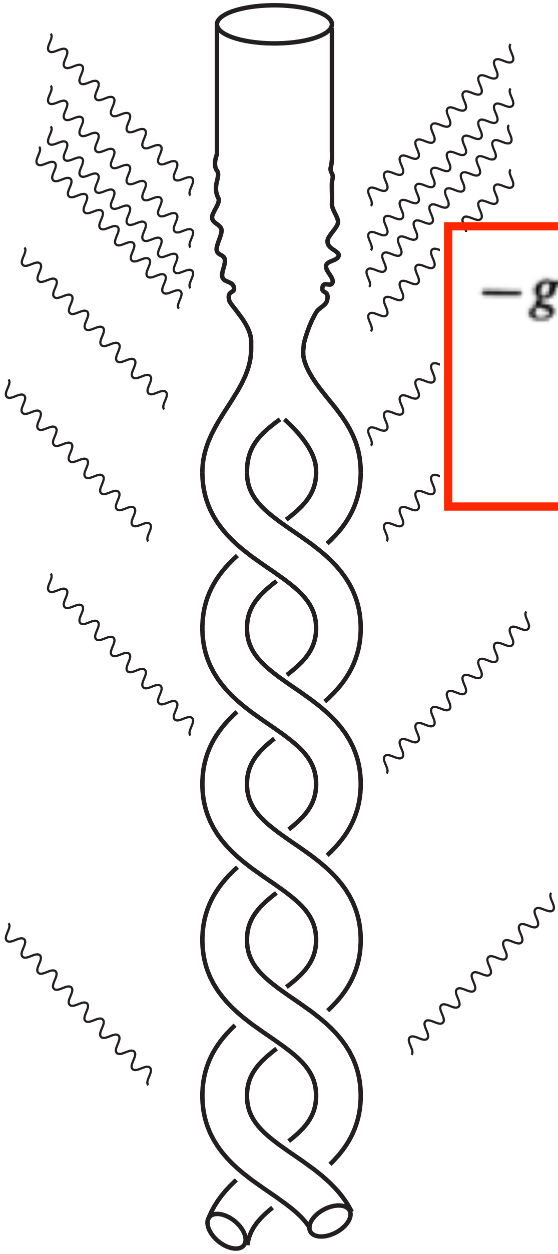
Radiation and Radiation-Reaction in Gravitational Scattering

Thibault Damour
Institut des Hautes Etudes Scientifiques

IHES
Institut des Hautes Études Scientifiques

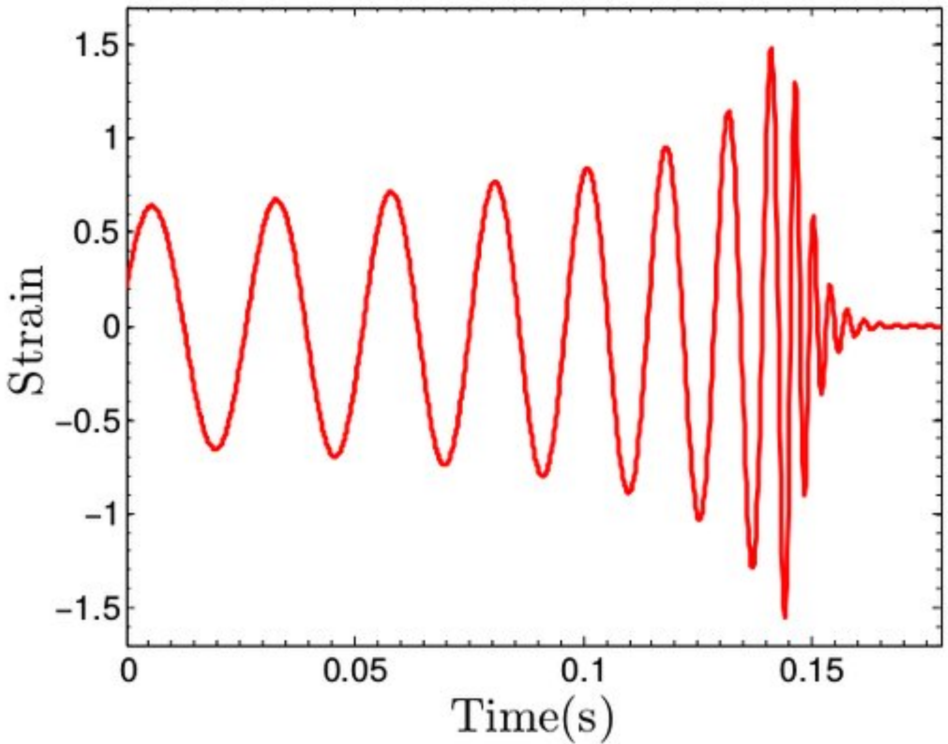
Annual Workshop of the Niels Bohr International Academy on Current Themes in High Energy Physics, Gravity and Cosmology, Niels Bohr Institute
21-25 August 2023, Copenhagen, Denmark

$$R_{\mu\nu} - \frac{1}{2} R\, g_{\mu\nu} = \frac{8\pi G}{c^4} T_{\mu\nu} \qquad R_{\mu\nu} = 0$$



$$ds^2 = g_{\mu\nu}(x^\lambda)\, dx^\mu dx^\nu$$

$$-g^{\mu\nu} g_{\alpha\beta,\,\mu\nu} + g^{\mu\nu} g^{\rho\sigma}\big(g_{\alpha\mu,\,\rho} g_{\beta\nu,\,\sigma} - g_{\alpha\mu,\,\rho} g_{\beta\sigma,\,\nu} + g_{\alpha\mu,\,\rho} g_{\nu\sigma,\,\beta} + g_{\beta\mu,\,\rho} g_{\nu\sigma,\,\alpha} - \frac{1}{2} g_{\mu\rho,\,\alpha} g_{\nu\sigma,\,\beta}\big) = 0$$



The GR 2-body problem

LO relativistic (« post-Newtonian ») corrections to the dynamics:
v^2/c^2+ GM/(c^2 r)=1PN (Lorentz-Droste'17, Einstein-Infeld-Hoffmann'38)

LO gravitational radiation: Einstein's 1916,1918 quadrupole formula

$$h_{ij} \simeq \frac{2G}{c^4 r} \ddot{Q}_{ij}^{TT}(t - r/c) \qquad \frac{dE_{\rm GW}}{dt} = \frac{G}{5c^5}\left(\frac{d^3 Q_{ij}}{dt^3}\right)^2$$

LO gravitational radiation reaction: O(1/c^5)=2.5PN
Burke'69 Thorne'69 Chandrasekhar-Esposito'70 with IR issues

1978 Hulse-Taylor binary pulsar « quadrupole controversy »
Avoiding PN IR divergencies by using a « post-Minkowskian » approach
(Westpfahl'79, Bel et al'81, TD Deruelle'81, TD'82)

$$a^i = A_0^i(z-z') + c^{-2}A_2^i(z-z',v,v') + c^{-4}A_4^i(z-z',v,v',S,S') + c^{-5}A_5^i(z-z',v-v') + O(c^{-6})$$

2PN (conservative)

2.5PN (radiation-reaction
G^2/c^5 + G^3/c^5)

1980's1990's LIGO-Virgo motivates higher-accuracy computations
of both dynamics and GW-emission of binary systems

Tools used for the GR 2-body pb

Post-Newtonian (PN) approximation (**expansion in 1/c; ie v^2/c^2 and GM/(c^2r)**)

Post-Minkowskian (PM) approximation (**expansion in G; ie in GM/(c^2b)**) and its recent **Worldline EFT avatars**

Multipolar post-Minkowskian (MPM) approximation theory to the GW emission of binary systems

Matched Asymptotic Expansions useful both for the motion of strongly self-gravitating bodies, and for the nearzone-wavezone matching

Gravitational Self-Force (SF): expansion in m1/m2, with « first law of BH mechanics » (LeTiec-Blanchet-Whiting'12,…)

Effective One-Body (EOB) Approach

Numerical Relativity (NR)

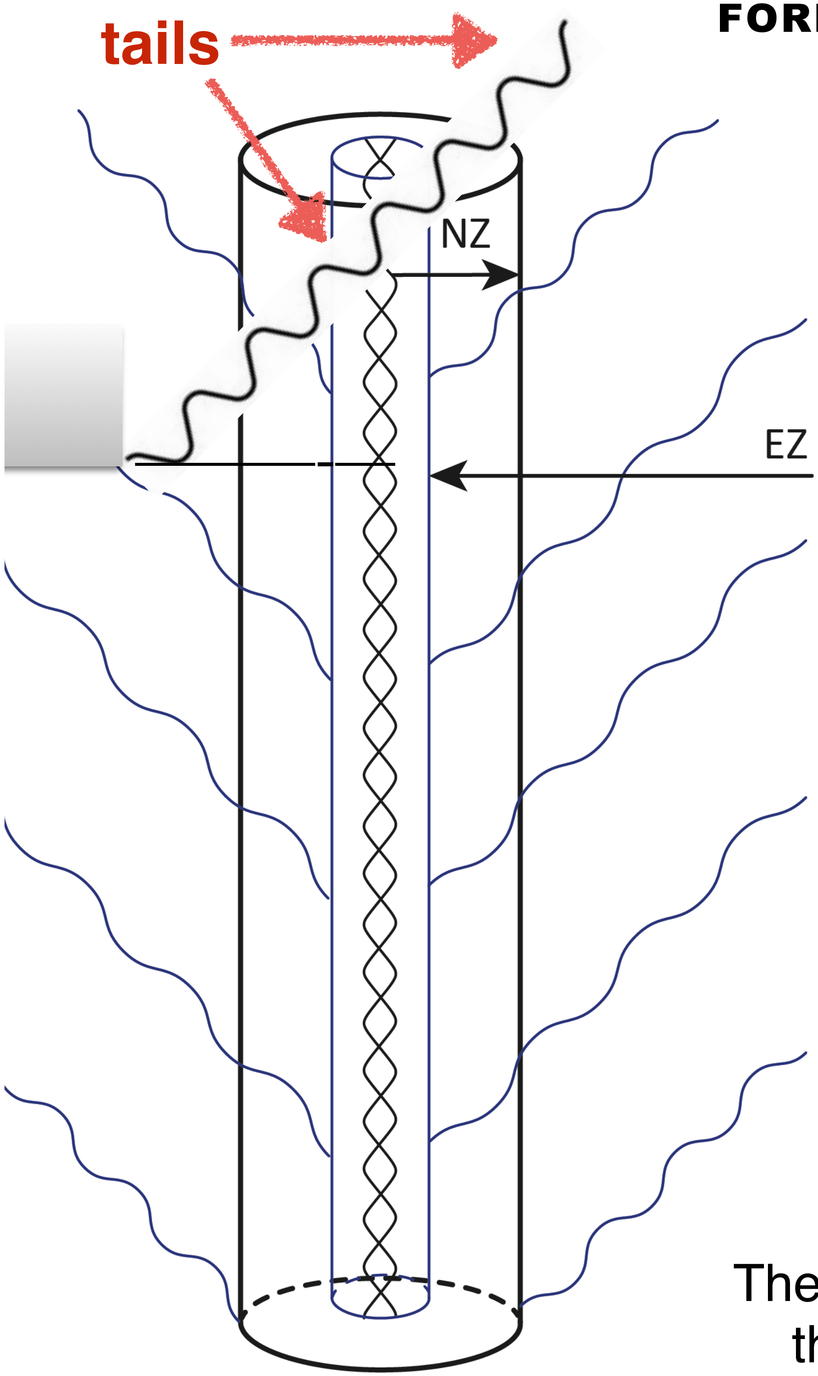
Effective Field Theory (EFT)

Recent

Quantum scattering amplitude aided by Double-Copy, Generalized Unitarity,« Feynman-integral Calculus » (IBP, DE, regions, reverse unitarity,…), Kosower-Maybee-O'Connell, eikonal, exponentiated representation **+ Worldline QFT**

Tutti Frutti method

GRAVITATIONAL WAVE GENERATION: MULTIPOLAR POST-MINKOWSKIAN FORMALISM (BLANCHET-DAMOUR-IYER)



Decomposition of space-time in various overlapping regions:

1. **near-zone: r << lambda : PN**
2. **exterior zone: r >> r_source: MPM**
3. **far wave-zone: Bondi-type expansion**

then matching between the zones

in exterior zone, iterative solution of Einstein's vacuum field equations by means of a double expansion in non-linearity and in multipoles, with crucial use of analytic continuation (complex B) for dealing with formal UV divergences at r=0

$$g = \eta + Gh_1 + G^2h_2 + G^3h_3 + \ldots,$$
$$\Box h_1 = 0,$$
$$\Box h_2 = \partial\partial h_1 h_1,$$
$$\Box h_3 = \partial\partial h_1 h_1 h_1 + \partial\partial h_1 h_2,$$
$$h_1 = \sum_\ell \partial_{i_1 i_2 \ldots i_\ell}\left(\frac{M_{i_1 i_2 \ldots i_\ell}(t-r/c)}{r}\right) + \partial\partial\ldots.\partial\left(\frac{\epsilon_{j_1 j_2 k} S_{k j_3 \ldots j_\ell}(t-r/c)}{r}\right),$$
$$h_2 = FP_B \Box^{-1}_{\text{ret}}\left(\left(\frac{r}{r_0}\right)^B \partial\partial h_1 h_1\right) + \ldots,$$
$$h_3 = FP_B \Box^{-1}_{\text{ret}} \ldots.$$

STF tensors encoding multipole moments

mass-type and spin-type multipole moments

The PN-matched MPM formalism has allowed to compute the GW emission to very high accuracy (Blanchet et al)

Perturbative computation of GW flux from binary system

- lowest order : Einstein 1918 Peters-Mathews 63
- 1 + (v^2/c^2) : Wagoner-Will 76
- … + (v^3/c^3) : Blanchet-Damour 92, Wiseman 93
- … + (v^4/c^4) : Blanchet-Damour-Iyer Will-Wiseman 95
- … + (v^5/c^5) : Blanchet 96
- … + (v^6/c^6) : Blanchet-Damour-Esposito-Farèse-Iyer 2004
- … + (v^7/c^7) : Blanchet
- … + (v^8/c^8) + (v^9/c^9) : Blanchet et al 2023

$$\nu = \frac{m_1 m_2}{(m_1+m_2)^2}$$

$$x = \left(\frac{v}{c}\right)^2 = \left(\frac{G(m_1+m_2)\Omega}{c^3}\right)^{\frac{2}{3}} = \left(\frac{\pi G(m_1+m_2)f}{c^3}\right)^{\frac{2}{3}}$$

LO quadrupole radiation

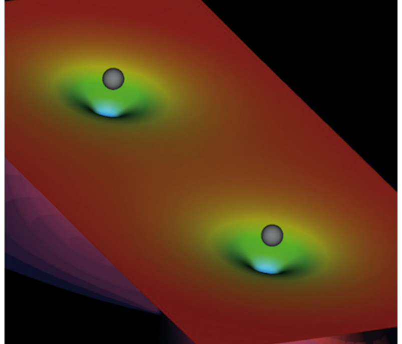
$$\begin{aligned}\mathcal{F} = \frac{32c^5}{5G}\nu^2 x^5 \Big\{ & 1 + \left(-\frac{1247}{336} - \frac{35}{12}\nu\right)x + 4\pi x^{3/2} + \left(-\frac{44711}{9072} + \frac{9271}{504}\nu + \frac{65}{18}\nu^2\right)x^2 + \left(-\frac{8191}{672} - \frac{583}{24}\nu\right)\pi x^{5/2} \\ & + \left[\frac{6643739519}{69854400} + \frac{16}{3}\pi^2 - \frac{1712}{105}\gamma_E - \frac{856}{105}\ln(16\,x) + \left(-\frac{134543}{7776} + \frac{41}{48}\pi^2\right)\nu - \frac{94403}{3024}\nu^2 - \frac{775}{324}\nu^3\right]x^3 \\ & + \left(-\frac{16285}{504} + \frac{214745}{1728}\nu + \frac{193385}{3024}\nu^2\right)\pi x^{7/2} \\ & + \Big[-\frac{323105549467}{3178375200} + \frac{232597}{4410}\gamma_E - \frac{1369}{126}\pi^2 + \frac{39931}{294}\ln 2 - \frac{47385}{1568}\ln 3 + \frac{232597}{8820}\ln x \\ & \quad + \left(-\frac{1452202403629}{1466942400} + \frac{41478}{245}\gamma_E - \frac{267127}{4608}\pi^2 + \frac{479062}{2205}\ln 2 + \frac{47385}{392}\ln 3 + \frac{20739}{245}\ln x\right)\nu \\ & \quad + \left(\frac{1607125}{6804} - \frac{3157}{384}\pi^2\right)\nu^2 + \frac{6875}{504}\nu^3 + \frac{5}{6}\nu^4\Big]x^4 \\ & + \Big[\frac{265978667519}{745113600} - \frac{6848}{105}\gamma_E - \frac{3424}{105}\ln(16\,x) + \left(\frac{2062241}{22176} + \frac{41}{12}\pi^2\right)\nu \\ & \quad - \frac{133112905}{290304}\nu^2 - \frac{3719141}{38016}\nu^3\Big]\pi x^{9/2} + \mathcal{O}(x^5)\Big\}. \end{aligned} \tag{4}$$

