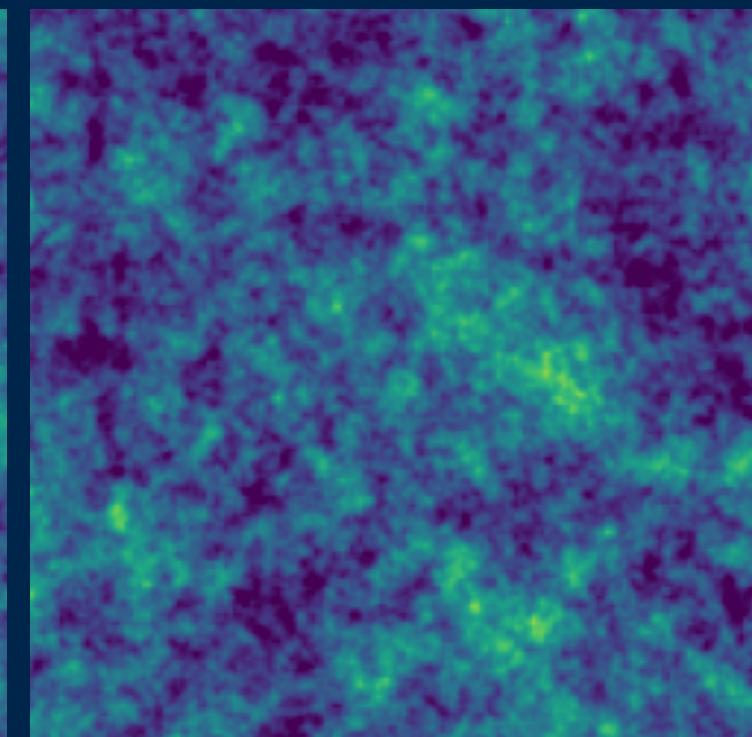
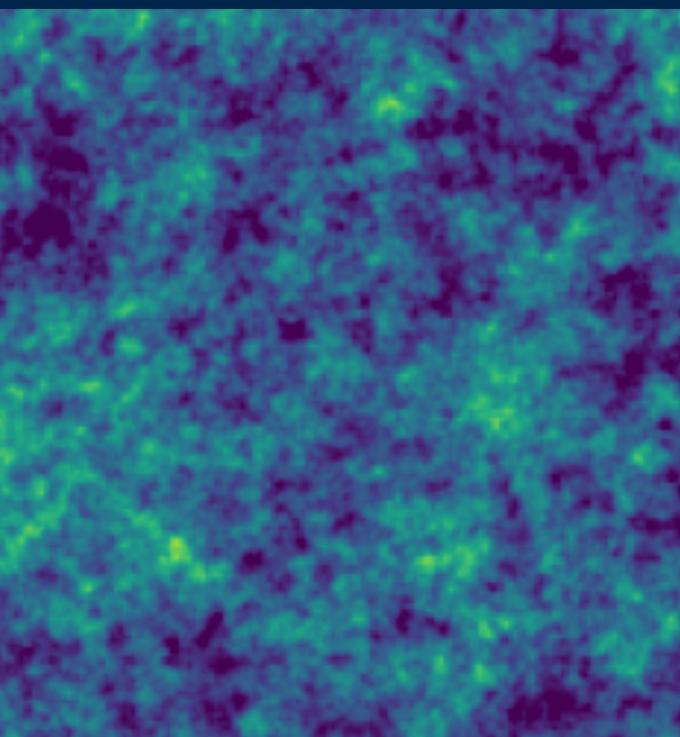
powerful,
but less interpretable or controllable

physical
parameters

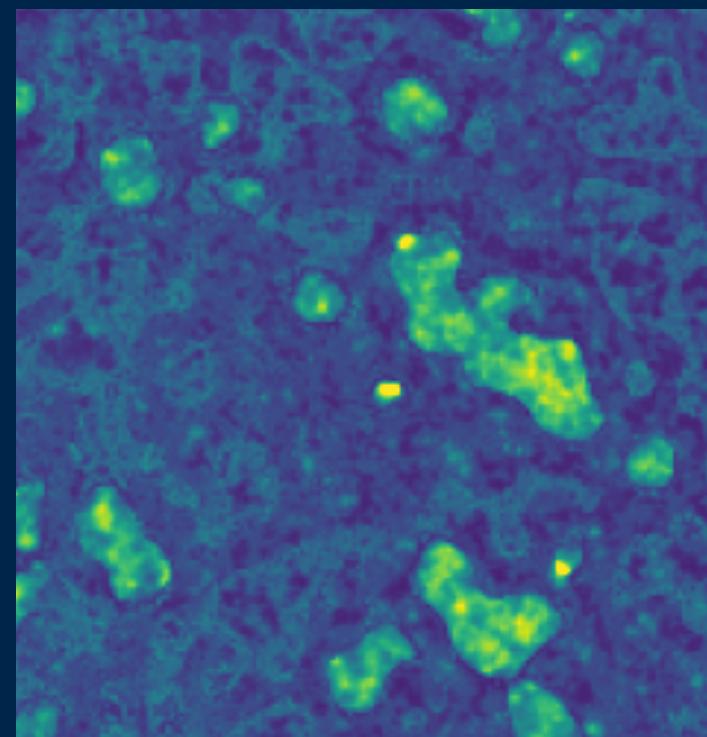
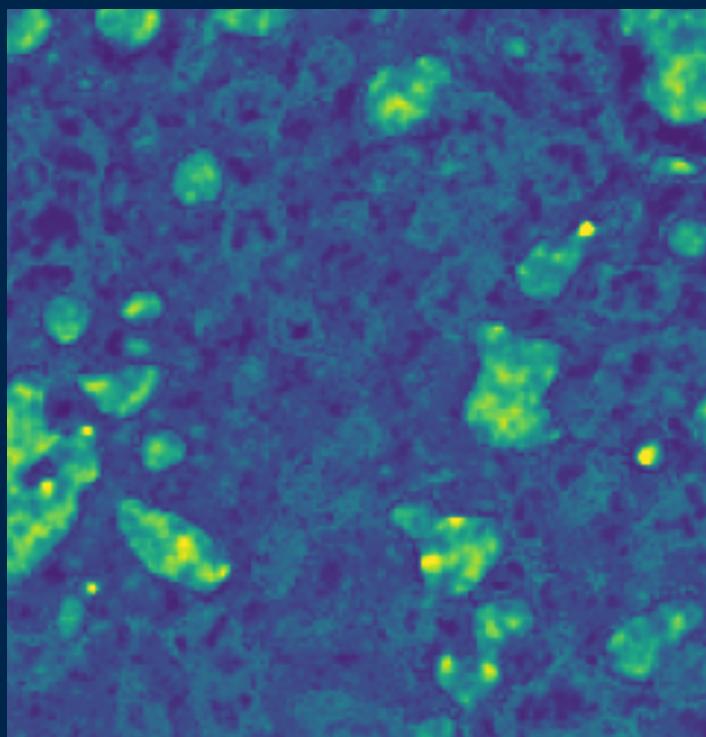


cosmic mass map

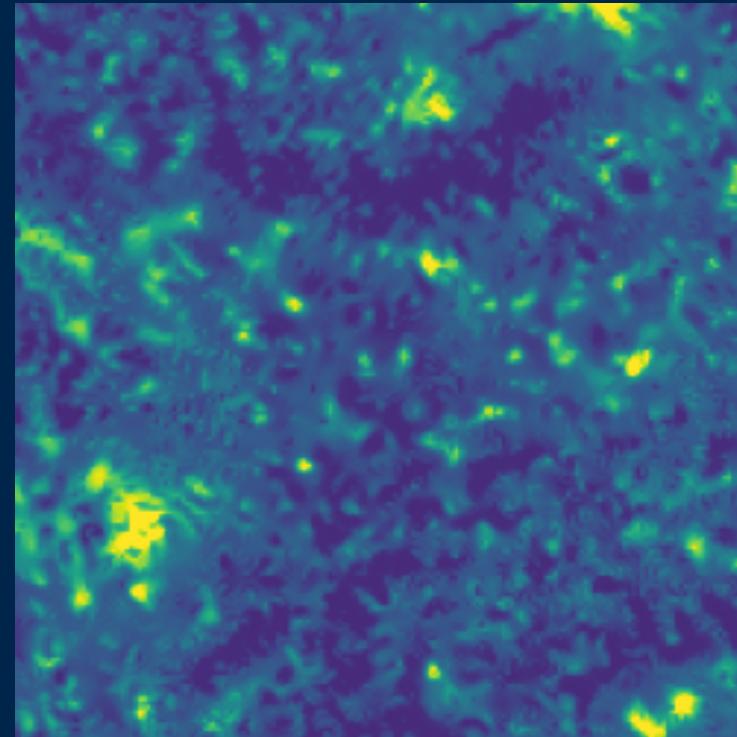
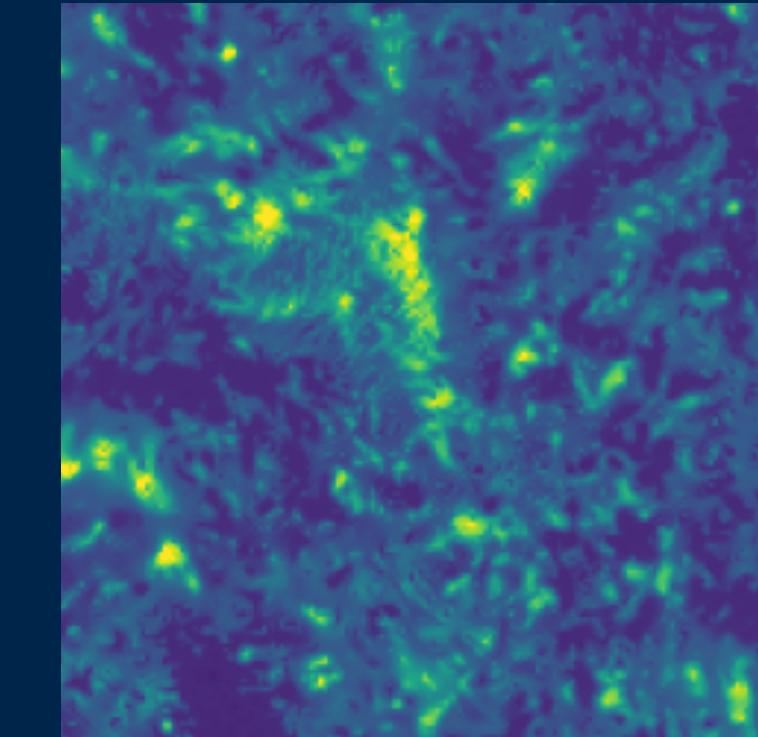
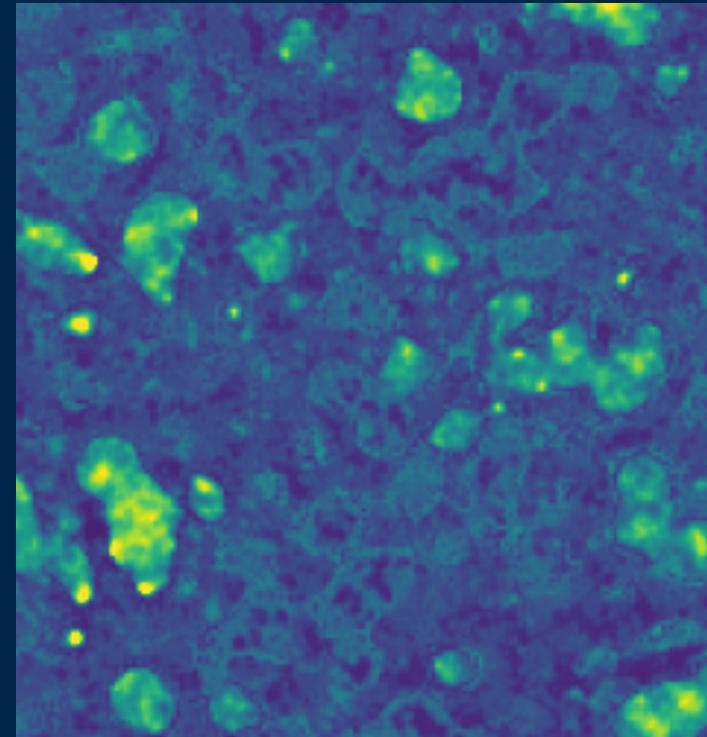
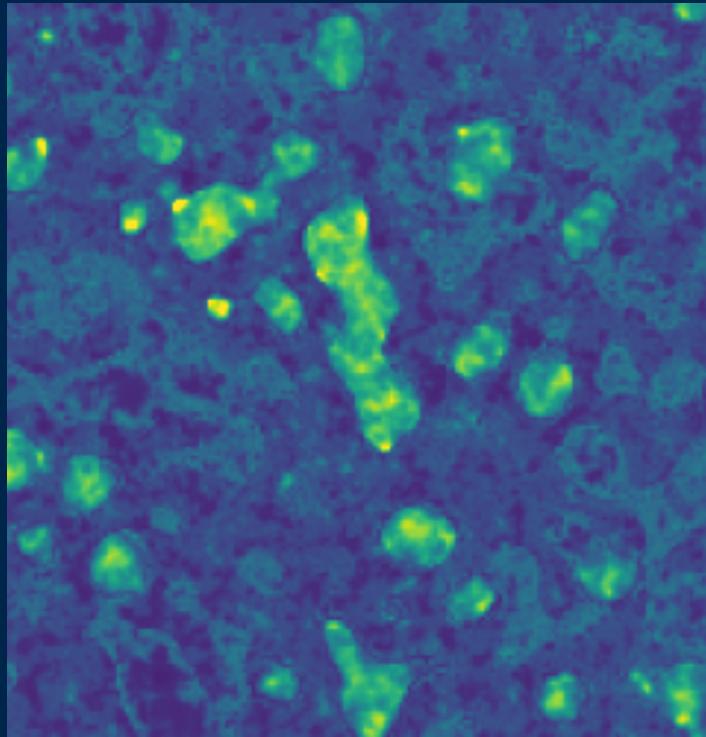
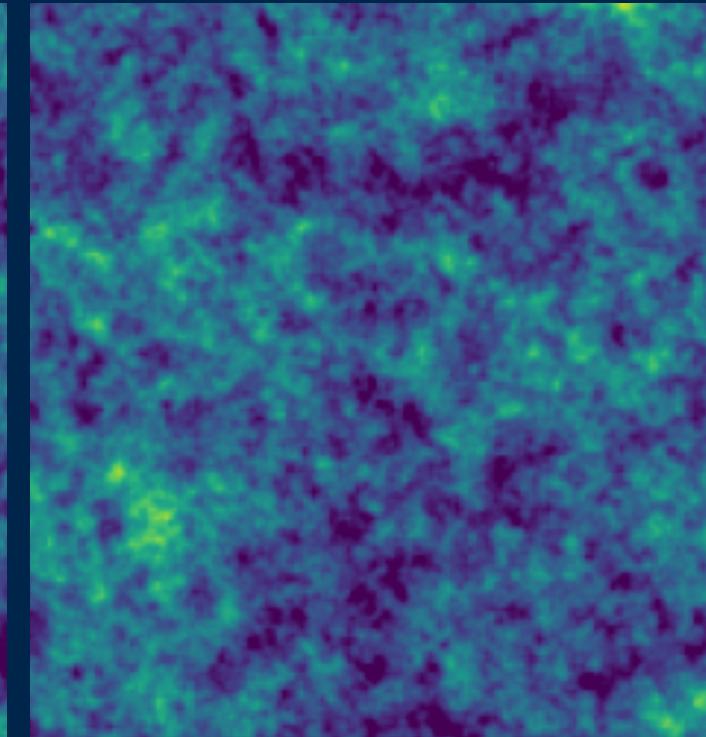
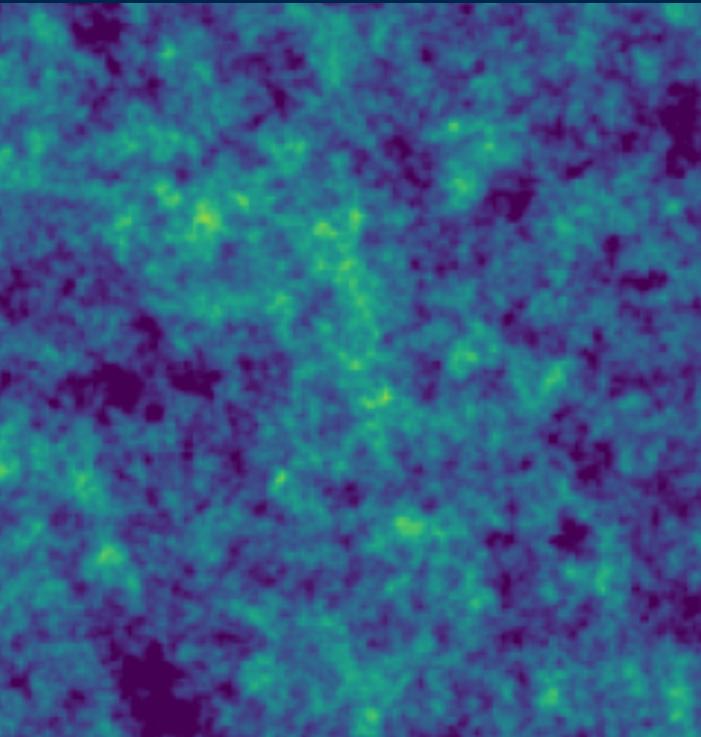
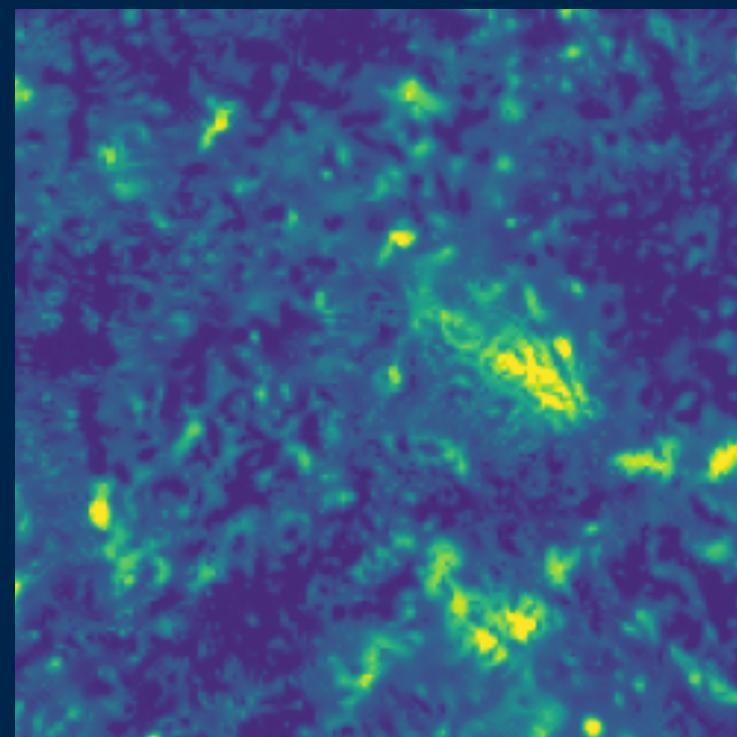
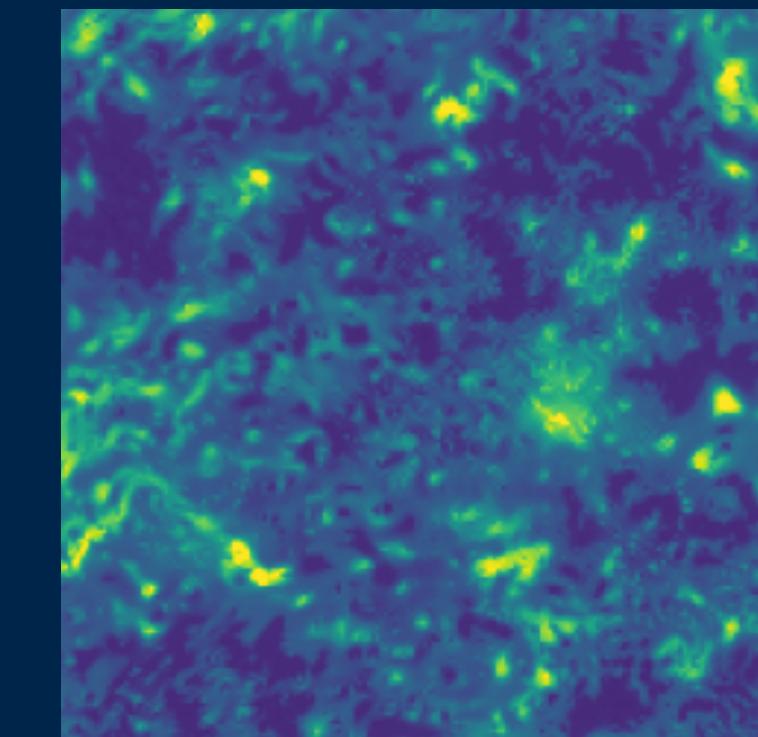
with power spectrum $P(l)$