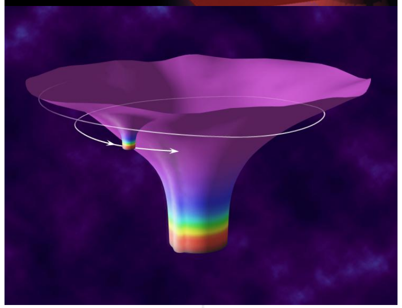
4PN (the x^4 term)

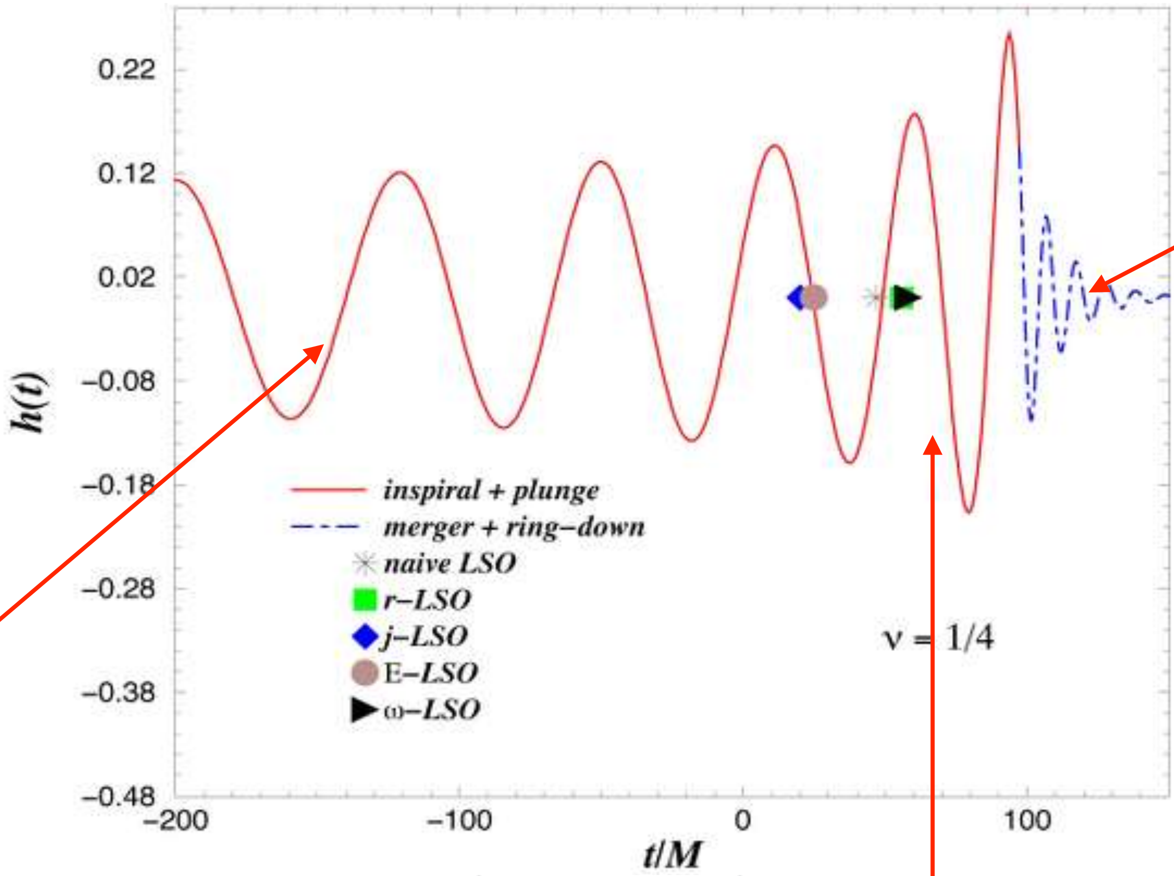
4.5PN (the $\pi x^{9/2}$ term)

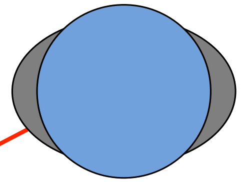
The Effective One-Body (EOB) approach to the GW signal emitted by the Merger of two Black Holes

Buonanno-TD'2000









Ringdown (BBH):
« vibration modes » of final BH (QNM);
perturbation of BHs à la Regge-Wheeler-Zerilli-Teukolsky +Vishveshwara

Inspiral:
perturbative computation of higher-order contributions to E=H and F (expansion in v^2/c^2 + tidal polarizability of NS)

Late inspiral, « plunge » and merger:
first estimated by the Effective One-Body method (AB-TD 2000)
later confirmed and improved by using numerical simulations (Pretorius…2005)

From EOB vs NR to EOB-NR waveforms

Buonanno-Cook-Pretorius 2007

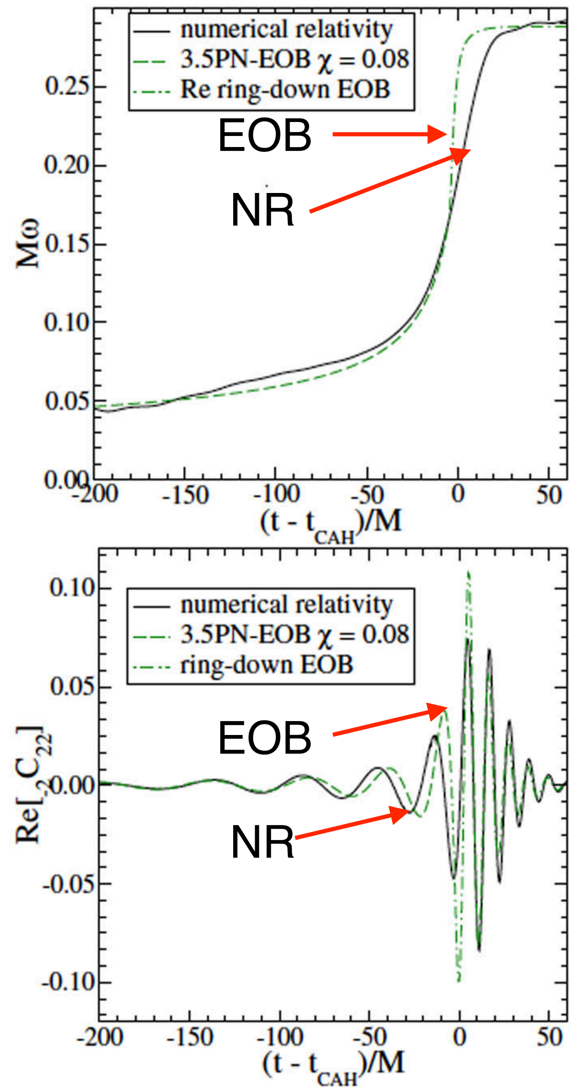
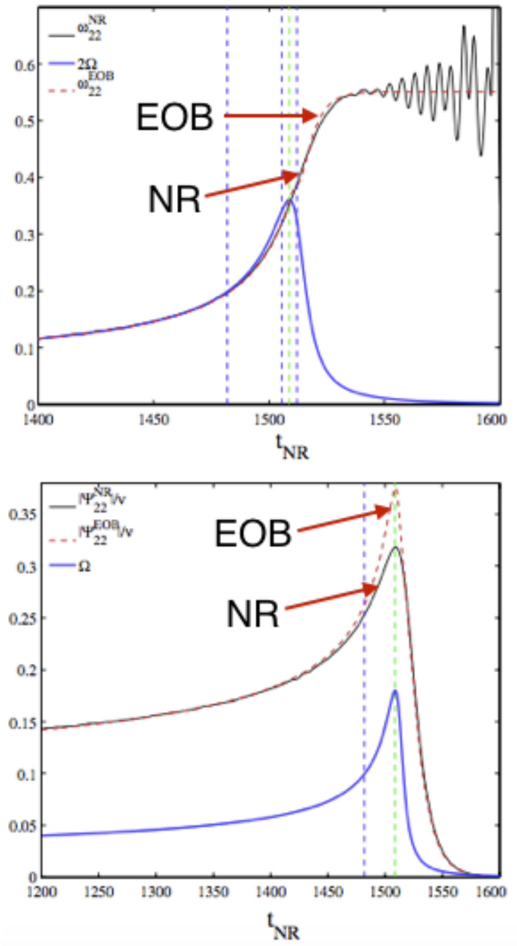
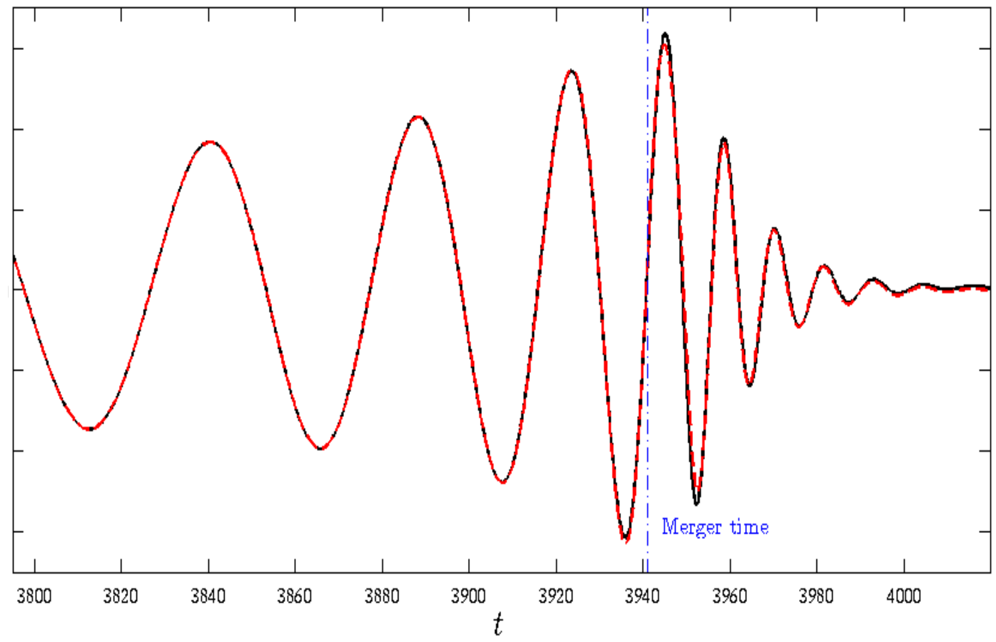


FIG. 21 (color online). We compare the NR and EOB frequency and Re[$_{-2}C_{22}$] waveforms throughout the entire inspiral–merger–ring-down evolution. The data refers to the $d = 16$ run.

TD-Nagar-Dorband-Pollney-Rezzolla 2008



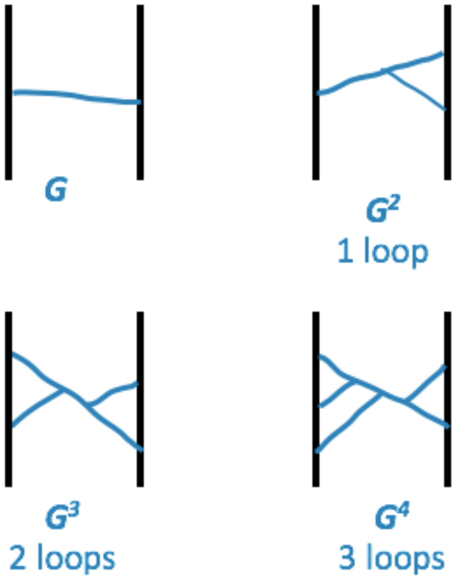
EOB-NR vs NR



Effective One-Body (EOB) approach: H + Rad-Reac Force

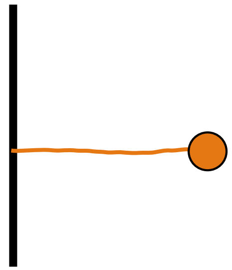
Historically rooted in QM: Brezin-Itzykson-ZinnJustin'70
eikonal scattering amplitude+ Wheeler's:`Think quantum mechanically'

Real 2-body system
(in the c.o.m. frame)



An effective particle of mass mu in some effective metric

$$\mu = \frac{m_1 m_2}{m_1 + m_2}$$



$$g_{\mathrm{eff}}^{\mu\nu}(X)$$

1:1 map

mass-shell constraint

$$0 = g_{\mathrm{eff}}^{\mu\nu}(X) P_\mu P_\nu + \mu^2 + Q(X, P)$$

Level correspondence
in the semi-classical limit:
Bohr-Sommerfeld ->
identification of
quantized action variables

$$J = \ell\hbar = \frac{1}{2\pi}\oint p_\varphi d\varphi$$

$$N = n\hbar = I_r + J$$

$$I_r = \frac{1}{2\pi}\oint p_r dr$$

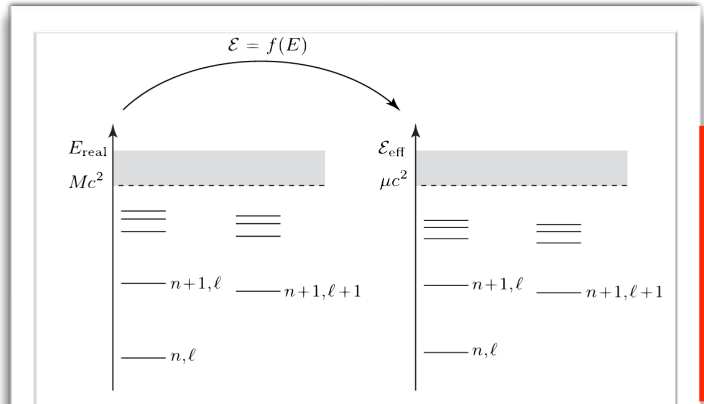


Figure 1: Sketch of the correspondence between the quantized energy levels of the real and effective conservative dynamics. n denotes the 'principal quantum

Crucial energy map

$$\mathcal{E}_{\mathrm{eff}} = \frac{(\mathcal{E}_{\mathrm{real}})^2 - m_1^2 - m_2^2}{2(m_1 + m_2)}$$

as functions of I_r and I_phi=J

New Angle of Attack on Two-Body Dynamics: Classical and/or Quantum Two-Body Scattering

TD 2016, 2017:

Gravitational scattering, post-Minkowskian approximation, and effective-one-body theory

High-energy gravitational scattering and the general relativistic two-body problem

A technique for translating the classical scattering function of two gravitationally interacting bodies into a corresponding (effective one-body) Hamiltonian description has been recently introduced [Phys. Rev. D **94**, 104015 (2016)]. Using this technique, we derive, for the first time, to second-order in Newton's constant (i.e. one classical loop) the Hamiltonian of two point masses having an arbitrary (possibly relativistic) relative velocity. The resulting (second post-Minkowskian) Hamiltonian is found to have a tame high-energy structure which we relate both to gravitational self-force studies of large mass-ratio binary systems, and to the ultra high-energy quantum scattering results of Amati, Ciafaloni and Veneziano. We derive several consequences of our second post-Minkowskian Hamiltonian: (i) the need to use special phase-space gauges to get a tame high-energy limit; and (ii) predictions about a (rest-mass independent) linear Regge trajectory behavior of high-angular-momenta, high-energy circular orbits. Ways of testing these predictions by dedicated numerical simulations are indicated. We finally indicate a way to connect our classical results to the quantum gravitational scattering amplitude of two particles, and we urge amplitude experts to use their novel techniques to compute the two-loop scattering amplitude of scalar masses, from which one could deduce the third post-Minkowskian effective one-body Hamiltonian.

one-loop
G^2

two-loop
G^3+G^4

Cheung-Rothstein-Solon 2018

From Scattering Amplitudes to Classical Potentials in the Post-Minkowskian Expansion

We combine tools from effective field theory and generalized unitarity to construct a map between on-shell scattering amplitudes and the classical potential for interacting spinless particles. For general relativity, we obtain analytic expressions for the classical potential of a binary black hole system at second order in the gravitational constant and all orders in velocity. Our results exactly match all known results up to fourth post-Newtonian order, and offer a simple check of future higher order calculations. By design, these methods should extend to higher orders in perturbation theory.