with $P(l)$ and bispectrum



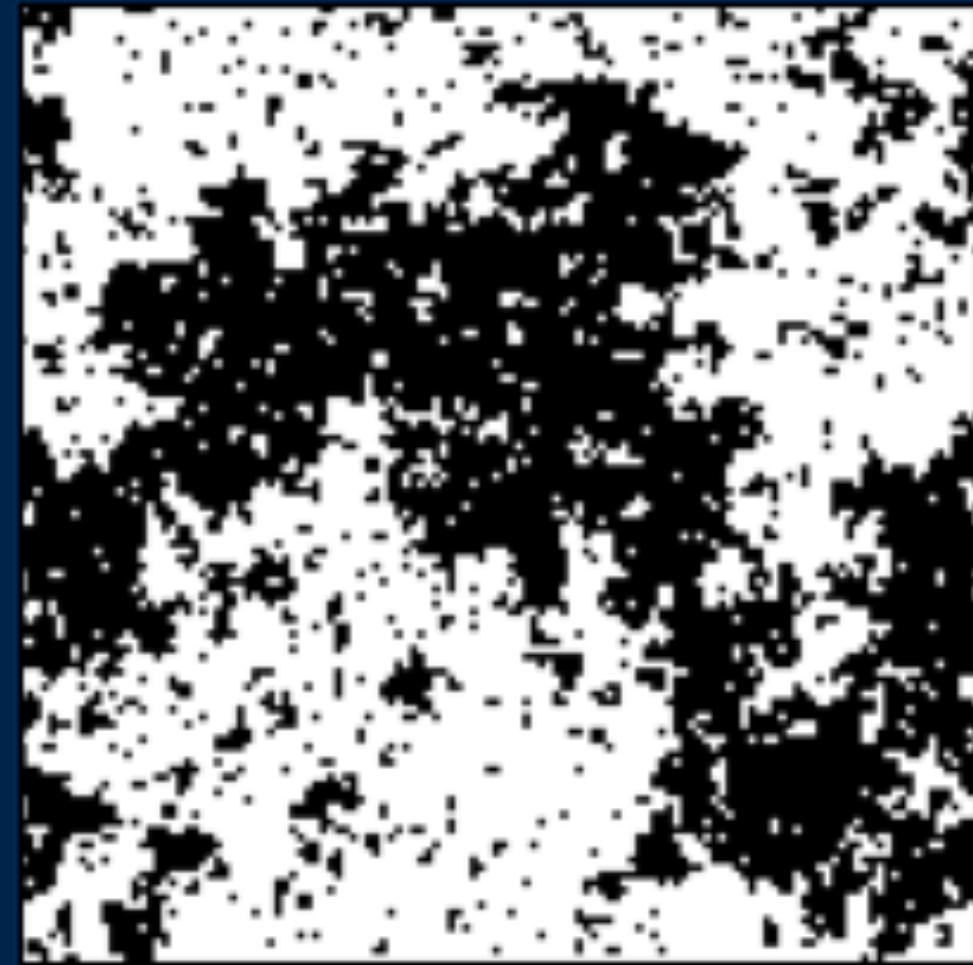
with scattering statistics



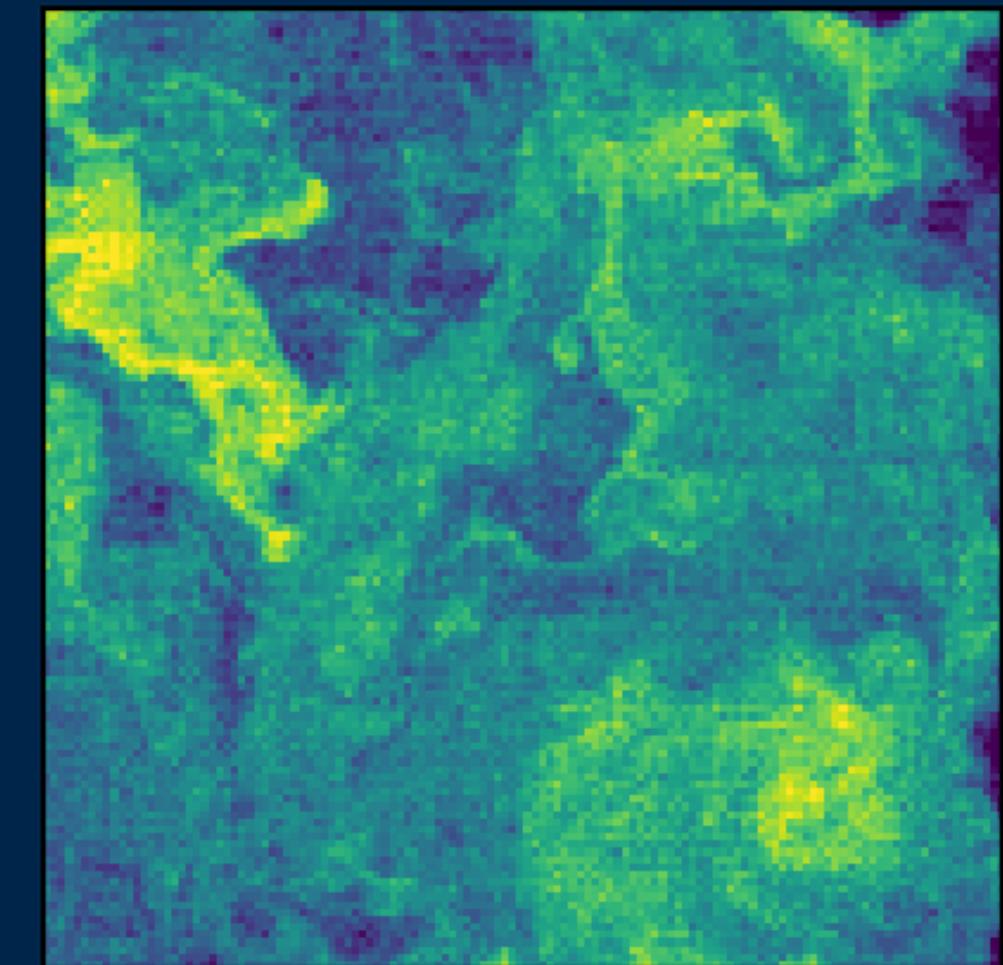
Turing pattern



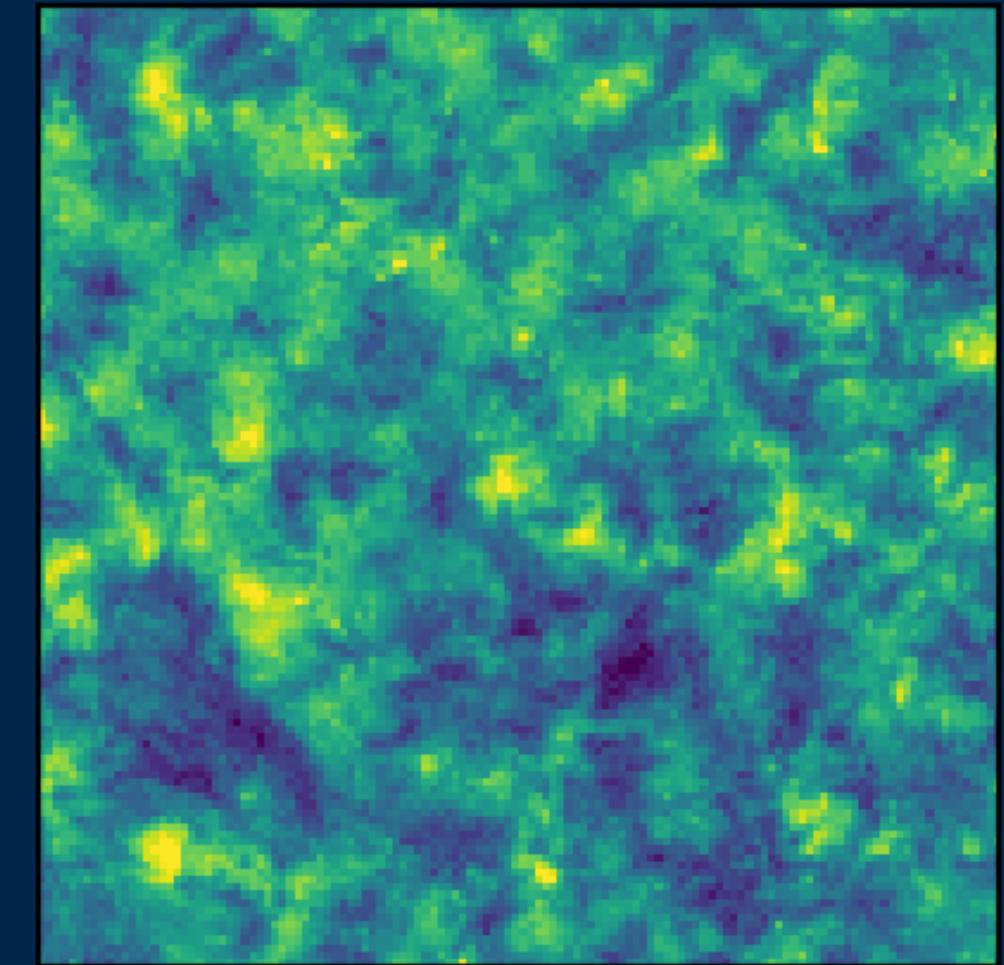
Ising model



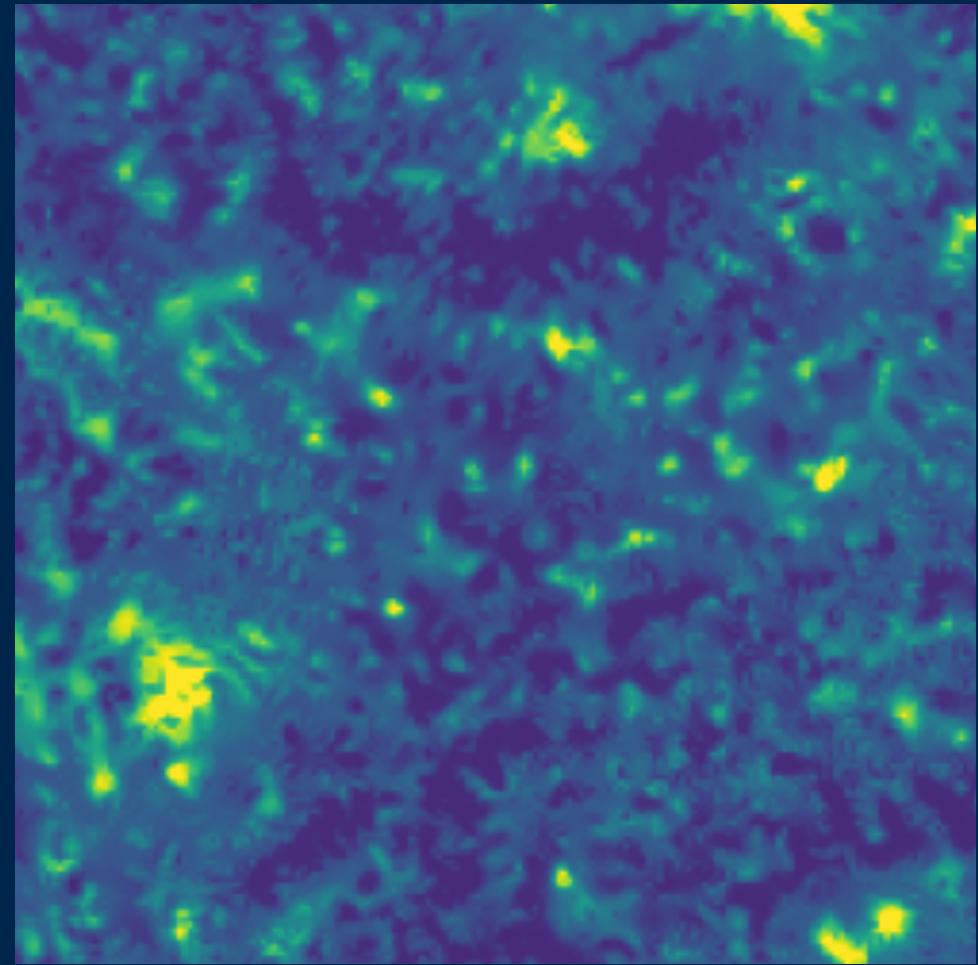
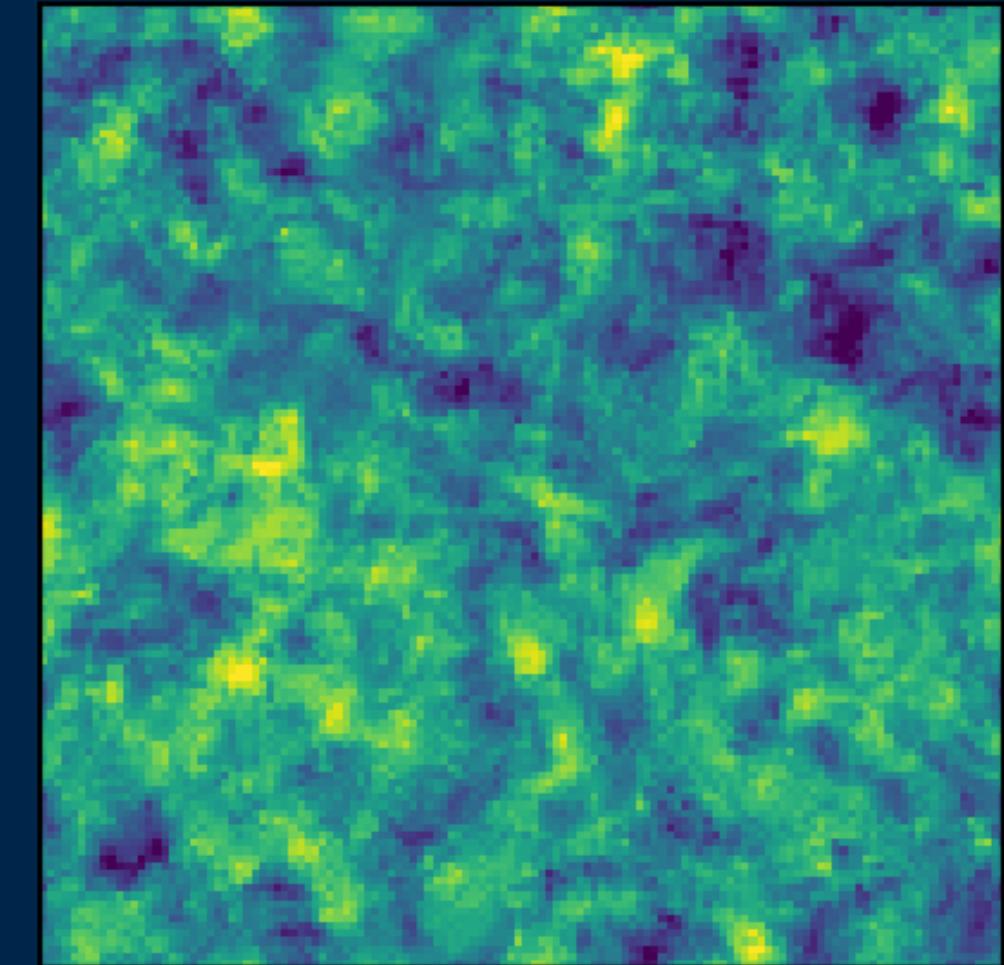
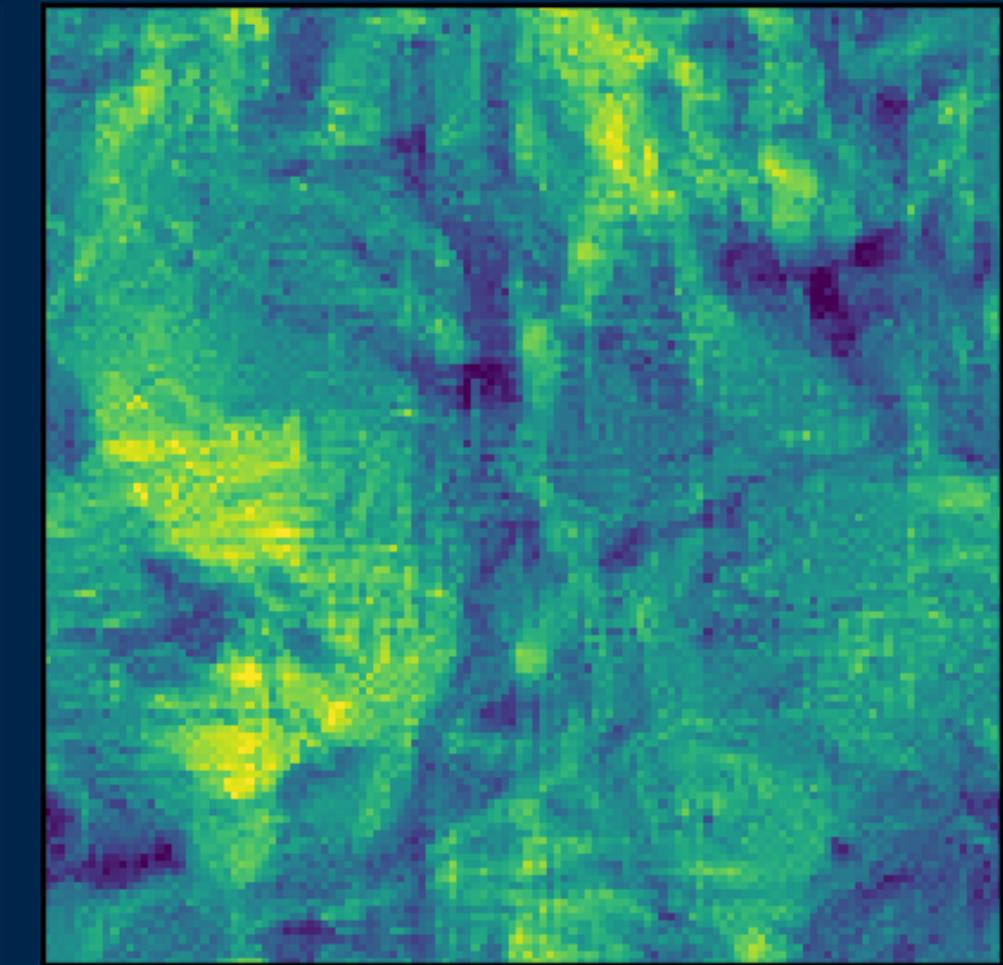
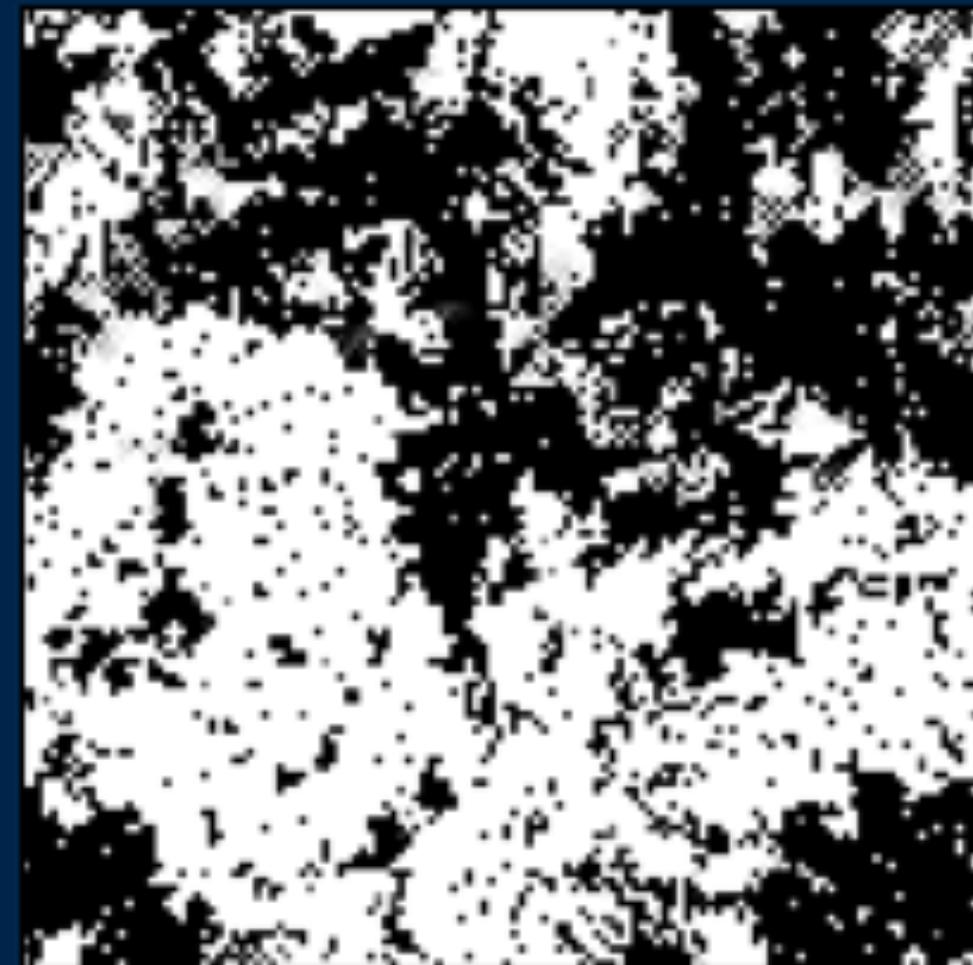
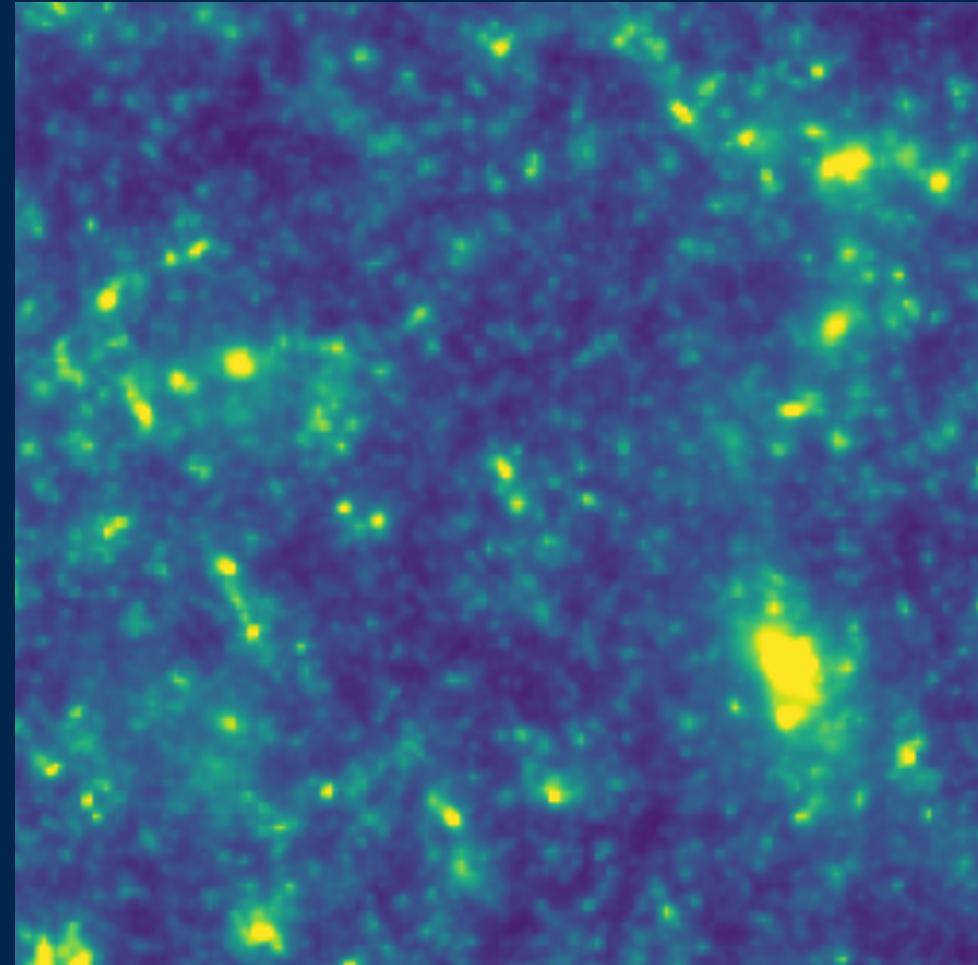
sea temperature