one-loop
G^2

Simple Map: ConservativeScattering angle <-> EOB dynamics

scattering angle, and its expansion in:

$$\frac{1}{j} = \frac{G m_1 m_2}{J}$$

$$\frac{1}{2}\chi = \Phi(E_{\rm real}, J; m_1, m_2, G)$$

TD'16-18
Bini-TD-Geralico'20

$$\frac{1}{2}\chi_{\rm class}(E,J) = \frac{1}{j}\chi_1(\hat{E}_{\rm eff},\nu) + \frac{1}{j^2}\chi_2(\hat{E}_{\rm eff},\nu) + O(G^3)$$

$$0 = g^{\mu\nu}_{\rm eff} P_\mu P_\nu + \mu^2 + Q$$

$$\chi_1(\hat{\mathcal{E}}_{\rm eff},\nu) = \frac{2\hat{\mathcal{E}}^2_{\rm eff} - 1}{\sqrt{\hat{\mathcal{E}}^2_{\rm eff} - 1}},$$

$$\chi_2(\hat{\mathcal{E}}_{\rm eff},\nu) = \frac{3\pi}{8}\frac{5\hat{\mathcal{E}}^2_{\rm eff} - 1}{\sqrt{1 + 2\nu(\hat{\mathcal{E}}_{\rm eff} - 1)}}$$

Westpfahl'85

$$g^{\mu\nu}_{\rm eff}$$

Schwarzschild metric M=m1+m2

$$Q = \left(\frac{GM}{R}\right)^2 q_2(E) + \left(\frac{GM}{R}\right)^3 q_3(E) + O(G^4)$$

equivalent to $g^{\mu\nu}_{\rm eff}(E)$

$$q_2(\hat{\mathcal{E}}_{\rm eff},\nu) = -\frac{4}{\pi}[\chi_2(\hat{\mathcal{E}}_{\rm eff},\nu) - \chi_2^{\rm Schw}(\hat{\mathcal{E}}_{\rm eff})].$$

$$q_3(\hat{\mathcal{E}}_{\rm eff},\nu) = \frac{4}{\pi}\frac{2\hat{\mathcal{E}}^2_{\rm eff} - 1}{\hat{\mathcal{E}}^2_{\rm eff} - 1}(\chi_2(\hat{\mathcal{E}}_{\rm eff},\nu) - \chi_2^{\rm Schw}(\hat{\mathcal{E}}_{\rm eff})) - \frac{\chi_3(\hat{\mathcal{E}}_{\rm eff},\nu) - \chi_3^{\rm Schw}(\hat{\mathcal{E}}_{\rm eff})}{\sqrt{\hat{\mathcal{E}}^2_{\rm eff} - 1}}.$$

$$\mathcal{E}_{\rm eff} = \frac{(\mathcal{E}_{\rm real})^2 - m_1^2 - m_2^2}{2(m_1 + m_2)},$$

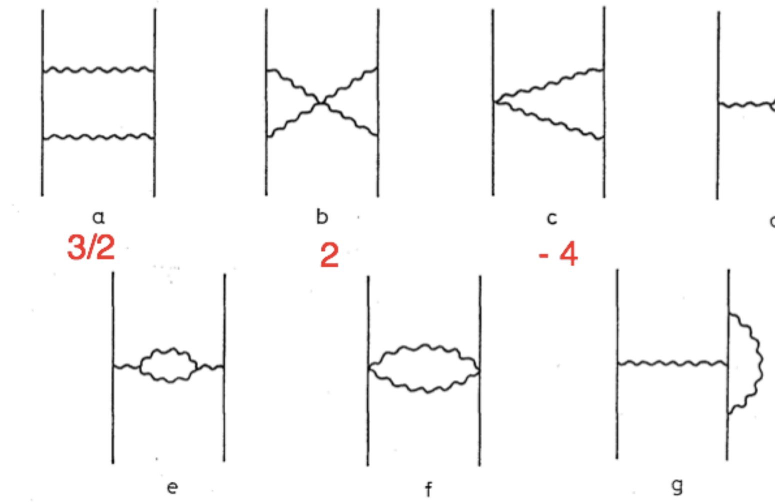
$$\frac{\mathcal{E}_{\rm eff}}{\mu} = \gamma = -\frac{p_1 \cdot p_2}{m_1 m_2}$$

$$q_2(\gamma,\nu) = \frac{3}{2}(5\gamma^2 - 1)\left(1 - \frac{1}{h(\gamma,\nu)}\right)$$

$$h(\gamma,\nu) = \sqrt{1 + 2\nu(\gamma - 1)}$$

Linear combinations of the scattering coefficients!

Quantum Scattering Amplitudes and 2-body Dynamics



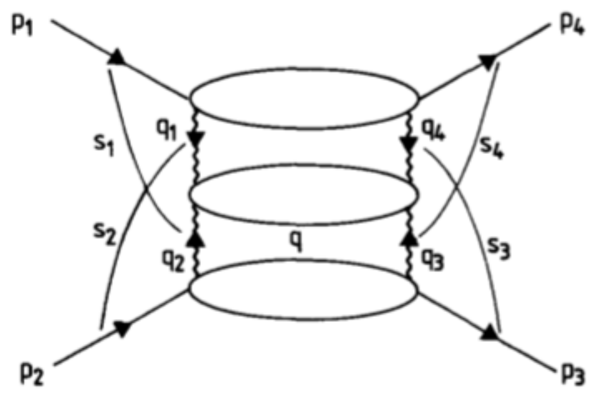


Fig. 3. The "H" diagram that provides the leading correction to the eikonal.

- Quantum Scattering Amplitudes —> Potential

one-graviton exchange :
Corinaldesi '56 '71,
Barker-Gupta-Haracz 66,
Barker-O'Connell 70, Hiida-Okamura72

Nonlinear: Iwasaki 71 [First post-Newtonian approx.],
Okamura-Ohta-Kimura-Hiida 73[2 PN]

Amati-Ciafaloni-Veneziano 1987-2008

Ultra-High-Energy (s >> M_Planck^2)
Four-graviton Scattering at 2 loops

Eikonal phase \delta in D=4
with one- and two-loop corrections
using the Regge-Gribov approach

confirmed by DiVecchia+'19

$$\delta = \frac{Gs}{\hbar}\left(\log\left(\frac{L_{IR}}{b}\right) + \frac{6\ell_s^2}{\pi b^2} + \frac{2G^2 s}{b^2}\left(1 + \frac{2i}{\pi}\log(\cdots)\right)\right)$$

Modern techniques for amplitudes (generalized unitarity; double copy; method of regions; IBPs; differential eqs; Bern, Dixon, Dunbar, Carrasco, Johansson, Cachazo et al., Bjerrum-Bohr et al., Cachazo-Guevara, Damgaard, Vanhove,…) can be used (Damour '17CheungRothsteinSolon'18) to improve the classical 2-body dynamics: need a quantum/classical dictionary.

Application to the ACV eikonal scattering phase (massless or ultra-relativistic scattering)

Amati-Ciafaloni-Veneziano'90+ Ciafaloni-Colferai'14+ Bern et al'20+ DiVecchia et al'20

$$\delta^{\text{eikonal}} = \frac{1}{\hbar}(\delta^{\text{R}} + i\delta^{\text{I}}) + \text{quantum corr.}$$

$$\frac{1}{2}\chi^{\text{eikonal}} = 2\frac{\gamma}{j} + \frac{16}{3}\frac{\gamma^3}{j^3} + \cdots$$

valid in the HE limit gamma-> infty

Using the chi—> Q dictionary this corresponds to the HE limits:

$$q_2^{\text{HE}} = \frac{15}{2}\gamma^2$$

$$q_3^{\text{HE}} = \gamma^2$$

i.e. an HE limit for the EOB mass-shell condition (TD'18)

$$0 = g_{\text{eff}}^{\mu\nu}(X)P_\mu P_\nu + \mu^2 + Q(X,P)$$

$$0 = g_{\text{Schw}}^{\mu\nu}P_\mu P_\nu + \left(\frac{15}{2}\left(\frac{GM}{R}\right)^2 + \left(\frac{GM}{R}\right)^3\right)P_0^2$$

Translating quantum scattering amplitudes into classical dynamical information (1)

The domain of validity of the Born-Feynman expansion

$$\mathcal{M}(s,t) = \mathcal{M}^{(\frac{G}{\hbar})}(s,t) + \mathcal{M}^{(\frac{G^2}{\hbar^2})}(s,t) + \cdots, \qquad \mathcal{M}^{(\frac{G}{\hbar})}(s,t) = 16\pi \frac{G}{\hbar} \frac{2(p_1 \cdot p_2)^2 - p_1^2 p_2^2}{-t}.$$

is

$$\frac{Gs}{\hbar v} \sim \frac{GE_1E_2}{\hbar v} \ll 1$$

while the domain of validity of classical scattering is (Bohr 1948)

$$\frac{Gs}{\hbar v} \sim \frac{GE_1E_2}{\hbar v} \gg 1$$

Amati-Ciafaloni-Veneziano faced this issue by assuming **eikonalization in b space**

$$\widetilde{\mathcal{A}}(s,b) = \int \frac{d^{D-2}q}{(2\pi)^{D-2}} \frac{\mathcal{A}(s,q^2)}{4pE} e^{-ib\cdot q}$$

$$1 + i\widetilde{\mathcal{A}}(s,b) = (1 + 2i\Delta(s,b))\, e^{2i\delta(s,b)}$$

classical phase

$$i\,\frac{\mathcal{A}(s,Q^2)}{4pE} = \int d^{D-2}b \left(e^{2i\delta(s,b)} - 1\right) e^{ib\cdot Q}$$

$$2\delta(s,b) = \frac{\Delta S_r(s,J)}{\hbar}$$

subtracted radial action of potential scattering

total classical momentum transfer:

$$Q^\mu = -\frac{\partial\, \mathrm{Re}\, 2\delta(s,b)}{\partial b^\mu}$$

Other approaches to extracting classical info: Bern et al., KMOC, Porto, Plefka, Damgaard,…

Translating quantum scattering amplitudes into classical dynamical information (2)

Damour'17: EOB potential Q(R,E) or W(R,E)
Cheung-Rothstein-Solon'18, Bern et al'19
different EFT potential V(R,P^2) and methods for
taking the classical limit at the integrand level,
and extracting the « classical part » of the scattering amplitude

EOB

$$Q^E(u,\mathcal{E}_{\rm eff}) = u^2 q_2(\mathcal{E}_{\rm eff}) + u^3 q_3(\mathcal{E}_{\rm eff}) + u^4 q_4^E(\mathcal{E}_{\rm eff}) + O(G^5)$$

$$w(r,p_\infty) = \frac{w_1(\gamma)}{r} + \frac{w_2(\gamma)}{r^2} + \frac{w_3(\gamma)}{r^3} + \frac{w_4(\gamma)}{r^4} + \cdots$$

EFT

$$H(\mathbf{P},\mathbf{X}) = \sqrt{m_1^2+\mathbf{P}^2} + \sqrt{m_2^2+\mathbf{P}^2} + V(R,\mathbf{P}^2)$$

$$V(R,\mathbf{P}^2) = G\frac{c_1(\mathbf{P}^2)}{R} + G^2\frac{c_2(\mathbf{P}^2)}{R^2} + G^3\frac{c_3(\mathbf{P}^2)}{R^3} + \cdots$$

non-relativistic potential scattering !

$$-\widehat{\hbar}^2\Delta_{\mathbf{x}}\psi(\mathbf{x}) = \left[p_\infty^2 + \frac{w_1}{r} + \frac{w_2}{r^2} + \frac{w_3}{r^3} + O\left(\frac{1}{r^4}\right)\right]\psi(\mathbf{x})$$

$$\mathcal{M}^{QFT}_{\rm classical} = \frac{8\pi G s}{\hbar} f^{EOB} = \mathcal{M}^{EFT}$$

issue: extracting the « classical » piece of the amplitude

Scattering Amplitudes and the Conservative Hamiltonian for Binary Systems at Third Post-Minkowskian Order

Zvi Bern,[1] Clifford Cheung,[2] Radu Roiban,[3] Chia-Hsien Shen,[1] Mikhail P. Solon,[2] and Mao Zeng[4]

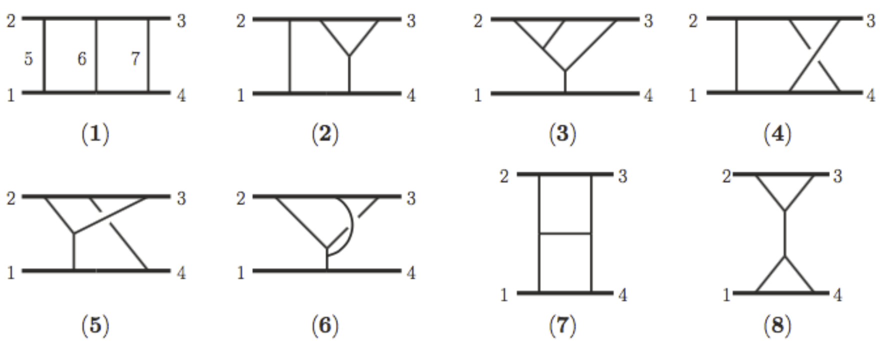
[1]*Mani L. Bhaumik Institute for Theoretical Physics, University of California at Los Angeles, Los Angeles, California 90095, USA*
[2]*Walter Burke Institute for Theoretical Physics, California Institute of Technology, Pasadena, California 91125*
[3]*Institute for Gravitation and the Cosmos, Pennsylvania State University, University Park, Pennsylvania 16802, USA*
[4]*Institute for Theoretical Physics, ETH Zürich, 8093 Zürich, Switzerland*

(Received 28 January 2019; published 24 May 2019)

two-loop level G^3

We present the amplitude for classical scattering of gravitationally interacting massive scalars at third post-Minkowskian order. Our approach harnesses powerful tools from the modern amplitudes program such as generalized unitarity and the double-copy construction, which relates gravity integrands to simpler gauge-theory expressions. Adapting methods for integration and matching from effective field theory, we extract the conservative Hamiltonian for compact spinless binaries at third post-Minkowskian order. The resulting Hamiltonian is in complete agreement with corresponding terms in state-of-the-art expressions at fourth post-Newtonian order as well as the probe limit at all orders in velocity. We also derive the scattering angle at third post-Minkowskian order and find agreement with known results.

the eight 2-loop diagrams contributing to the O(G^3/r^3) classical potential



two-loop level

$$\mathcal{M}_3 = \frac{\pi G^3 \nu^2 m^4 \log \boldsymbol{q}^2}{6\gamma^2 \xi}\Bigg[3 - 6\nu + 206\nu\sigma - 54\sigma^2 + 108\nu\sigma^2 + 4\nu\sigma^3 - \frac{48\nu(3+12\sigma^2-4\sigma^4)\operatorname{arcsinh}\sqrt{\frac{\sigma-1}{2}}}{\sqrt{\sigma^2-1}} - \frac{18\nu\gamma(1-2\sigma^2)(1-5\sigma^2)}{(1+\gamma)(1+\sigma)}\Bigg] + \frac{8\pi^3 G^3 \nu^4 m^6}{\gamma^4 \xi}\big[3\gamma(1-2\sigma^2)(1-5\sigma^2)F_1 - 32m^2\nu^2(1-2\sigma^2)^3 F_2\big], \quad (8)$$

arcsinh

3PM computation (Bern-Cheung-Roiban-Shen-Solon-Zeng'19)

using a combination of techniques: generalized unitarity; BCJ double-copy; 2-loop amplitude of quasi-classical diagrams; **EFT transcription** (Cheung-Rothstein-Solon'18); **resummation of PN-expanded integrals for potential-gravitons**

$$\chi_3^{\rm cons} = \chi_3^{\rm Schw} - \frac{2\nu\sqrt{\gamma^2-1}}{h^2(\gamma,\nu)}\bar{C}^{\rm cons}(\gamma)$$

G^3 contrib. to H_EOB

$$q_3^{\rm cons} = \frac{3}{2}\frac{(2\gamma^2-1)(5\gamma^2-1)}{\gamma^2-1}\left(\frac{1}{h(\gamma,\nu)}-1\right) + \frac{2\nu}{h^2(\gamma,\nu)}\bar{C}^{\rm cons}(\gamma)$$

$$\overline{C}^{\rm cons}(\gamma) = \frac{2}{3}\gamma(14\gamma^2+25) + 2(4\gamma^4-12\gamma^2-3)\frac{\mathcal{A}(v)}{\sqrt{\gamma^2-1}}$$

$$h(\gamma,\nu) \equiv \frac{\sqrt{s}}{M} = \sqrt{1+2\nu(\gamma-1)}$$

$$\mathcal{A}(v) \equiv \text{arctanh}(v) = \frac{1}{2}\ln\frac{1+v}{1-v} = 2\,\text{arcsinh}\sqrt{\frac{\gamma-1}{2}}$$

puzzling HE limits when compared to ACV and Akcay et al'12

$$\frac{1}{2}\chi^{\rm cons} = 2\frac{\gamma}{j} + (12 - 8\ln(2\gamma))\frac{\gamma^3}{j^3} + O(G^4)$$

$q_3^{\rm cons} \approx +8\ln(2\gamma)\gamma^2$ instead of $q_3^{\rm ACV} \approx +1\gamma^2$

confirmations: 5PN (Bini-TD-Geralico'19); 6PN (Blümlein-Maier-Marquard-Schäfer'20, Bini-TD-Geralico'20); 3PM (Cheung-Solon'20, Kälin-Porto'20)

Conservative and Radiative Aspects of the Dynamics

selected by Bern's integration method of regions

1PM

2PM

$$\dot{\mathbf{p}}_a \sim G\left(1 + \frac{1}{c^2} + \frac{1}{c^4} + \frac{1}{c^6} + \frac{1}{c^8} + \frac{1}{c^{10}}\right) +$$
$$+ G^2\left(\frac{1}{c^2} + \frac{1}{c^4} + \frac{1}{c^5} + \frac{1}{c^6} + \frac{1}{c^7} + \frac{1}{c^8} + \frac{1}{c^9} + \frac{1}{c^{10}}\right) +$$
$$+ G^3\left(\frac{1}{c^4} + \frac{1}{c^5} + \frac{1}{c^6} + \frac{1}{c^7} + \frac{1}{c^8} + \frac{1}{c^9} + \frac{1}{c^{10}}\right) +$$
$$+ G^4\left(\frac{1}{c^6} + \frac{1}{c^7} + \frac{1}{c^8} + \frac{1}{c^9} + \frac{1}{c^{10}}\right) +$$
$$+ G^5\left(\frac{1}{c^8} + \frac{1}{c^9} + \frac{1}{c^{10}}\right) +$$
$$+ G^6 \frac{1}{c^{10}}$$

1PN

2PN

2.5PN LO rad reac

4PN (G^4,G^5) separation well-defined

5PN same order as F_radreac^2!