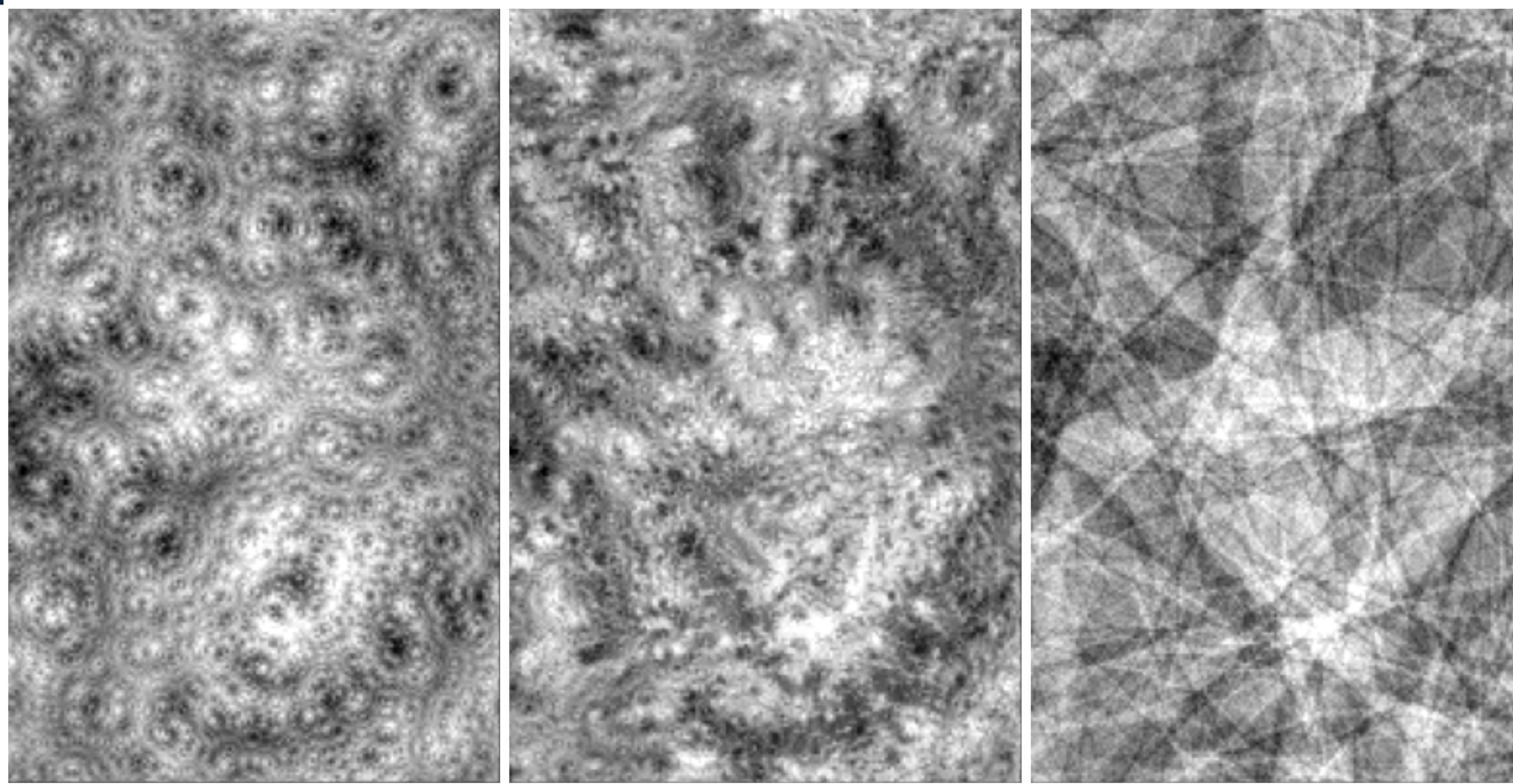
solar UV image



cosmic matter



Swirls vs. Origami $S_2^{\parallel} / S_2^{\perp}$



<1

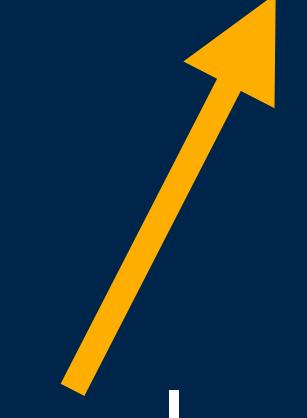
=1

>1

Fourier oscillations

$$P(\vec{k}) = \langle |I \star e^{k \cdot x}|^2 \rangle$$


local kernels (wavelets)

$$S(\vec{k}) = \langle |I \star \psi_k| \rangle$$


$$P(\vec{k}) = \langle |I \star e^{k \cdot x}|^2 \rangle$$

Fourier oscillations



$$S(\vec{k}) = \langle |I \star \psi_k| \rangle$$

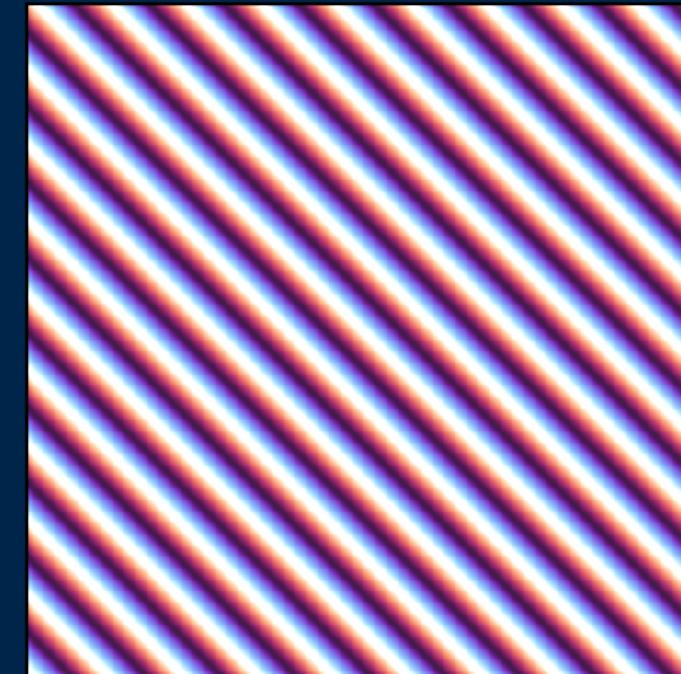
local kernels (wavelets)



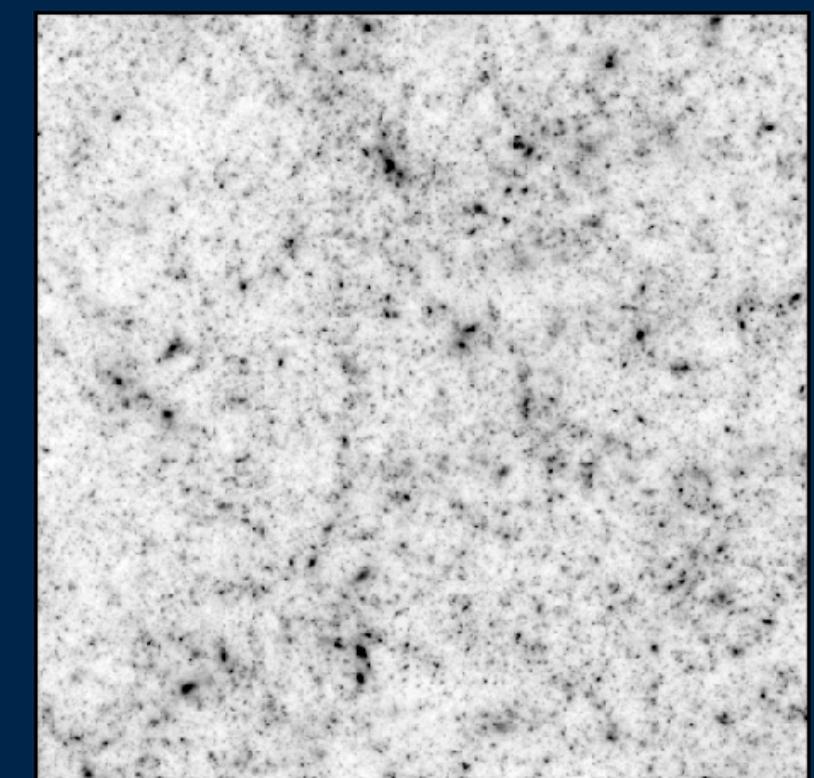
convolution



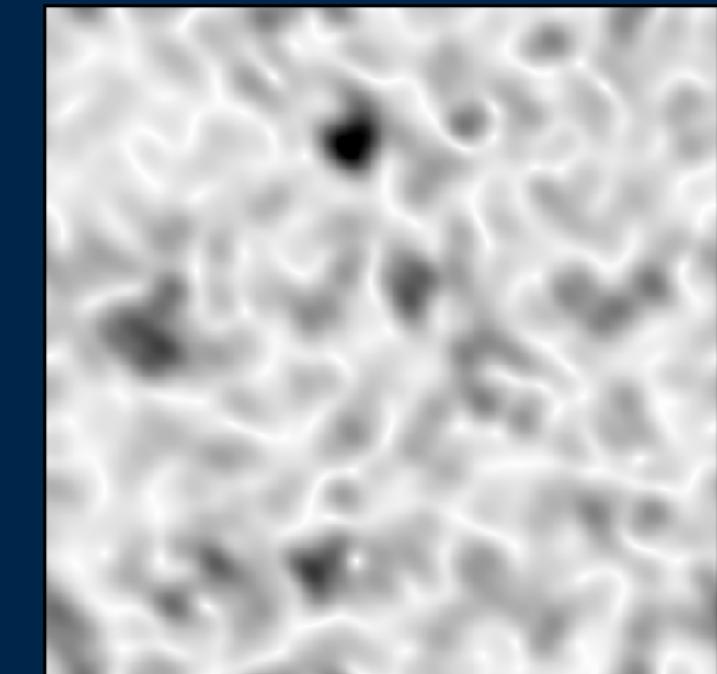
modulus



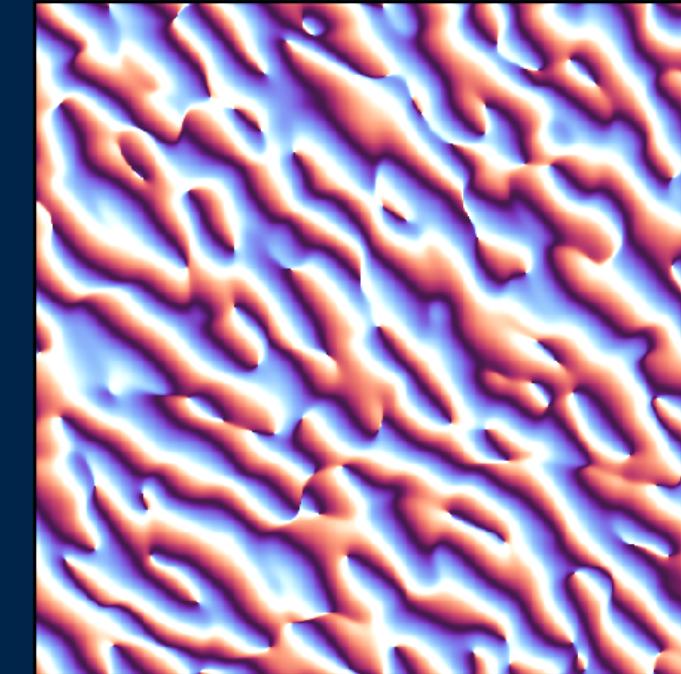
phase



convolution



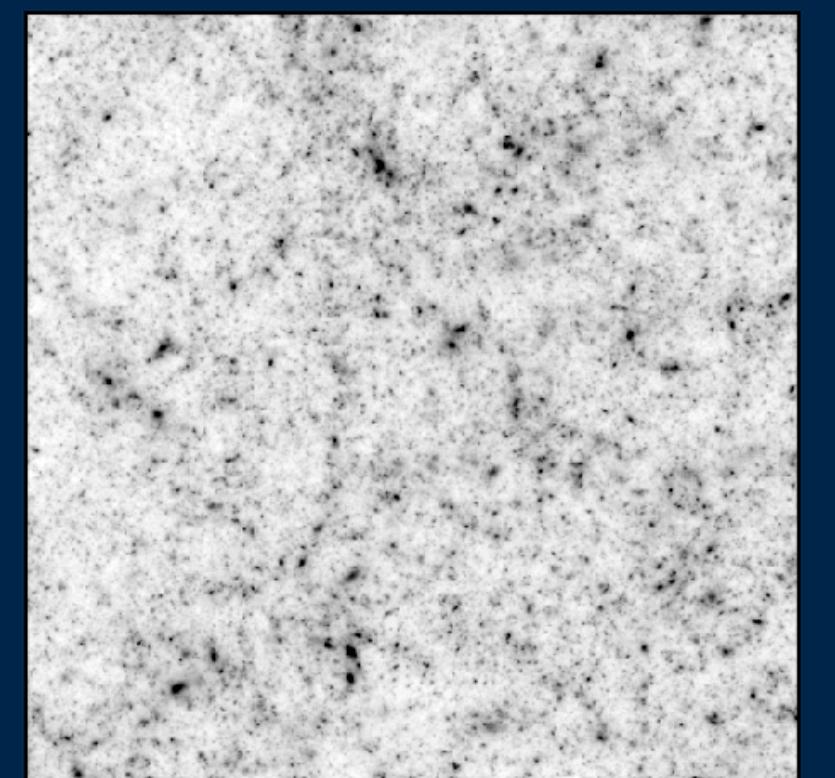
modulus



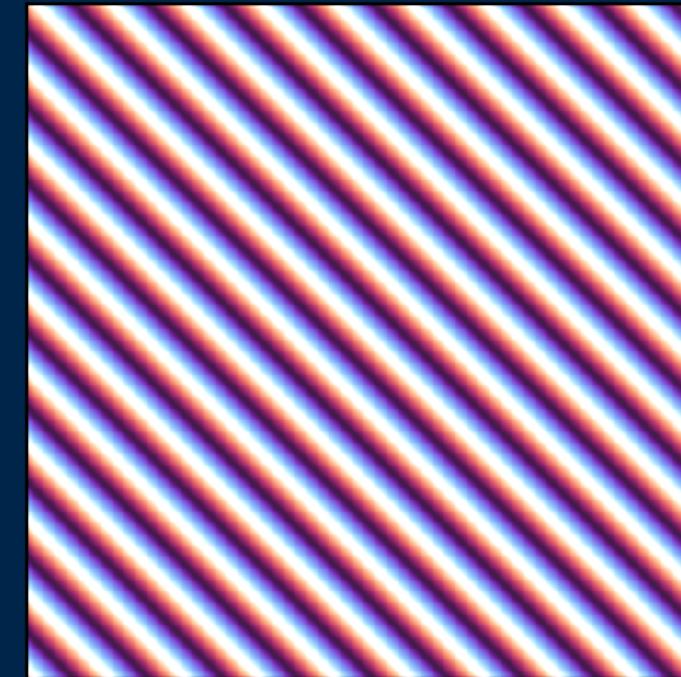
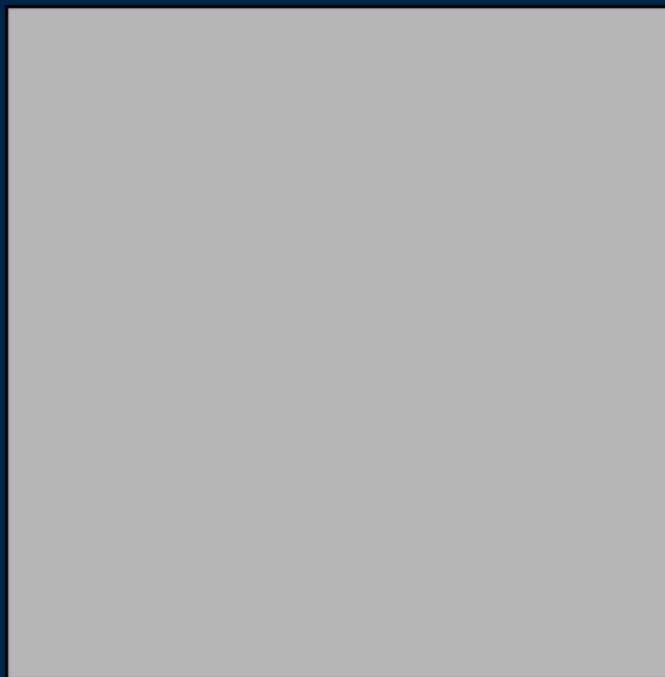
phase

Fourier oscillations

$$P(\vec{k}) = \langle |I \star e^{k \cdot x}|^2 \rangle$$



convolution



modulus

phase

local kernels (wavelets)