Black: time-even and conservative

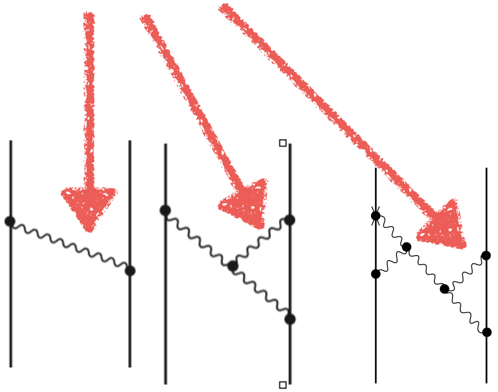
Red: time-odd and dissipative

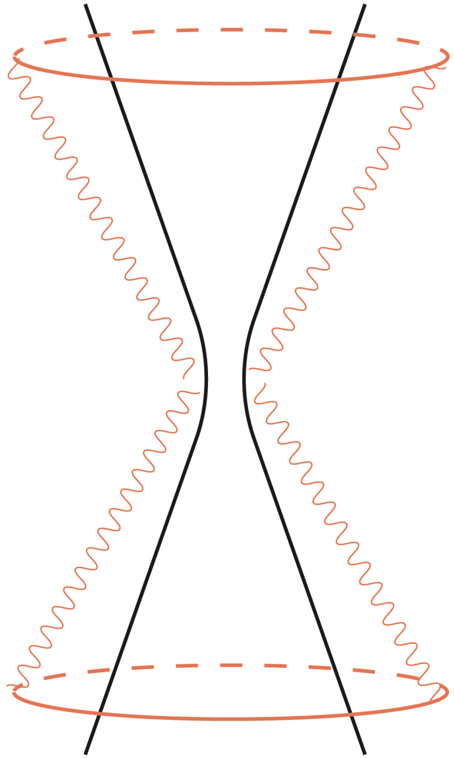
Blue: nonlocal-in-time but decomposable in conservative and dissipative

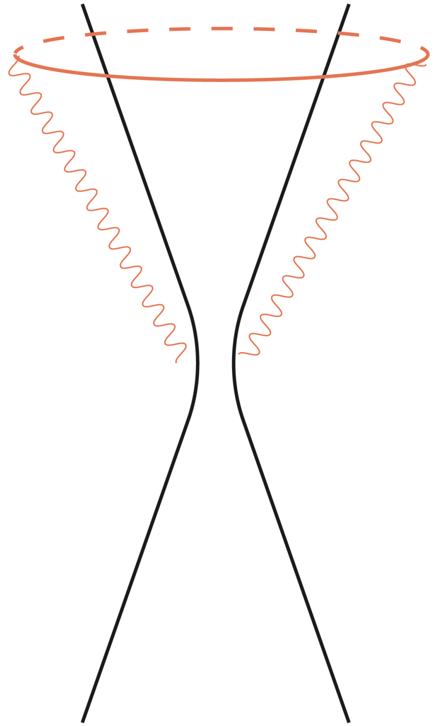
Purple: ambiguous in various ways ??

Conservative vs Radiation-reacted
Classical Gravitational Scattering

Fokker-Wheeler-Feynman
conservative action using
G_sym=1/2(G_ret+Gadv)
=Re[G_F]=PV(1/p^2)







Radiation-reaction effects enter scattering at G^3/c^5 (Bini-TD'12)

$$\frac{1}{2}\chi^{\rm rad} = +\frac{8G^3}{5c^5}\frac{m_1^3 m_2^3}{J^3}\nu v^2 + \cdots$$

Radiation-reaction effects in scattering play a crucial role at **high-energy**

(DiVecchia-Heissenberg-Russo-Veneziano'20, TD'21, Hermann-Parra-Martinez-Ruf-Zeng'21,….)

they resolve the O(G^3) puzzle of the discrepancy between the HE limit of Amati-Ciafaloni-Veneziano'90(+ Ciafaloni-Colferai'14), and the G^3 result of Bern et al'19,20

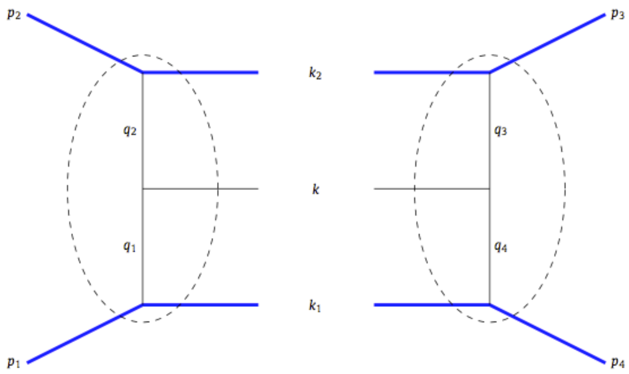
Universality of ultra-relativistic gravitational scattering from analyticity/crossing (DiVecchia-Heissenberg-Russo-Veneziano'20)

ultra-relativistic eikonal phase: $\delta(s,b) = \delta_0(s,b) + \delta_2(s,b)$

$$\mathrm{Re}(2\delta_2) = \frac{\pi}{2\log s}\,\mathrm{Im}(2\delta_2) - \frac{\delta_0}{s}(\nabla 2\delta_0)^2 + \mathcal{O}\left(\frac{1}{\log s}\right)$$

IR finite

IR divergent



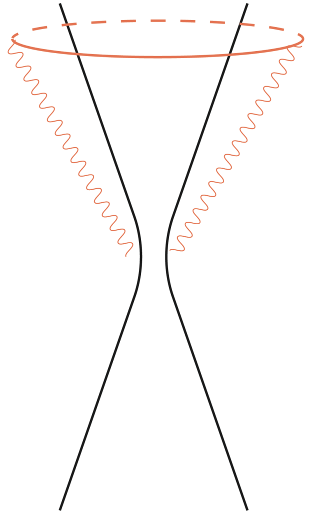
$$\mathrm{Im}\,\widetilde{A_2^{(3p)}}(s,b) \simeq \frac{1}{2s}\frac{(8G_N s)^3 \log s\,\Gamma^3(1-\epsilon)}{16(\pi b^2)^{1-3\epsilon}}\left[-\frac{1}{4\epsilon}+\frac{1}{2}+\mathcal{O}(\epsilon)\right]$$

$$\mathrm{Re}(2\delta_2) \simeq \frac{4G_N^3 s^2}{\hbar b^2}$$

universality of HE result = ACV, thanks to radiative effects

+ DiVecchia-Heissenberg-Russo-Veneziano'21: **Radiation Reaction from Soft Theorems**

Radiation-Reaction Contribution to the (transverse) Classical Scattering Angle at G^3 (TD 2010.01641)



$$\chi^{\rm tot} = \chi^{\rm cons} + \chi^{\rm rad}$$

where, to first order in Rad-Reac, one has (Bini-TD'12)

linear response formula

$$\chi^{\rm rad}(E,J) = -\frac{1}{2}\frac{\partial \chi^{\rm cons}}{\partial E}E^{\rm rad} - \frac{1}{2}\frac{\partial \chi^{\rm cons}}{\partial J}J^{\rm rad}$$

O(G^2) [TD-Deruelle'81]

chi^cons=O(G^1)

O(G^3)

O(G^4)

O(G^3)

$$h_{ij}^{\rm TT} = \frac{f_{ij}(t-r,\theta,\phi)}{r} + O\left(\frac{1}{r^2}\right)$$

$$J_k^{\rm rad} = \frac{\epsilon_{kij}}{16\pi G}\int du\, d\Omega \left[f_{ia}\partial_u f_{ja} - \frac{1}{2}x^i \partial_j f_{ab}\partial_u f_{ab}\right]$$

DeWitt'71, Thorne'80
Kovacs-Thorne'77, Bel et al'81,
Westpfahl'85

$$\mathcal{I}(v) = -\frac{16}{3} + \frac{2}{v^2} + \frac{2(3v^2-1)}{v^3}\mathcal{A}(v)$$

$$\mathcal{A}(v) \equiv \mathrm{arctanh}(v) = \frac{1}{2}\ln\frac{1+v}{1-v}$$

$$\frac{1}{2}\chi^{\rm rad}(\gamma,j,\nu) = +\frac{\nu}{h^2(\gamma,\nu)j^3}(2\gamma^2-1)^2\,\mathcal{I}(v) + O(G^4)$$

$$\frac{1}{2}(\chi^{\rm cons}+\chi^{\rm rad}) = 2\frac{\gamma}{j} + \frac{16}{3}\frac{\gamma^3}{j^3} = \chi^{\rm ACV}$$

Radiation-reaction and angular momentum loss at the second post-Minkowskian order

(Bini-TD'22)

The linear-response formula assumes a balance between E and J GW-losses at infinity and the mechanical E and J of the 2-body system. It also relies on using the « standard » Peters-DeWitt-Thorne formula for the GW J-flux:

$$J_k^{\rm rad} = \frac{\epsilon_{kij}}{16\pi G} \int du\, d\Omega \left[f_{ia}\partial_u f_{ja} - \frac{1}{2} x^i \partial_j f_{ab} \partial_u f_{ab} \right]$$

which is valid only in the (instantaneous) c.m. frame of the system, and which crucially depends on the zero-frequency part of the waveform

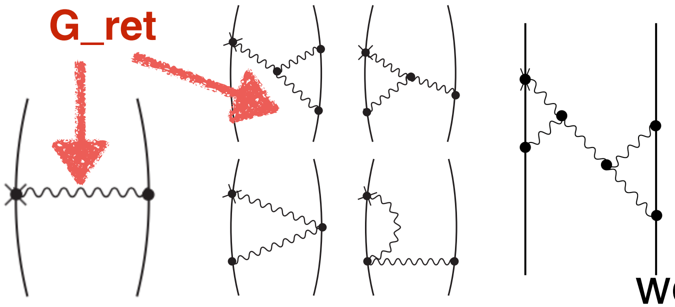
$$f_{ij}(u,\theta,\phi) = G f_{ij}^{(1)}(\theta,\phi) + G^2 f_{ij}^{(2)}(u,\theta,\phi) + O(G^3)$$

This raises several subtle issues: Ashtekar et al'20, Veneziano-Vilkovisky'22, Riva-Vernizzi, Riva-Vernizzi-Wong'23

Similarly to the resolution of the binary-pulsar « quadrupole controversy », it is useful to clarify the physics by a direct **EOM-based approach**, fully within a **retarded** PM approach (Bini-TD'22)

$$m_a \frac{d^2 z_a^\mu(\tau_a)}{d\tau_a^2} = F^\mu_{a\rm R}[z_a(\tau_a), u_a(\tau_a); z_{b\rm R}(\tau_a), u_{b\rm R}(\tau_a)]$$

Considering an auxiliary Fokker-Wheeler-Feynman-type PM dynamics allows one to define Noether« conserved » quantities



$$P_\mu^{\rm sys}(\tau_1,\tau_2) = P_\mu^{\rm kin}(\tau_1,\tau_2) + P_\mu^{\rm int}(\tau_1,\tau_2),$$
$$J_{\mu\nu}^{\rm sys}(\tau_1,\tau_2) = J_{\mu\nu}^{\rm kin}(\tau_1,\tau_2) + J_{\mu\nu}^{\rm int}(\tau_1,\tau_2).$$

$$I = -\sum_a m_a \int d\tau_a (-\dot z_{a\mu} \dot z_a^\mu)^{1/2} + \sum_{a<b} \iint d\tau_a d\tau_b \Lambda_{ab} + \sum_{a<b<c} \iiint d\tau_a d\tau_b d\tau_c \Lambda_{abc} + O(G^3),$$

Poincaré-covariant interaction terms Noether-derived from the Fokker action

well-defined O(G^2) PM rad-reac force

$$dP^\mu_{\rm sys} = \sum_a \mathcal{F}^\mu_{a\,\rm rr}(\tau_a) \frac{d\tau_a}{d\sigma} d\sigma + O(G^3),$$

$$dJ^{\mu\nu}_{\rm sys} = \sum_a (z_a^\mu \mathcal{F}^\nu_{a\,\rm rr}(\tau_a) - z_a^\nu \mathcal{F}^\mu_{a\,\rm rr}(\tau_a)) \frac{d\tau_a}{d\sigma} d\sigma + O(G^3).$$

Poincaré-covariant final results: $[P^{\rm sys}]^{+\infty}_{-\infty} = O(G^3).$

$$[J^{\mu\nu}_{\rm sys}]^{+\infty}_{-\infty} = \frac{G^2 m_1 m_2}{b_{12}^2} c_I(v) \mathcal{I}(v) [(p_1 - p_2) \wedge b_{12}]^{\mu\nu} + O(G^3).$$

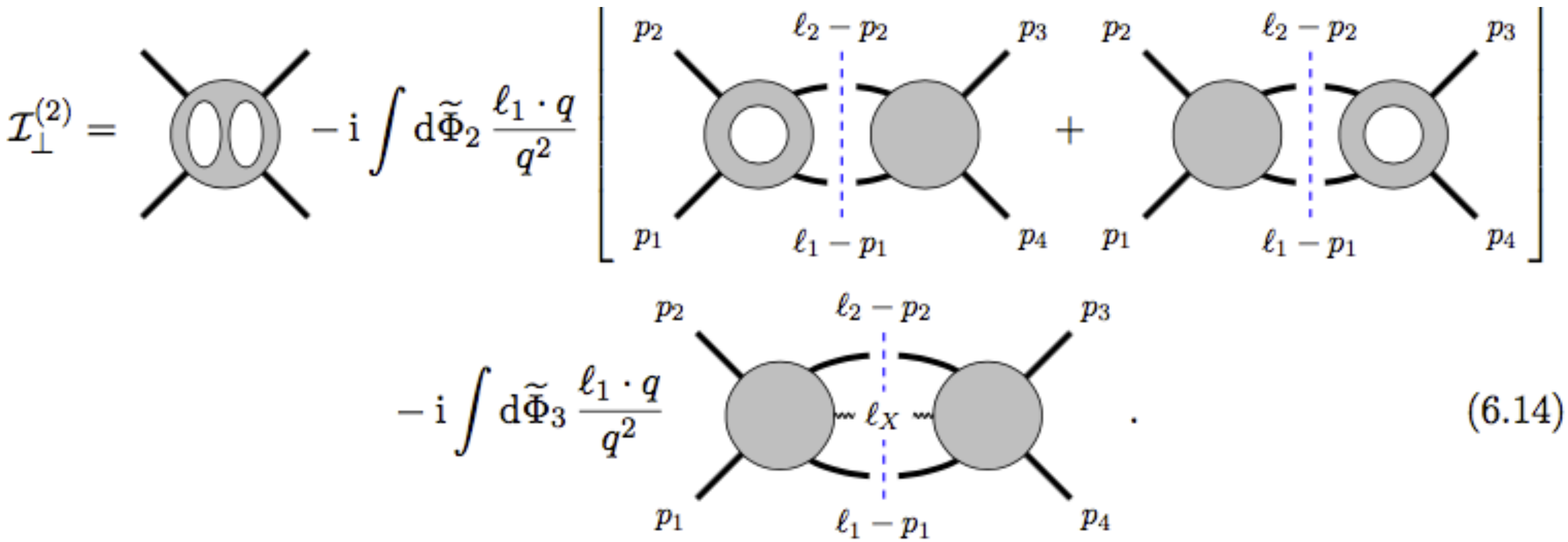
Translating quantum scattering amplitudes into classical dynamical information (3)

Kosower-Maybee-O'Connell'19 formalism for any observable O

$\Delta O = \langle \text{out}|\mathbb{O}|\text{out}\rangle - \langle \text{in}|\mathbb{O}|\text{in}\rangle$ with |out> = S |in> and S=1+i T

$$\Delta O = \langle \text{in}|i[O,T]|\text{in}\rangle + \langle \text{in}|T^\dagger[O,T]|\text{in}\rangle$$

Hermann-Parra-Martinez-Ruf-Zeng'21 making use of: generalized unitarity, reverse unitarity (for phase-space integrals), method of regions, integration by parts canonical differential eqs applied KMOC to O= p_1^mu and p_rad^mu



momentum transfer (impulse)

(Hermann-Parra-Martinez-Ruf-Zeng'21)

$$\Delta p^{\mu,(2)}_{1,\perp,\text{cons}} = \frac{G^3 M^4 \nu}{|b|^3} \frac{2}{\sqrt{\sigma^2-1}} \frac{b^\mu}{|b|} \Bigg[h^2(\sigma,\nu) \left(16\sigma^2 - \frac{1}{(\sigma^2-1)^2} \right) - \frac{4}{3}\nu\,\sigma \left(14\sigma^2+25\right) - 8\nu \left(4\sigma^4 - 12\sigma^2 - 3\right) \frac{\text{arcsinh}\sqrt{\frac{\sigma-1}{2}}}{\sqrt{\sigma^2-1}} \Bigg] \quad (7$$

$$\Delta p^{\mu,(2)}_{1,u,\text{cons}} = \frac{G^3 M^5 \nu^2}{|b|^3} \frac{3\pi \left(2\sigma^2-1\right)\left(5\sigma^2-1\right)}{2\left(\sigma^2-1\right)} \left[\frac{1}{m_1} \check{u}^\mu_1 - \frac{1}{m_2} \check{u}^\mu_2 \right]$$

$$\Delta p^{\mu,(2)}_{1,\text{rad}} = \frac{G^3 M^4 \nu^2}{|b|^3} \Bigg\{ \frac{4}{\sqrt{\sigma^2-1}} \frac{b^\mu}{|b|} \left[f^{\text{LS}}_1(\sigma) + f^{\text{LS}}_3(\sigma) \frac{\sigma\,\text{arcsinh}\sqrt{\frac{\sigma-1}{2}}}{\sqrt{\sigma^2-1}} \right] + \pi\,\check{u}^\mu_2 \left[f_1(\sigma) + f_2(\sigma) \log\left(\frac{\sigma+1}{2}\right) + f_3(\sigma) \frac{\sigma\,\text{arcsinh}\sqrt{\frac{\sigma-1}{2}}}{\sqrt{\sigma^2-1}} \right] \Bigg\} .$$

$$f^{\text{LS}}_1(\sigma) = -\frac{(2\sigma^2-1)^2(5\sigma^2-8)}{3(\sigma^2-1)^{3/2}},$$

$$f^{\text{LS}}_3(\sigma) = \frac{2(2\sigma^2-1)^2(2\sigma^2-3)}{(\sigma^2-1)^{3/2}},$$

$$f_1(\sigma) = \frac{210\sigma^6 - 552\sigma^5 + 339\sigma^4 - 912\sigma^3 + 3148\sigma^2 - 3336\sigma + 1151}{48\left(\sigma^2-1\right)^{3/2}},$$

$$f_2(\sigma) = -\frac{35\sigma^4 + 60\sigma^3 - 150\sigma^2 + 76\sigma - 5}{8\sqrt{\sigma^2-1}},$$

$$f_3(\sigma) = \frac{\left(2\sigma^2-3\right)\left(35\sigma^4 - 30\sigma^2 + 11\right)}{8(\sigma^2-1)^{3/2}} .$$

radiated 4-momentum

$$\Delta R^\mu = \frac{G^3 m_1^2 m_2^2}{|b|^3} \frac{u^\mu_1 + u^\mu_2}{\sigma+1} \mathcal{E}(\sigma) + \mathcal{O}(G^4)$$

$$\frac{\mathcal{E}(\sigma)}{\pi} = f_1(\sigma) + f_2(\sigma) \log\left(\frac{\sigma+1}{2}\right) + f_3(\sigma) \frac{\sigma\,\text{arcsinh}\sqrt{\frac{\sigma-1}{2}}}{\sqrt{\sigma^2-1}}$$

High-energy puzzle

$$\mathcal{E}(\gamma) \propto \gamma^3$$

see also: DiVecchia et al, Riva-Vernizzi'21,Bjerrum-Bohr-Plante-Vanhove-Damgaard, useful results concerning the **waveform** (using QFT integration methods)…

potential gravitons only

Scattering Amplitudes and Conservative Binary Dynamics at $\mathcal{O}(G^4)$

Zvi Bern,[1] Julio Parra-Martinez,[2] Radu Roiban,[3] Michael S. Ruf,[4]
Chia-Hsien Shen,[5] Mikhail P. Solon,[1] and Mao Zeng[6]
[1]Mani L. Bhaumik Institute for Theoretical Physics, University of California at Los Angeles

three-loop level G^4

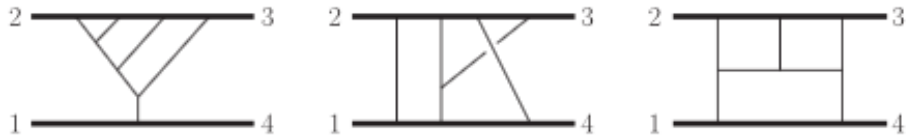


FIG. 2. Sample diagrams at $\mathcal{O}(G^4)$. From left to right: a contribution in the probe limit, a nonplanar diagram that contains iteration terms, and a diagram that contains contributions related to the tail effect.

IR-divergent because lacks tail effects

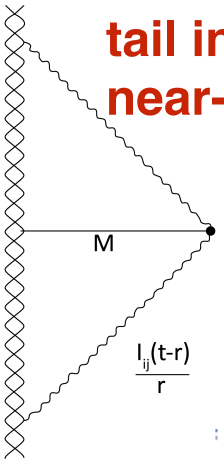
$$\mathcal{M}_4(\boldsymbol{q}) = G^4 M^7 \nu^2 |\boldsymbol{q}| \left(\frac{\boldsymbol{q}^2}{4^{1/3}\tilde{\mu}^2}\right)^{-3\epsilon} \pi^2 \left[\mathcal{M}_4^p + \nu\left(\frac{\mathcal{M}_4^t}{\epsilon} + \mathcal{M}_4^f\right)\right] + \int_\ell \frac{\tilde{I}_{r,1}^4}{Z_1 Z_2 Z_3} + \int_\ell \frac{\tilde{I}_{r,1}^2 \tilde{I}_{r,2}}{Z_1 Z_2} + \int_\ell \frac{\tilde{I}_{r,1}\tilde{I}_{r,3}}{Z_1} + \int_\ell \frac{\tilde{I}_{r,2}^2}{Z_1},$$

$$\mathcal{M}_4^p = -\frac{35(1-18\sigma^2+33\sigma^4)}{8(\sigma^2-1)}, \quad \mathcal{M}_4^t = h_1 + h_2 \log\left(\frac{\sigma+1}{2}\right) + h_3 \frac{\operatorname{arccosh}(\sigma)}{\sqrt{\sigma^2-1}},$$

$$\mathcal{M}_4^f = h_4 + h_5 \log\left(\frac{\sigma+1}{2}\right) + h_6 \frac{\operatorname{arccosh}(\sigma)}{\sqrt{\sigma^2-1}} + h_7 \log(\sigma) - h_2 \frac{2\pi^2}{3} + h_8 \frac{\operatorname{arccosh}^2(\sigma)}{\sigma^2-1} + h_9\left[\operatorname{Li}_2\left(\frac{1-\sigma}{2}\right) + \frac{1}{2}\log^2\left(\frac{\sigma+1}{2}\right)\right]$$
$$+ h_{10}\left[\operatorname{Li}_2\left(\frac{1-\sigma}{2}\right) - \frac{\pi^2}{6}\right] + h_{11}\left[\operatorname{Li}_2\left(\frac{1-\sigma}{1+\sigma}\right) - \operatorname{Li}_2\left(\frac{\sigma-1}{\sigma+1}\right) + \frac{\pi^2}{3}\right] + h_2 \frac{2\sigma(2\sigma^2-3)}{(\sigma^2-1)^{3/2}}\left[\operatorname{Li}_2\left(\sqrt{\frac{\sigma-1}{\sigma+1}}\right) - \operatorname{Li}_2\left(-\sqrt{\frac{\sigma-1}{\sigma+1}}\right)\right]$$
$$+ \frac{2h_3}{\sqrt{\sigma^2-1}}\left[\operatorname{Li}_2(1-\sigma-\sqrt{\sigma^2-1}) - \operatorname{Li}_2(1-\sigma+\sqrt{\sigma^2-1}) + 5\operatorname{Li}_2\left(\sqrt{\frac{\sigma-1}{\sigma+1}}\right) - 5\operatorname{Li}_2\left(-\sqrt{\frac{\sigma-1}{\sigma+1}}\right)\right.$$
$$\left. + 2\log\left(\frac{\sigma+1}{2}\right)\operatorname{arccosh}(\sigma)\right] + h_{12} K^2\left(\frac{\sigma-1}{\sigma+1}\right) + h_{13} K\left(\frac{\sigma-1}{\sigma+1}\right) E\left(\frac{\sigma-1}{\sigma+1}\right) + h_{14} E^2\left(\frac{\sigma-1}{\sigma+1}\right), \tag{6}$$

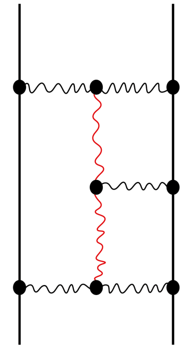
soft (radiationlike) gravitons

Scattering Amplitudes, the Tail Effect, and Conservative Binary Dynamics at $\mathcal{O}(G^4)$



tail in near-zone

Zvi Bern,[1] Julio Parra-Martinez,[2] Radu Roiban,[3] Michael S. Ruf,[1]
Chia-Hsien Shen,[4] Mikhail P. Solon,[1] and Mao Zeng[5]



Conservative Dynamics of Binary Systems at Fourth Post-Minkowskian Order in the Large-eccentricity Expansion

Christoph Dlapa,[1] Gregor Kälin,[1] Zhengwen Liu,[1] and Rafael A. Porto[1]

$$G^4 M^7 \nu^2 |\boldsymbol{q}| \pi^2 \left[\mathcal{M}_4^{\rm p} + \nu\left(4\mathcal{M}_4^{\rm t}\log\left(\tfrac{p_\infty}{2}\right) + \mathcal{M}_4^{\pi^2} + \mathcal{M}_4^{\rm rem}\right)\right] + \int_\ell \frac{\tilde{I}_{r,1}^4}{Z_1 Z_2 Z_3} + \int_\ell \frac{\tilde{I}_{r,1}^2 \tilde{I}_{r,2}}{Z_1 Z_2} + \int_\ell \frac{\tilde{I}_{r,1}\tilde{I}_{r,3}}{Z_1} + \int_\ell \frac{\tilde{I}_{r,2}^2}{Z_1},$$

$$\begin{aligned}
\mathcal{M}_4^{\rm p} &= -\frac{35\left(1-18\sigma^2+33\sigma^4\right)}{8\left(\sigma^2-1\right)},\\
\mathcal{M}_4^{\rm t} &= r_1 + r_2\log\left(\tfrac{\sigma+1}{2}\right) + r_3\frac{\operatorname{arccosh}(\sigma)}{\sqrt{\sigma^2-1}},\\
\mathcal{M}_4^{\pi^2} &= r_4\pi^2 + r_5\,{\rm K}\left(\tfrac{\sigma-1}{\sigma+1}\right){\rm E}\left(\tfrac{\sigma-1}{\sigma+1}\right) + r_6\,{\rm K}^2\left(\tfrac{\sigma-1}{\sigma+1}\right) + r_7\,{\rm E}^2\left(\tfrac{\sigma-1}{\sigma+1}\right),\\
\mathcal{M}_4^{\rm rem} &= r_8 + r_9\log\left(\tfrac{\sigma+1}{2}\right) + r_{10}\tfrac{\operatorname{arccosh}(\sigma)}{\sqrt{\sigma^2-1}} + r_{11}\log(\sigma) + r_{12}\log^2\left(\tfrac{\sigma+1}{2}\right) + r_{13}\tfrac{\operatorname{arccosh}(\sigma)}{\sqrt{\sigma^2-1}}\log\left(\tfrac{\sigma+1}{2}\right) + r_{14}\tfrac{\operatorname{arccosh}^2(\sigma)}{\sigma^2-1}\\
&\quad + r_{15}\,{\rm Li}_2\left(\tfrac{1-\sigma}{2}\right) + r_{16}\,{\rm Li}_2\left(\tfrac{1-\sigma}{1+\sigma}\right) + r_{17}\tfrac{1}{\sqrt{\sigma^2-1}}\left[{\rm Li}_2\left(-\sqrt{\tfrac{\sigma-1}{\sigma+1}}\right) - {\rm Li}_2\left(\sqrt{\tfrac{\sigma-1}{\sigma+1}}\right)\right].
\end{aligned} \tag{3}$$

$$\begin{aligned}
\mathcal{M}_4^{\rm radgrav,f} = {} & \frac{12044}{75}p_\infty^2 + \frac{212077}{3675}p_\infty^4 + \frac{115917979}{793800}p_\infty^6\\
& - \frac{9823091209}{76839840}p_\infty^8 + \frac{115240251793703}{1038874636800}p_\infty^{10}\\
& - \frac{411188753665637}{4155498547200}p_\infty^{12} + \cdots,
\end{aligned} \tag{6}$$

whose first three terms match the sixth PN order result in Eq. (6.20) of Ref. [42].

matches the 6PN result of the Tutti Frutti approach (Bini-D-Geralico'21)

Tutti-Frutti method

(**Bini-TD-Geralico** '19,'20'21) **combines PN MPM EOB Delaunay Self-Force Scattering properties**

SIXTH POST-NEWTONIAN LOCAL-IN-TIME DYNAMIC

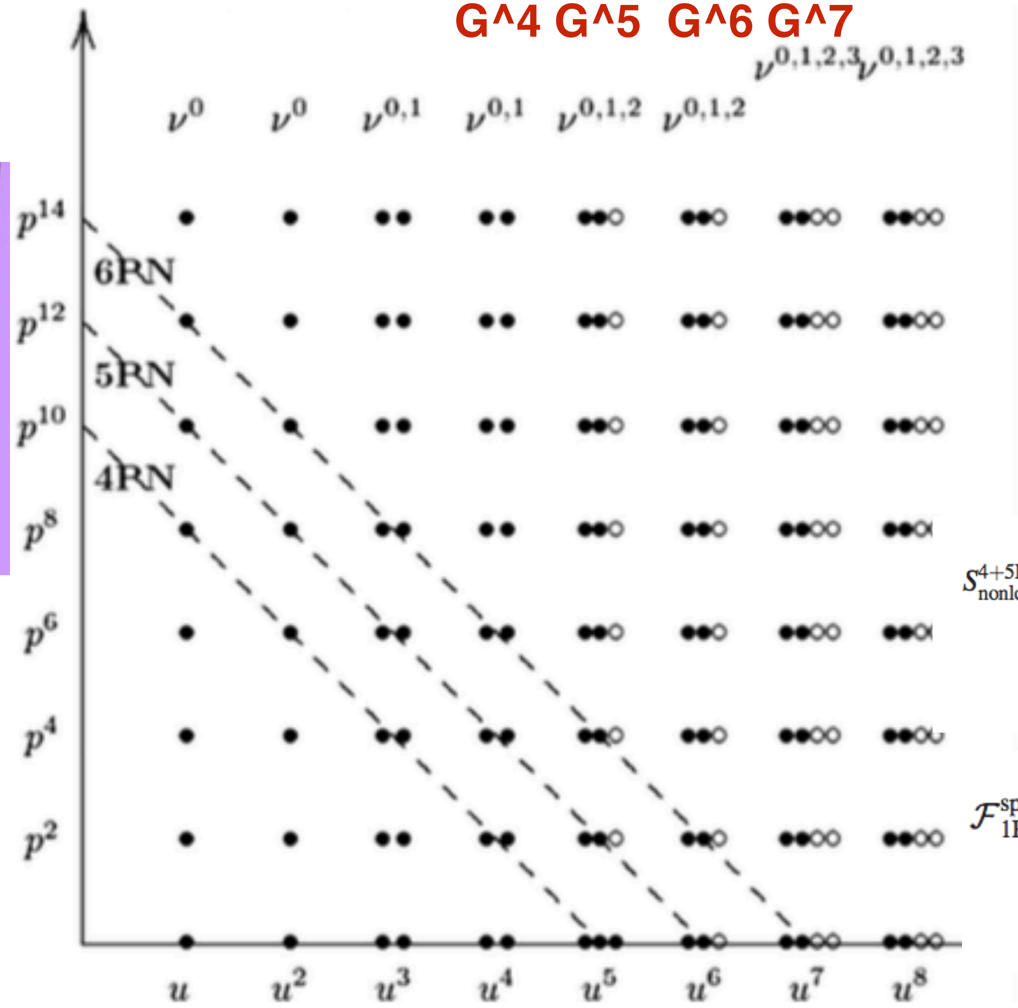


FIG. 1. Schematic representation of the irreducible information contained, at each post-Minkowskian level (keyed by a power of $u = GM/r$), in the local dynamics. Each vertical column of dots describes the post-Newtonian expansion (keyed by powers of p^2) of an energy-dependent function parametrizing the scattering angle. The various columns at a given post-Minkowskian level correspond to increasing powers of the symmetric mass-ratio ν. See text for details.