$$S_0 = \langle I \rangle$$

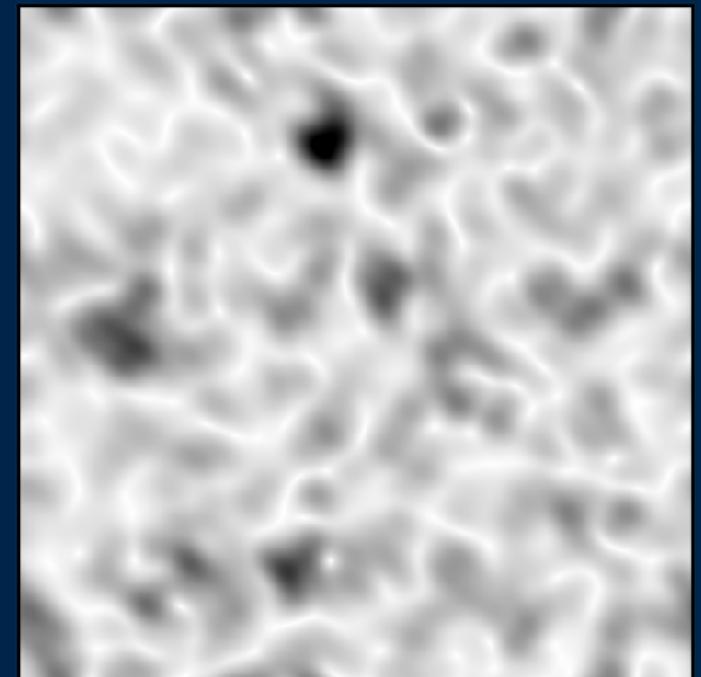


$$S_1(\vec{k}) = \langle |I \star \psi_k| \rangle$$

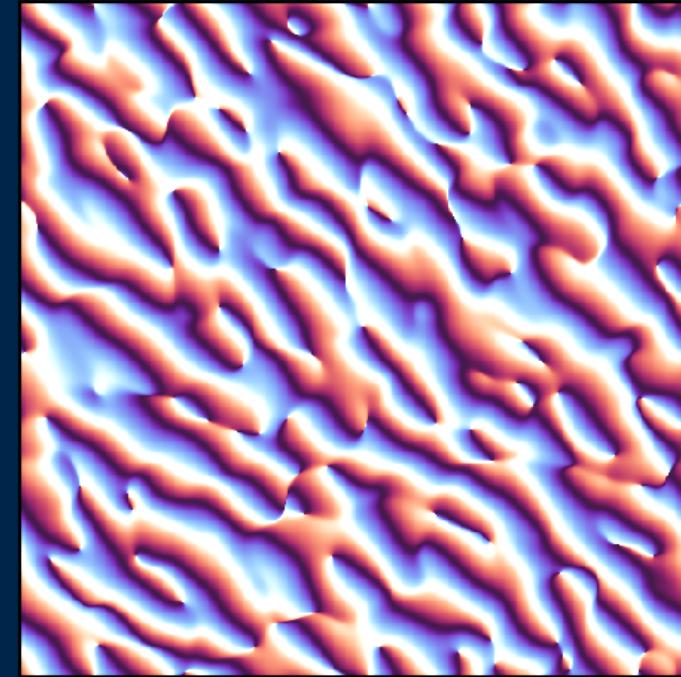
$$S_2(\vec{k}_1, \vec{k}_2) = \langle | |I \star \psi_{k_1}| \star \psi_{k_2} | \rangle$$

...

convolution

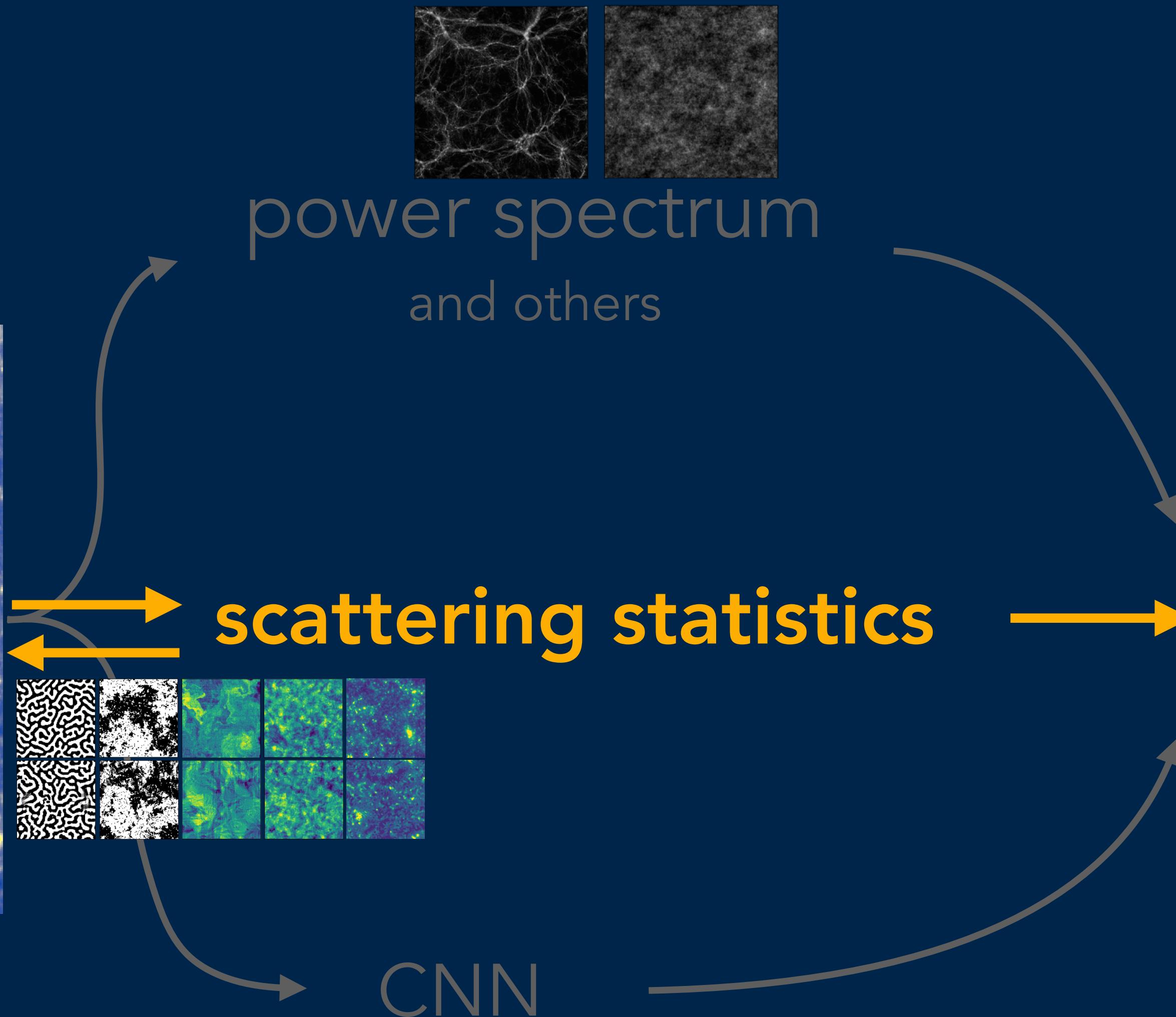
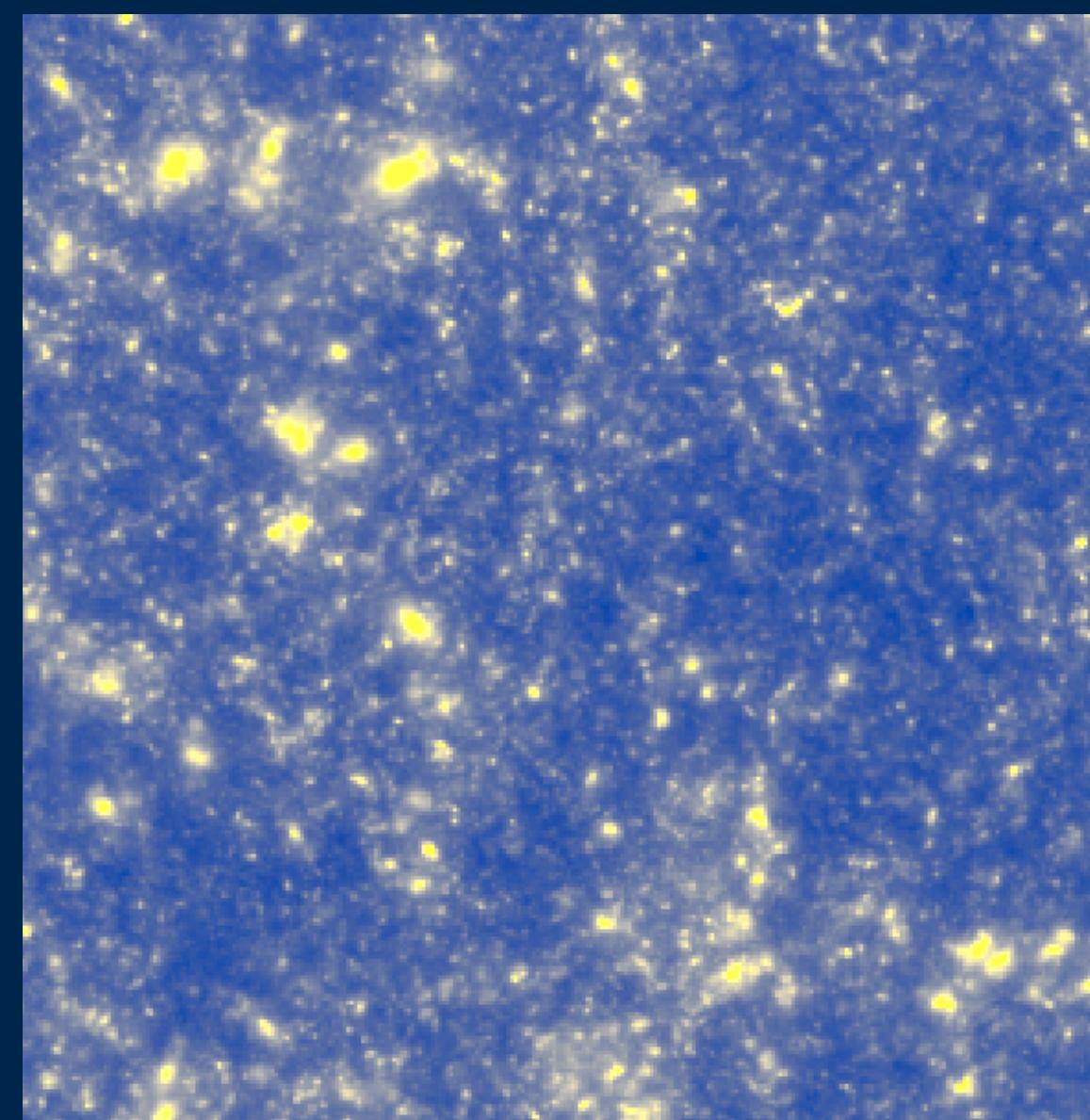


modulus

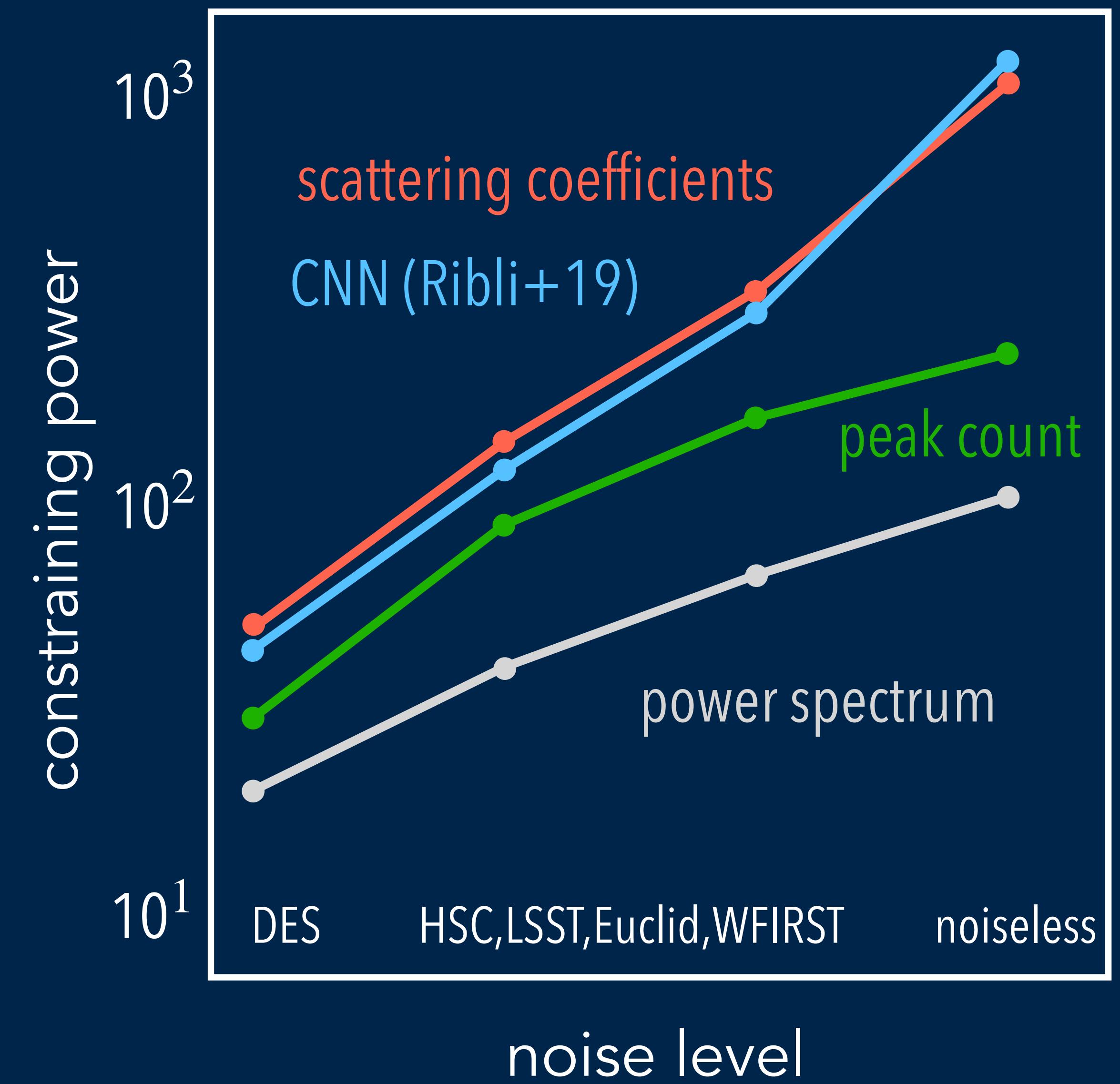
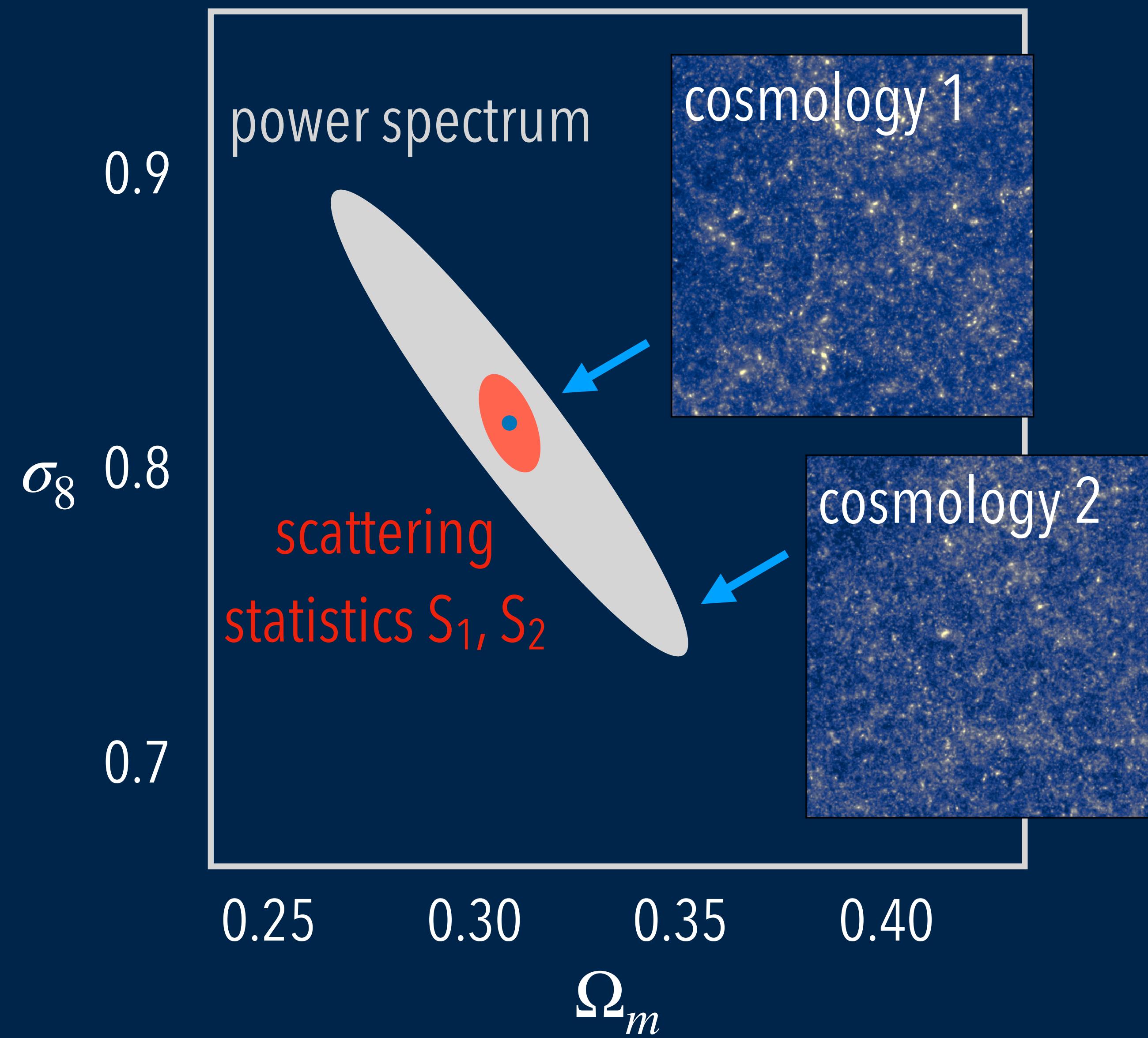


phase

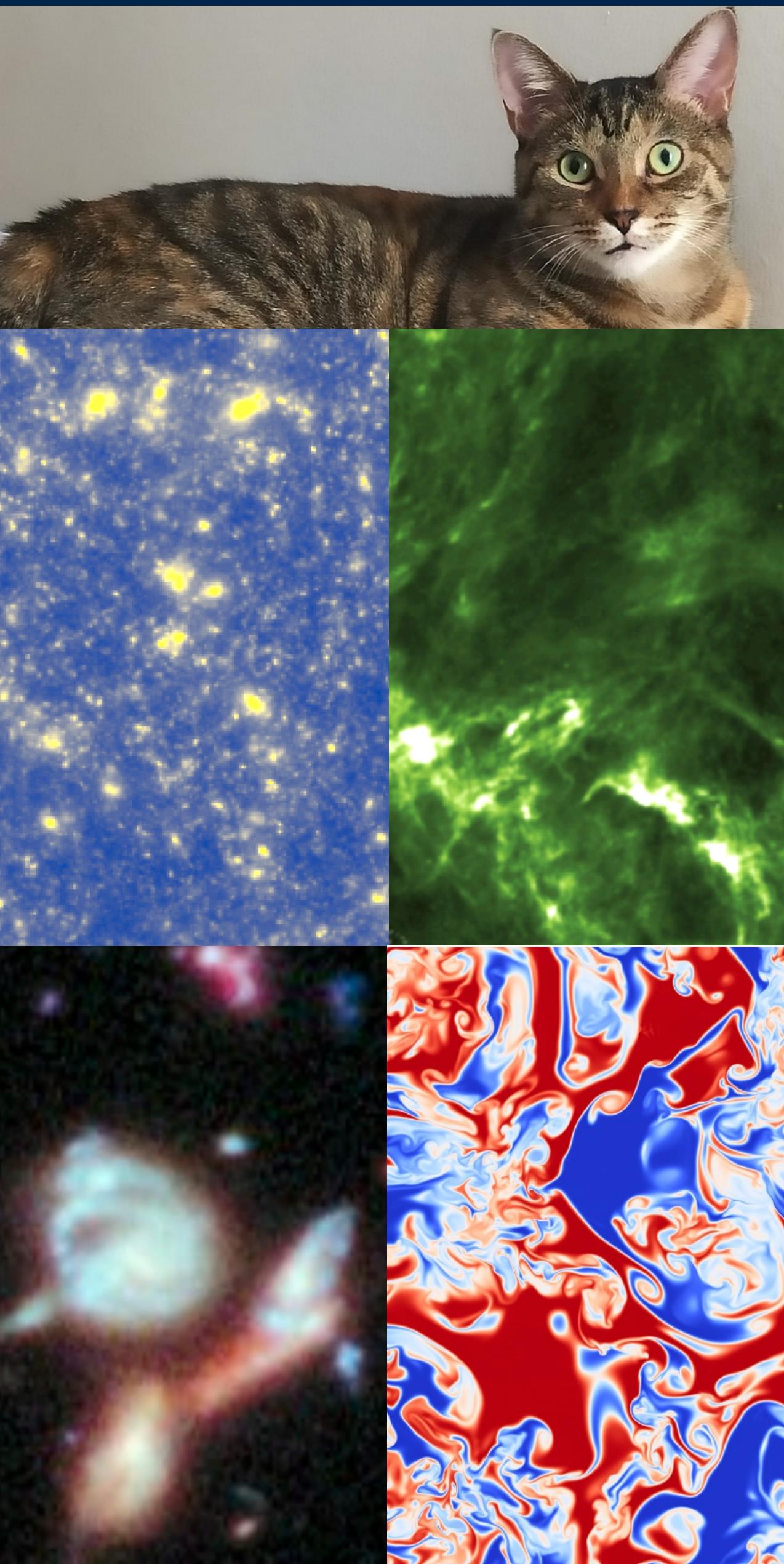
How do we characterize a field?



scale range: 1 arcmin to 3.5 deg



How do we characterize a field?



power spectrum

power spectrum

σ_8

scattering

Ω_m

physical
parameters

scattering transform

efficient, interpretable, robust

CNN

arXiv: 2006.08561

arXiv: 2103.09247

interpretations (coming soon)