6PN conservative dynamics complete at 3PM and 4PM

$$S^{4+5\mathrm{PN}}_{\mathrm{nonloc}}[x_1(s_1),x_2(s_2)] = \frac{G^2\mathcal{M}}{c^3}\int dt\,\mathrm{PF}_{2r^h_{12}(t)/c} \times \int \frac{dt'}{|t-t'|}\mathcal{F}^{\mathrm{split}}_{1\mathrm{PN}}(t,t').$$

$$\mathcal{F}^{\mathrm{split}}_{1\mathrm{PN}}(t,t') = \frac{G}{c^5}\left(\frac{1}{5}I^{(3)}_{ab}(t)I^{(3)}_{ab}(t') + \frac{1}{189c^2} \right. \cdots$$
$$\left. + \frac{16}{45c^2}J^{(3)}_{ab}(t)J^{(3)}_{ab}(t')\right).$$

Classical scattering worldline perturbation theory enhanced by using QFT integration methods

$$\frac{dx_a^\mu}{d\sigma_a} = g^{\mu\nu}(x_a)p_{a\nu},$$

$$\frac{dp_{a\mu}}{d\sigma_a} = -\frac{1}{2}\partial_\mu g^{\alpha\beta}(x_a)p_{a\alpha}p_{a\beta}.$$

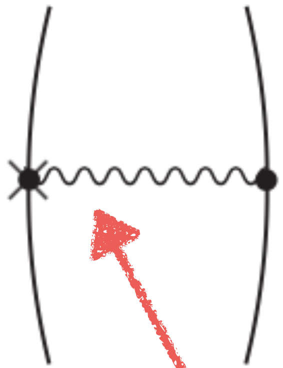
$$R^{\mu\nu} - \tfrac{1}{2}Rg^{\mu\nu} = 8\pi G T^{\mu\nu}$$

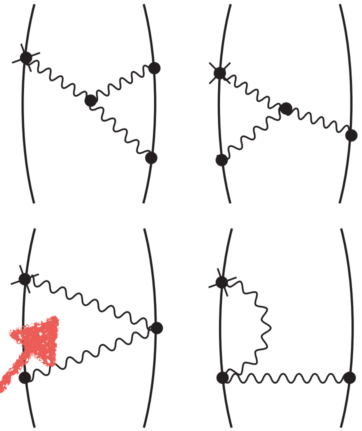
$$T^{\mu\nu}(x) = \sum \int d\sigma_a p_a^\mu p_a^\nu \frac{\delta^4(x - x_a(\sigma_a))}{\sqrt{g}}$$

$$\Delta p_{a\mu} = \int_{-\infty}^{+\infty} d\sigma_a \frac{dp_{a\mu}}{d\sigma_a}$$

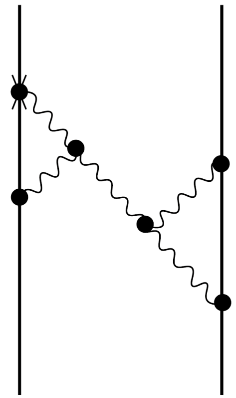
$$= -\frac{1}{2}\int_{-\infty}^{+\infty} d\sigma_a \partial_\mu g^{\alpha\beta}(x_a)p_{a\alpha}p_{a\beta}.$$

$$\Delta p_{1\mu} = 2G\int d\sigma_1 d\sigma_2 p_{1\alpha}p_{1\beta} \times \partial_\mu \mathcal{P}^{\alpha\beta;\alpha'\beta'}(x_1(\sigma_1) - x_2(\sigma_2))p_{2\alpha'}p_{2\beta'}$$





with retarded propagator



$$\mathcal{P}^{\alpha\beta;\alpha'\beta'}(x-y) = \left(\eta^{\alpha\alpha'}\eta^{\beta\beta'} - \frac{1}{2}\eta^{\alpha\beta}\eta^{\alpha'\beta'}\right)\mathcal{G}(x-y)$$

with retarded propagator

Approach initiated long ago: Rosenblum'78 Westpfahl'79,'85 Portilla'80 Bel et al.'81 limited by the technical difficulty of computing the integrals beyond G^2, ie at G^2=2-loop.

Recently developed to compete with quantum-scattering approach: Kalin-Porto, Porto et al, Plefka et al, Dlapa-Kalin-Liu-Porto,…

Radiation Reaction and Gravitational Waves at Fourth Post-Minkowskian Order

Christoph Dlapa,[1] Gregor Kälin,[1] Zhengwen Liu,[2,1] Jakob Neef,[3,4] and Rafael A. Porto[1] (PRL 10 March 2023)

confirmed by recent QFT computation
Damgaard-Hansen-Plant'e-Vanhove'23

G^4 term

$$\Delta^{(n)} p_1^\mu = c_{1b}^{(n)} \frac{\hat{b}^\mu}{b^n} + \frac{1}{b^n} \sum_a c_{1\check{u}_a}^{(n)} \check{u}_a^\mu,$$

$$\begin{aligned}\frac{c_{1b}^{(4)\,\mathrm{tot}}}{\pi} = & -\frac{3h_1 m_1 m_2 (m_1^3+m_2^3)}{64(\gamma^2-1)^{5/2}} + m_1^2 m_2^2 (m_1+m_2)\left[\frac{21 h_2 \mathrm{E}^2\left(\frac{\gamma-1}{\gamma+1}\right)}{32(\gamma-1)\sqrt{\gamma^2-1}} + \frac{3h_3 \mathrm{K}^2\left(\frac{\gamma-1}{\gamma+1}\right)}{16(\gamma^2-1)^{3/2}} - \frac{3h_4 \mathrm{E}\left(\frac{\gamma-1}{\gamma+1}\right)\mathrm{K}\left(\frac{\gamma-1}{\gamma+1}\right)}{16(\gamma^2-1)^{3/2}} + \frac{\pi^2 h_5}{8\sqrt{\gamma^2-1}}\right.\\
& \left. + \frac{h_6 \log\left(\frac{\gamma-1}{2}\right)}{16(\gamma^2-1)^{3/2}} + \frac{3h_7 \mathrm{Li}_2\left(\sqrt{\frac{\gamma-1}{\gamma+1}}\right)}{(\gamma-1)(\gamma+1)^2} - \frac{3h_7 \mathrm{Li}_2\left(\frac{\gamma-1}{\gamma+1}\right)}{4(\gamma-1)(\gamma+1)^2}\right] + m_1^3 m_2^2 \left[\frac{h_8}{48(\gamma^2-1)^3} + \frac{\sqrt{\gamma^2-1}\, h_9}{768(\gamma-1)^3\gamma^9(\gamma+1)^4} + \frac{h_{10}\log\left(\frac{\gamma+1}{2}\right)}{8(\gamma^2-1)^2}\right.\\
& - \frac{h_{11}\log\left(\frac{\gamma+1}{2}\right)}{32(\gamma^2-1)^{5/2}} + \frac{h_{12}\log(\gamma)}{16(\gamma^2-1)^{5/2}} - \frac{h_{13}\mathrm{arccosh}(\gamma)}{8(\gamma-1)(\gamma+1)^4} + \frac{h_{14}\mathrm{arccosh}(\gamma)}{16(\gamma^2-1)^{7/2}} - \frac{3h_{15}\log\left(\frac{\gamma+1}{2}\right)\log\left(\frac{\gamma-1}{\gamma+1}\right)}{8\sqrt{\gamma^2-1}} + \frac{3h_{16}\mathrm{arccosh}(\gamma)\log\left(\frac{\gamma-1}{\gamma+1}\right)}{16(\gamma^2-1)^2}\\
& \left. - \frac{3h_{17}\mathrm{Li}_2\left(\frac{\gamma-1}{\gamma+1}\right)}{64\sqrt{\gamma^2-1}} - \frac{3}{32}\sqrt{\gamma^2-1}\, h_{18}\mathrm{Li}_2\left(\frac{1-\gamma}{\gamma+1}\right)\right] + m_1^2 m_2^3 \left[\frac{3h_{15}\log\left(\frac{2}{\gamma-1}\right)\log\left(\frac{\gamma+1}{2}\right)}{8\sqrt{\gamma^2-1}} + \frac{3h_{16}\log\left(\frac{\gamma-1}{2}\right)\mathrm{arccosh}(\gamma)}{16(\gamma^2-1)^2} + \frac{h_{19}}{48(\gamma^2-1)^3}\right.\\
& + \frac{h_{20}}{192\gamma^7(\gamma^2-1)^{5/2}} + \frac{h_{21}\log\left(\frac{\gamma+1}{2}\right)}{8(\gamma^2-1)^2} + \frac{h_{22}\log\left(\frac{\gamma+1}{2}\right)}{16(\gamma^2-1)^{3/2}} + \frac{h_{23}\log(\gamma)}{2(\gamma^2-1)^{3/2}} - \frac{h_{24}\mathrm{arccosh}(\gamma)}{16(\gamma^2-1)^3} + \frac{h_{25}\mathrm{arccosh}(\gamma)}{16(\gamma^2-1)^{7/2}} - \frac{3h_{26}\mathrm{arccosh}^2(\gamma)}{32(\gamma^2-1)^{7/2}}\\
& \left. + \frac{3h_{27}\log^2\left(\frac{\gamma+1}{2}\right)}{2\sqrt{\gamma^2-1}} + \frac{3h_{28}\log\left(\frac{\gamma+1}{2}\right)\mathrm{arccosh}(\gamma)}{16(\gamma^2-1)^2} + \frac{h_{29}\mathrm{Li}_2\left(\frac{1-\gamma}{\gamma+1}\right)}{4\sqrt{\gamma^2-1}} + \frac{3h_{30}\mathrm{Li}_2\left(\frac{\gamma-1}{\gamma+1}\right)}{8\sqrt{\gamma^2-1}}\right],\end{aligned}$$

Its PN expansion agrees with Bini-TD-Geralico'23 notably for the **nu^2 =O(RR^2) contribution.**
Moreover, Bini-TD-Geralico'23 went beyond the linear-response formula by using balance+mass-polynomiality

$$\Delta p_{a\mu} = \Delta p_{a\mu}^{\mathrm{cons}} + \Delta p_{a\mu}^{\mathrm{rr\,lin}} + \Delta p_{a\mu}^{\mathrm{rr\,nonlin}}$$

$$\Delta p^{\mathrm{rr}}_{1\,\mu G^4} = \Delta p^{\mathrm{rr\,lin-odd}}_{1\,\mu G^4} + \frac{G^4}{b^4} m_1^3 m_2^2 p_x^{G^4}(\gamma)\hat{b}^\mu_{12},$$

$$\Delta p^{\mathrm{rr}}_{1\,\mu G^4} = \Delta p^{\mathrm{rr\,lin-odd}}_{1\,\mu G^4} + \frac{m_1}{m_2-m_1} P^{\mathrm{rad}}_{x\,G^4}\hat{b}^\mu_{12}.$$

relation between the rad-reac^2 term and P^rad_x

Strong-field scattering of two black holes: Numerical relativity meets post-Minkowskian gravity

Thibault Damour[1] and Piero Rettegno[2,3]

(PRD March 2023)

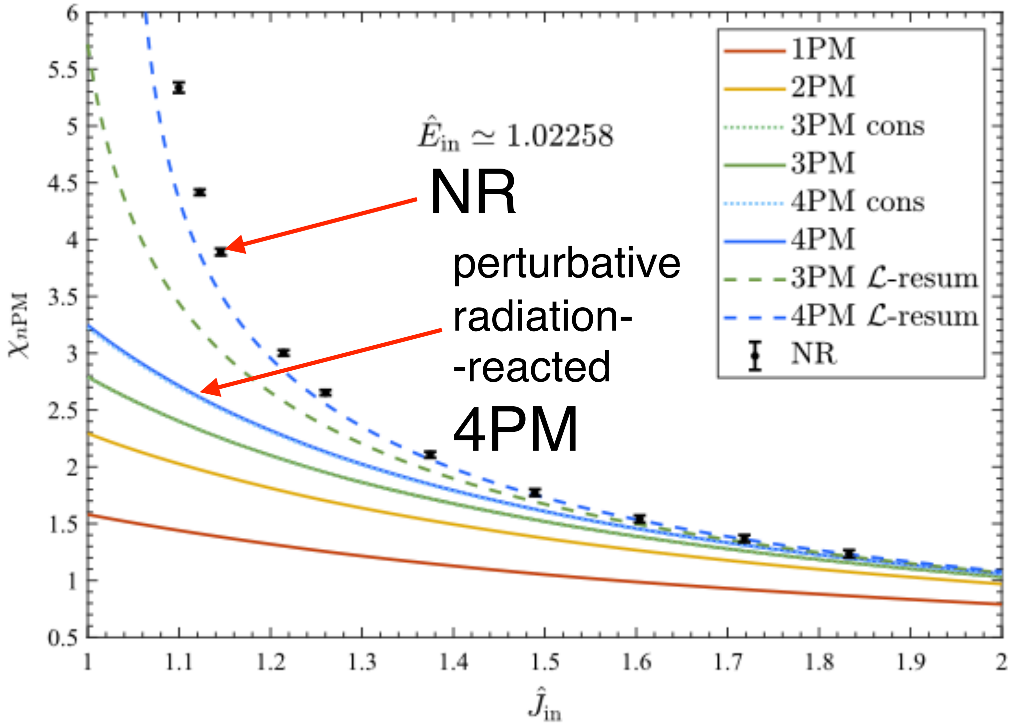
$$\chi_{n\mathrm{PM}}(\gamma, j) \equiv \sum_{i=1}^{n} 2\frac{\chi_i(\gamma)}{j^i} \quad \longrightarrow \quad \mu^2 + g_{\mathrm{eff}}^{\mu\nu} P_\mu P_\nu + Q(X^\mu, P_\mu) = 0,$$

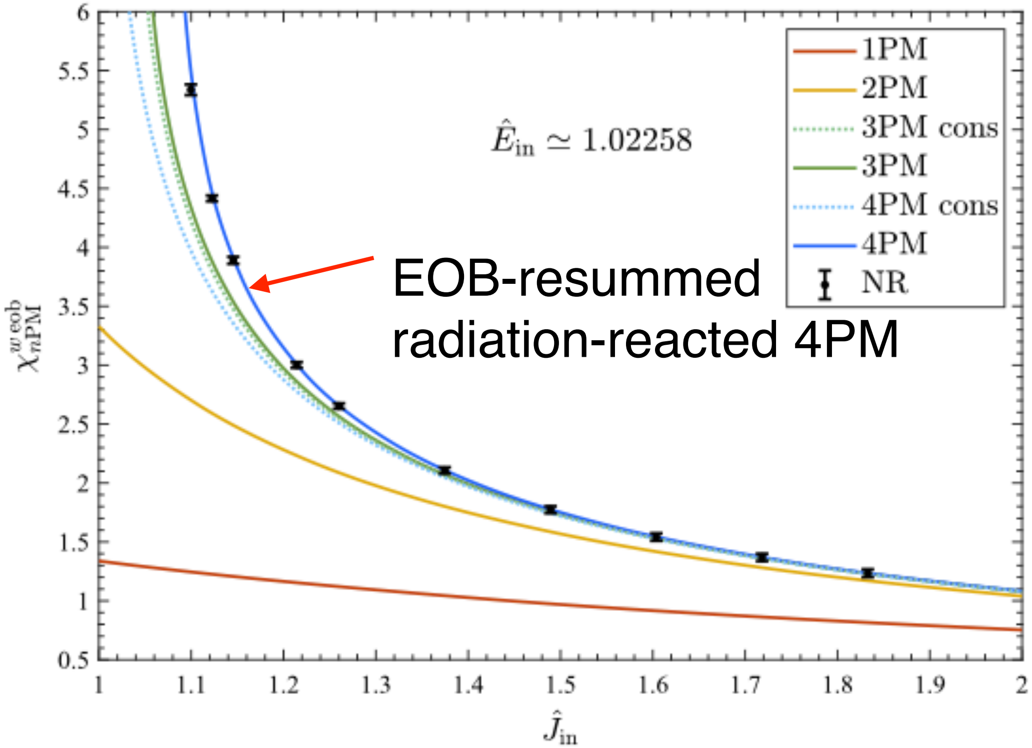
$$p_{\bar r}^2 + \frac{j^2}{\bar r^2} = p_\infty^2 + w(\bar r, \gamma).$$

$$w(\bar r, \gamma) = \frac{w_1(\gamma)}{\bar r} + \frac{w_2(\gamma)}{\bar r^2} + \frac{w_3(\gamma)}{\bar r^3} + \frac{w_4(\gamma)}{\bar r^4} + O\left[\frac{1}{\bar r^5}\right]$$

Newtonianlike EOB radial potential

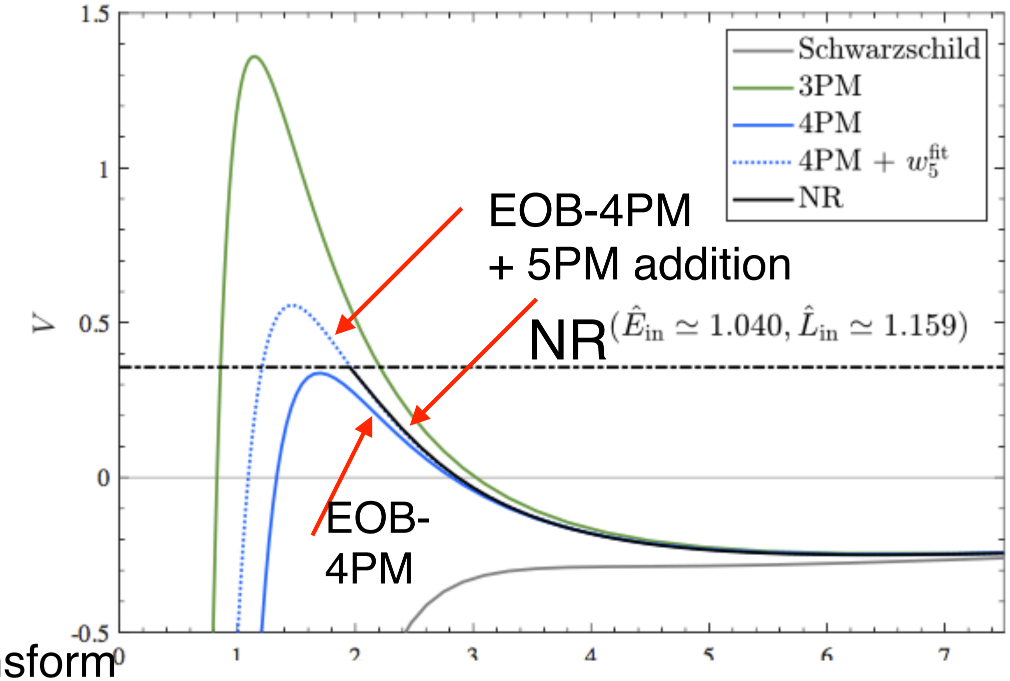
$$\chi_{n\mathrm{PM}}^{w\,\mathrm{eob}}(\gamma, j) \equiv 2j \int_0^{\bar u_{\max}(\gamma, j)} \frac{d\bar u}{\sqrt{p_\infty^2 + w_{n\mathrm{PM}}(\bar u, \gamma) - j^2 \bar u^2}} - \pi.$$





Strong-field scattering of two spinning black holes: Numerical Relativity versus post-Minkowskian gravity (Rettegno et al.'23)

Higher-energy non-spinning: Comparison between the effective potential V=L^2/r^2-w(r) extracted from NR simulations and its EOB-PM equivalent

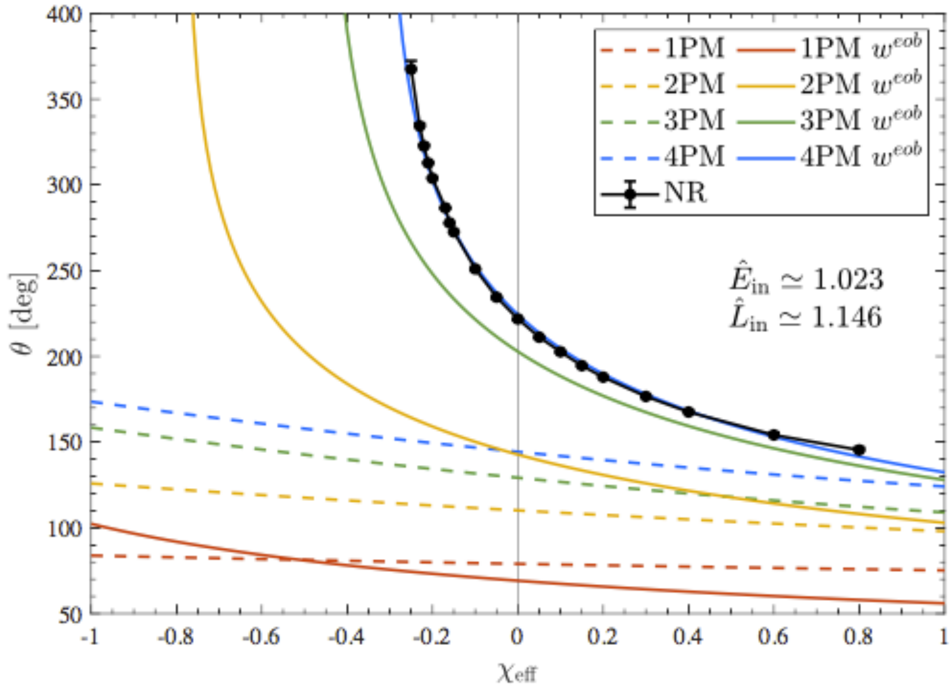


In the spin-aligned case, one can transform the PM-expanded scattering angle

$$\theta_{n\rm PM}(\gamma,\ell,S_i) \equiv \sum_{k=1}^{n} 2\frac{\theta_k(\gamma,\ell,S_i)}{\ell^k}$$

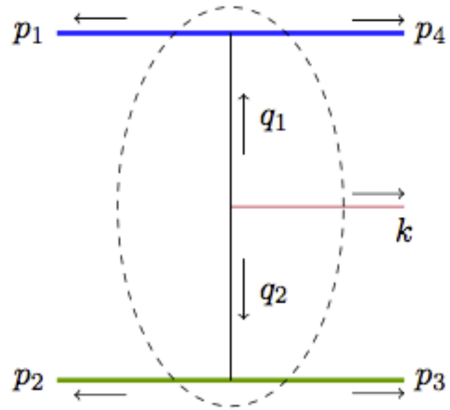
into an equivalent spin-dependent EOB potential

$$\begin{aligned} w_{n\rm PM}(\bar{r},\gamma,\ell,S_i) &= w^{\rm orb}(\bar{r},\gamma) \\ &+ \frac{\ell\, w^{\rm S}_{n\rm PM}(\bar{r},\gamma)}{\bar{r}^2} + \frac{w^{\rm S^2}_{n\rm PM}(\bar{r},\gamma)}{\bar{r}^2} \\ &+ \frac{\ell\, w^{\rm S^3}_{n\rm PM}(\bar{r},\gamma)}{\bar{r}^4} + \frac{w^{\rm S^4}_{n\rm PM}(\bar{r},\gamma)}{\bar{r}^4}. \end{aligned}$$



PM waveform computation $W(k^\mu) = \epsilon^\mu \epsilon^\nu h_{\mu\nu}(\omega, \theta, \phi)$

G^1=1PM (linearized,Einstein 1918) stationary $\propto \delta(\omega)$



LO (tree level) waveform

G^2=2PM: classical time-domain W(t,n): Kovacs-Thorne 1977

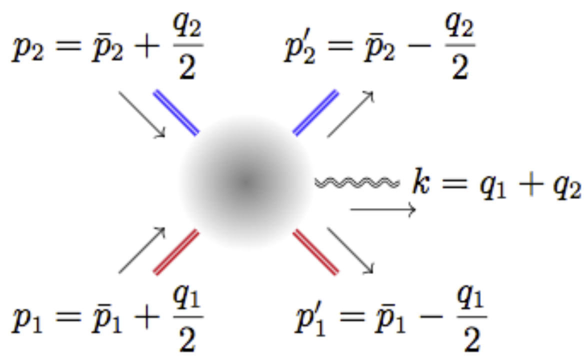
quantum-based: yields W(k,p1,p2,p3,p4)=W(k,p1,p2,q1)

Johansson-Ochirov'15, GoldbergerRidgway'17

Luna-Nicholson-OConnellWhite'18

Mougiakakos-Riva-Vernizzi'21,Bautista-Siemonsen'22

Recent NLO (one-loop) waveform



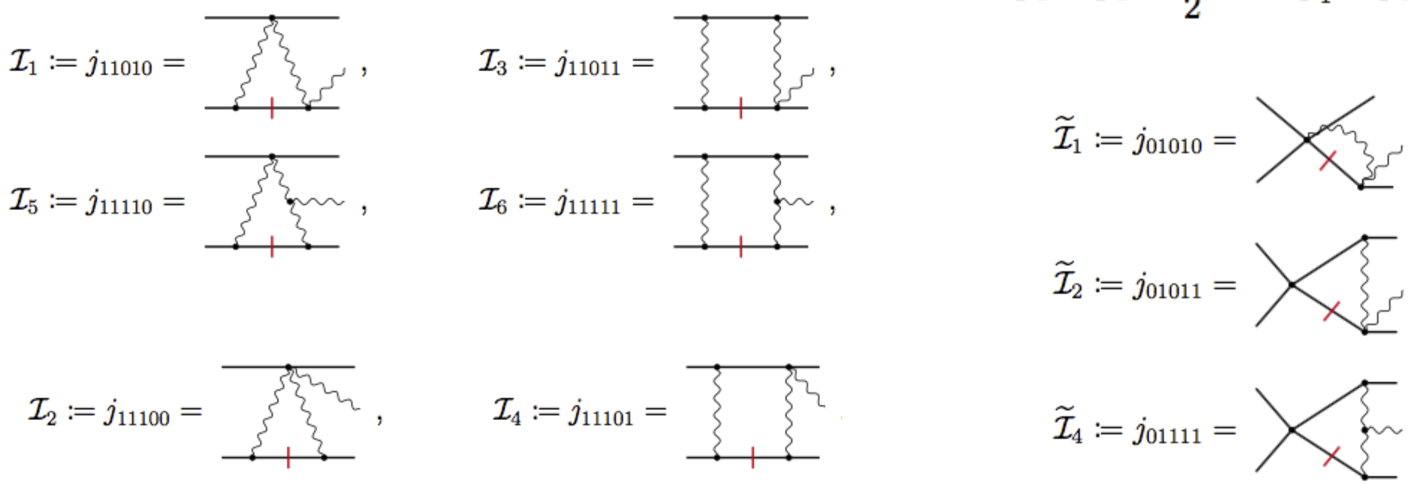
G^3=3PM

Andreas Brandhuber[a], Graham R. Brown[a], Gang Chen[b], Stefano De Angelis[c], Joshua Gowdy[a] and Gabriele Tr

Aidan Herderschee,[a] Radu Roiban[b,c] and Fei Teng[b,c]

Alessandro Georgoudis^a Carlo Heissenberg^b,a Ingrid Vazquez-Holm^b,a

5-point HEFT amplitude at one-loop O(G^3) waveform



$$\mathcal{M} = \kappa\, m_1 \mathcal{M}_1^{\rm lin} + \kappa^3 m_1 m_2 \mathcal{M}_1^{\rm tree} + \kappa^5 m_1^2 m_2 \mathcal{M}_1^{\rm one\,loop} + (1 \to 2)$$
$$= -i\frac{\kappa}{2}\epsilon^*_\mu \epsilon^*_\nu \tilde{T}^{\mu\nu}(k) = -i\frac{\kappa}{2}\epsilon^*_\mu \epsilon^*_\nu \int \mu_{12}(q_1, q_2)\tilde{T}^{\mu\nu}(k, q_1, q_2)$$

$$\mu_{1,2}(k) \equiv e^{i(q_1\cdot b_1 + q_2\cdot b_2)}\bar{\delta}^{(4)}(k - q_1 - q_2)\bar{\delta}(q_1\cdot u_1)\bar{\delta}(q_2\cdot u_2)$$

$\mathcal{M}_1^{\rm lin}$

$$\tilde{T}^{\mu\nu}_{\rm Fig.\,1a}(k) = \sum_a m_a u_a^\mu u_a^\nu e^{ik\cdot b_a}\bar{\delta}(\omega_a)$$

$\mathcal{M}_1^{\rm tree}$

to be integrated

$$\tilde{T}^{\mu\nu}_{\rm Fig.\,1b}(k) = \frac{m_1 m_2}{4m_{\rm Pl}^2}\int_{q_1,q_2}\mu_{1,2}(k)\frac{1}{q_2^2}\left[\frac{2\gamma^2-1}{\omega_1 + i\epsilon}q_2^{(\mu}u_1^{\nu)} - 4\gamma u_2^{(\mu}u_1^{\nu)} - \left(\frac{2\gamma^2-1}{2}\frac{k\cdot q_2}{(\omega_1+i\epsilon)^2} - \frac{2\gamma\omega_2}{\omega_1+i\epsilon} - 1\right)u_1^\mu u_1^\nu\right],$$

$$\tilde{T}^{\mu\nu}_{\rm Fig.\,1c}(k) = \frac{m_1 m_2}{4m_{\rm Pl}^2}\int_{q_1,q_2}\mu_{1,2}(k)\frac{1}{q_1^2 q_2^2}\left[\frac{2\gamma^2-1}{2}q_2^\mu q_2^\nu + (2\omega_2^2 - q_1^2)u_1^\mu u_1^\nu + 4\gamma\omega_2 q_2^{(\mu}u_1^{\nu)}\right.$$
$$\left. - \eta^{\mu\nu}\left(\gamma\omega_1\omega_2 + \frac{2\gamma^2-1}{4}q_2^2\right) + 2(\gamma q_1^2 - \omega_1\omega_2)u_1^{(\mu}u_2^{\nu)}\right],$$

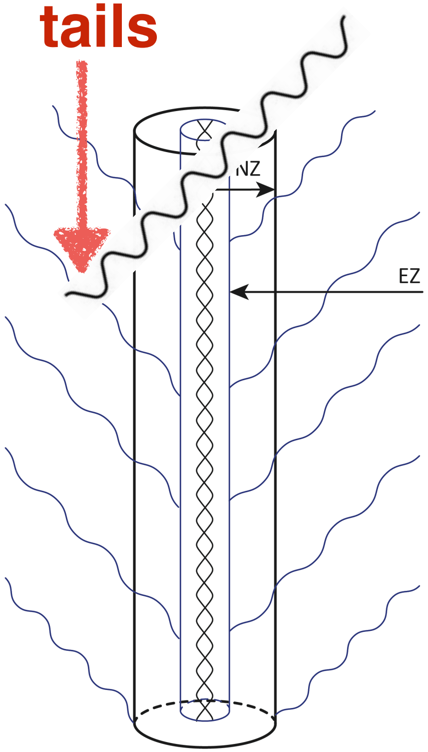
$\mathcal{M}_1^{\rm one\,loop}$

complicated rational functions of 8 variables with **spurious poles**

$$\mathcal{M}^{(1)}_{\bar{m}_1^3\bar{m}_2^2} = \frac{\mathfrak{d}_{\rm IR}}{\epsilon} + \mathfrak{R} + i\pi\,\mathfrak{i}_1 + \frac{i\pi}{\sqrt{\bar{y}^2-1}}\mathfrak{i}_2 + c_{1,0}\,\mathcal{I}_1 + c_{2,0}\,\mathcal{I}_2$$
$$+ \mathfrak{l}_{\bar{w}_1}\log\frac{\bar{w}_1^2}{\mu_{\rm IR}^2} + \mathfrak{l}_{\bar{w}_2}\log\frac{\bar{w}_2^2}{\mu_{\rm IR}^2} + \mathfrak{l}_q \log\frac{q_1^2}{q_2^2} + \mathfrak{l}_{\bar{y}}\frac{\log\left(\sqrt{\bar{y}^2-1}+\bar{y}\right)}{\sqrt{\bar{y}^2-1}} + \mathcal{O}(\epsilon^0)$$

Comparing one-loop amplitude to MPM waveform

tails



algorithmic STF tensors encoding multipole moments (related to the source moments I_L,J_L)

$$g = \eta + Gh_1 + G^2 h_2 + G^3 h_3 + \ldots,$$
$$\Box h_1 = 0,$$
$$\Box h_2 = \partial\partial h_1 h_1,$$
$$\Box h_3 = \partial\partial h_1 h_1 h_1 + \partial\partial h_1 h_2,$$
$$h_1 = \sum_\ell \partial_{i_1 i_2 \ldots i_\ell}\left(\frac{M_{i_1 i_2 \ldots i_\ell}(t-r/c)}{r}\right) + \partial\partial\ldots\partial\left(\frac{\epsilon_{j_1 j_2 k} S_{k j_3 \ldots j_\ell}(t-r/c)}{r}\right),$$
$$h_2 = FP_B \Box^{-1}_{\rm ret}\left(\left(\frac{r}{r_0}\right)^B \partial\partial h_1 h_1\right) + \ldots,$$
$$h_3 = FP_B \Box^{-1}_{\rm ret} \ldots$$

radiative multipole moments (observable at infinity) U_L, V_L

$$r h^{\rm TT}_{ij} = \frac{4G}{c^2} P(n)_{ijab} \sum_{l=2}^{\infty} \frac{1}{c^l}\frac{1}{l!}\left(U_{abL-2} n_{L-2} - \frac{2l}{c(l+1)} n_{cL-2}\epsilon_{cd(a} V_{b)dL-2}\right)$$

$$\mathcal{M}^{\rm MPM}(k,b,u_1,u_2,m_1,m_2) = -i\frac{\kappa}{2}\epsilon^\mu\epsilon^\nu h^{\rm MPM}_{\mu\nu}(\omega,\theta,\phi) = -i\frac{\kappa}{2}\int dt e^{i\omega t}\epsilon^\mu\epsilon^\nu h^{\rm MPM}_{\mu\nu}(t,\theta,\phi)$$

$$\mathcal{M}^{\rm HEFT}(k,b,u_1,u_2,m_1,m_2) =$$
$$e^{i\frac{b_1+b_2}{2}\cdot k}\int \frac{d^D q}{(2\pi)^{D-2}}\,\delta\left(2p_1\cdot\left(q+\frac{k}{2}\right)\right)\delta\left(2p_2\cdot\left(-q+\frac{k}{2}\right)\right)\, e^{iq\cdot(b_1-b_2)}\mathcal{M}^{(1)}_{5,\rm HEFT}\left(q+\frac{k}{2},-q+\frac{k}{2};h\right)$$

MPM radiative quadrupole moment

radiative quadrupole

$$\mathrm{U}_{ij} = \mathrm{U}_{ij}^{\text{inst}} + \mathrm{U}_{ij}^{\text{tail}} + \mathrm{U}_{ij}^{\text{tail-tail}} + \mathrm{U}_{ij}^{\text{mem}} + \mathcal{O}\left(\frac{1}{c^8}\right)$$

$$\begin{aligned}
\mathrm{U}_{ij}^{\text{inst}} &= \mathrm{M}_{ij}^{(2)} \\
&+ \frac{G}{c^5}\left[\frac{1}{7}\mathrm{M}^{(5)}_{a\langle i}\mathrm{M}_{j\rangle a} - \frac{5}{7}\mathrm{M}^{(4)}_{a\langle i}\mathrm{M}^{(1)}_{j\rangle a} - \frac{2}{7}\mathrm{M}^{(3)}_{a\langle i}\mathrm{M}^{(2)}_{j\rangle a} + \frac{1}{3}\epsilon_{ab\langle i}\mathrm{M}^{(4)}_{j\rangle a}\mathrm{S}_b\right] \\
&+ \frac{G}{c^7}\Bigg[-\frac{64}{63}\mathrm{S}^{(2)}_{a\langle i}\mathrm{S}^{(3)}_{j\rangle a} + \frac{1957}{3024}\mathrm{M}^{(3)}_{ijab}\mathrm{M}^{(4)}_{ab} + \frac{5}{2268}\mathrm{M}^{(3)}_{ab\langle i}\mathrm{M}^{(4)}_{j\rangle ab} + \frac{19}{648}\mathrm{M}^{(3)}_{ab}\mathrm{M}^{(4)}_{ijab} \\
&\quad + \frac{16}{63}\mathrm{S}^{(1)}_{a\langle i}\mathrm{S}^{(4)}_{j\rangle a} + \frac{1685}{1008}\mathrm{M}^{(2)}_{ijab}\mathrm{M}^{(5)}_{ab} + \frac{5}{126}\mathrm{M}^{(2)}_{ab\langle i}\mathrm{M}^{(5)}_{j\rangle ab} - \frac{5}{756}\mathrm{M}^{(2)}_{ab}\mathrm{M}^{(5)}_{ijab} \\
&\quad + \frac{80}{63}\mathrm{S}_{a\langle i}\mathrm{S}^{(5)}_{j\rangle a} + \frac{5}{42}\mathrm{S}_{a}\mathrm{S}^{(5)}_{ija} + \frac{41}{28}\mathrm{M}^{(1)}_{ijab}\mathrm{M}^{(6)}_{ab} + \frac{5}{189}\mathrm{M}^{(1)}_{ab\langle i}\mathrm{M}^{(6)}_{j\rangle ab} \\
&\quad + \frac{1}{432}\mathrm{M}^{(1)}_{ab}\mathrm{M}^{(6)}_{ijab} + \frac{91}{216}\mathrm{M}_{ijab}\mathrm{M}^{(7)}_{ab} - \frac{5}{252}\mathrm{M}_{ab\langle i}\mathrm{M}^{(7)}_{j\rangle ab} - \frac{1}{432}\mathrm{M}_{ab}\mathrm{M}^{(7)}_{ijab} \\
&\quad + \epsilon_{ac\langle i}\Bigg(\frac{32}{189}\mathrm{M}^{(3)}_{j\rangle bc}\mathrm{S}^{(3)}_{ab} - \frac{1}{6}\mathrm{M}^{(3)}_{ab}\mathrm{S}^{(3)}_{j\rangle bc} + \frac{3}{56}\mathrm{S}^{(2)}_{j\rangle bc}\mathrm{M}^{(4)}_{ab} + \frac{10}{189}\mathrm{S}^{(2)}_{ab}\mathrm{M}^{(4)}_{j\rangle bc} \\
&\quad + \frac{65}{189}\mathrm{M}^{(2)}_{j\rangle bc}\mathrm{S}^{(4)}_{ab} + \frac{1}{28}\mathrm{M}^{(2)}_{ab}\mathrm{S}^{(4)}_{j\rangle bc} + \frac{187}{168}\mathrm{S}^{(1)}_{j\rangle bc}\mathrm{M}^{(5)}_{ab} - \frac{1}{189}\mathrm{S}^{(1)}_{ab}\mathrm{M}^{(5)}_{j\rangle bc} \\
&\quad - \frac{5}{189}\mathrm{M}^{(1)}_{j\rangle bc}\mathrm{S}^{(5)}_{ab} + \frac{1}{24}\mathrm{M}^{(1)}_{ab}\mathrm{S}^{(5)}_{j\rangle bc} + \frac{65}{84}\mathrm{S}_{j\rangle bc}\mathrm{M}^{(6)}_{ab} + \frac{1}{189}\mathrm{S}_{ab}\mathrm{M}^{(6)}_{j\rangle bc} \\
&\quad - \frac{10}{63}\mathrm{M}_{j\rangle bc}\mathrm{S}^{(6)}_{ab} + \frac{1}{168}\mathrm{M}_{ab}\mathrm{S}^{(6)}_{j\rangle bc}\Bigg)\Bigg].
\end{aligned}$$

linear tails (inc. harmonic #)

$$\mathrm{U}_{ij}^{\text{tail}} = \frac{2GM}{c^3}\int_0^{+\infty} d\tau \left[\ln\left(\frac{c\tau}{2r_0}\right) + \frac{11}{12}\right]\mathrm{M}^{(4)}_{ij}(U-\tau)\,.$$

$$\mathrm{U}_{ij}^{\text{tail-tail}} = 2\left(\frac{GM}{c^3}\right)^2\int_0^{+\infty} d\tau \left[\ln^2\left(\frac{c\tau}{2r_0}\right) + \frac{57}{70}\ln\left(\frac{c\tau}{2r_0}\right) + \frac{124627}{44100}\right]$$

$$\mathrm{U}_{ij}^{\text{mem}} = \frac{G}{c^5}\left[-\frac{2}{7}\int_0^{+\infty} d\tau\, \mathrm{M}^{(3)}_{a\langle i}\mathrm{M}^{(3)}_{j\rangle a}(U-\tau)\right] + \ldots$$

$$\mathrm{M}_{ij} = \mathrm{I}_{ij} + \frac{4G}{c^5}\left[\mathrm{W}^{(2)}\mathrm{I}_{ij} - \mathrm{W}^{(1)}\mathrm{I}^{(1)}_{ij}\right] + \ldots$$

direct radiation-reaction term

$$I_{ij} = \nu m\left\{\left[A - \frac{24}{7}\frac{\nu}{c^5}\frac{G^2m^2}{r^2}\dot{r}\right]x_{\langle i}x_{j\rangle} + B\frac{r^2}{c^2}v_{\langle i}v_{j\rangle} + 2\left[C\frac{r\dot{r}}{c^2} + \frac{24}{7}\frac{\nu}{c^5}\frac{G^2m^2}{r}\right]x_{\langle i}v_{j\rangle}\right\},$$

source quadrupole

(to be evaluated along **rad-reacted** hyperbolic motion and Fourier-transformed to omega)

time-even PN corrections

Einstein 1918

$$\begin{aligned}
A = 1 &+ \frac{1}{c^2}\left[v^2\left(\frac{29}{42} - \frac{29\nu}{14}\right) + \frac{Gm}{r}\left(-\frac{5}{7} + \frac{8}{7}\nu\right)\right] + \frac{1}{c^4}\left[\frac{Gm}{r}v^2\left(\frac{2021}{756} - \frac{5947}{756}\nu - \frac{4883}{756}\nu^2\right) + \frac{G^2m^2}{r^2}\left(-\frac{355}{252} - \frac{953}{126}\nu + \frac{337}{252}\nu^2\right)\right. \\
&\left. + v^4\left(\frac{253}{504} - \frac{1835}{504}\nu + \frac{3545}{504}\nu^2\right) + \frac{Gm}{r}\dot{r}^2\left(-\frac{131}{756} + \frac{907}{756}\nu - \frac{1273}{756}\nu^2\right)\right] + \frac{1}{c^6}\left[v^6\left(\frac{4561}{11088} - \frac{7993}{1584}\nu + \frac{117067}{5544}\nu^2\right.\right.
\end{aligned}$$

Comparing one-loop amplitude to MPM waveform

(Bini-TD-Geralico, WIP)

$$W(t,\theta,\phi) \sim \frac{1}{c^4}\left(G\,(\text{stationary}) + G^2(1+\frac{1}{c^1}+\frac{1}{c^2}+\frac{1}{c^3}+\cdots) + G^3(1+\frac{1}{c^1}+\frac{1}{c^2}+\frac{1}{c^3}+\cdots) + O(G^4)\right)$$

tree-level **one-loop**

Aim: accuracy up to radiation-reaction effects: O(1/c^5) beyond LO quadrupole

$$U_{ij}(\omega) \sim \left(G(1+\frac{1}{c^2}+\frac{1}{c^4}) + G^2(1+\frac{1}{c^2}+\frac{1}{c^3}+\frac{1}{c^4}+\frac{1}{c^5}) + O(G^4)\right) + O(\frac{1}{c^6})$$

LO tail **rad-reac plus similar effects**

Current results (to be confirmed):

the terms linked to time-even PN corrections to multipoles agree
in the test-mass limit m1>>m2, the O(nu) 1/c^3 and 1/c^5 terms agree (inc. tail harmonic #)
in the equal-mass case m1=m2, the O(nu^2/c^5) terms disagree
in addition: we find that supposedly negigible O(hbar) terms are classically important

CaronHuot+'23 have argued that the current HEFT waveform is incomplete because some PV —> Ret.

Other Current Puzzles

high-energy limits?

G^3 energy loss too large

G^3 angular momentum loss too large (Manohar-Ridgway-Shen'22)

Rad-reacted G^4 scattering diverges (Porto..,Damgaard..)

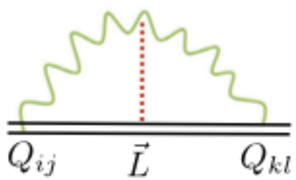
cf ACV motivation: BH formation in HE scattering

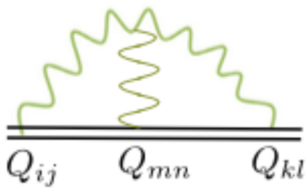
Subtleties in defining/computing angular momentum flux

(Ashtekar et al., Veneziano-Vilkovisky, Riva-Vernizzi,…)

low-energy discrepancy at 5PN between

Foffa-Sturani'19,21,22 Bluemlein et al'21 and Bini-TD-Geralico





$$
\begin{aligned}
S_{QQL} &= C_{QQL} G^2 \int dt I^{(4)}_{is} I^{(3)}_{js} \varepsilon_{ijk} L_k \\
S_{QQQ_1} &= C_{QQQ_1} G^2 \int dt I^{(4)}_{is} I^{(4)}_{js} I_{ij} \, , \\
S_{QQQ_2} &= C_{QQQ_2} G^2 \int dt I^{(3)}_{is} I^{(3)}_{js} I^{(2)}_{ij} \, .
\end{aligned}
$$

TF-constraint on 5PN O(nu^2) EFT radiative terms

$$0 = \frac{2973}{350} - \frac{69}{2} C_{QQL} + \frac{253}{18} C_{QQQ_1} + \frac{85}{9} C_{QQQ_2}$$

not solved by recent in-in results (Foffa-Sturani'22)

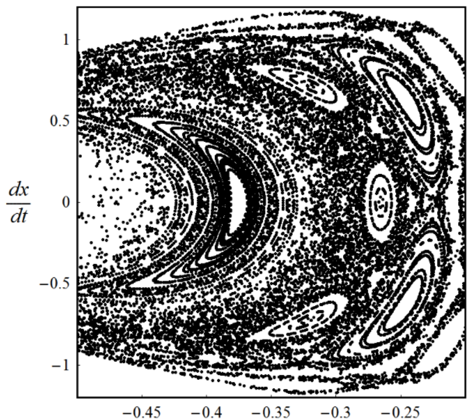
Conclusions

- Analytical approaches to GW signals play a crucial role (in conjunction with Numerical Relativity simulations) for the detection, interpretation and parameter estimation of coalescing binary systems (BBH and BNS). It is important to further improve our analytical knowledge for future GW detectors: second generation ground-based detectors, space detectors, second generation ground-based detectors.

- Quantum (and classical) scattering approaches have given new results of great conceptual interest, and also of potential interests for GW detection. The fruitful dialogue between QFT, EFT, PN, PM, EOB, Tutti-Frutti methods must be vigorously pursued. Discrepancies must be resolved to complete the determination of the **5PN** dynamics (of direct utility for LIGO-Virgo). Radiative effects are still puzzling. Quantum PM waveform computations are promising (though their G-accuracy wont compete with MPM).

Henri Poincaré

«Il n'y a pas de problèmes résolus, il y a seulement des problèmes plus ou moins résolus »



«There are no (definitely) solved problems, there are only more or less solved problems »